



17th August,
KIAS

Walls, Bubbles & Doom

the Cosmology of HET.

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in collaboration with Rodrigo Alonso, Juan Carlos Cárdenas & Rachel Houtz

Aims

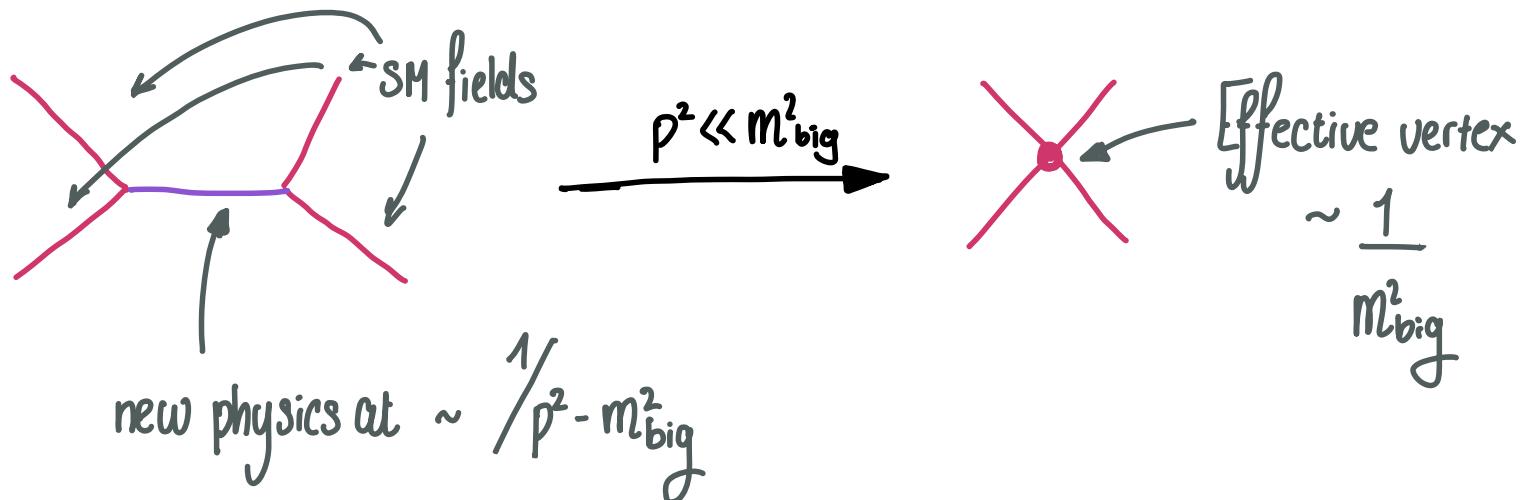
- ① What is HEFT & why should I care about it?
- ② What does the LHC not tell us?
- ③ What can cosmology tell us that the LHC can't?
- ④ Some pheno! (ie what could tEFT look like in gravitational wave format?)

1 ||| What is EFT?

Effective field Theory.

Take the SM to be a low-energy, effective theory: all new physics is at

$$E \gg \Lambda \sim \text{TeV}$$



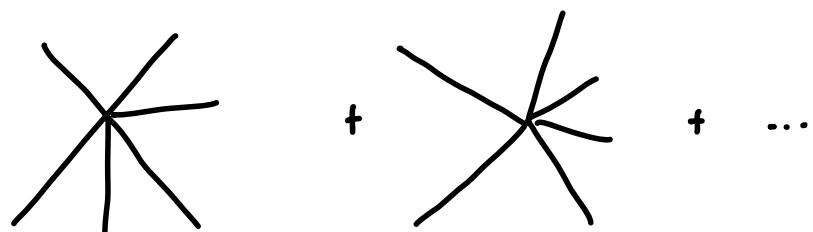
Think 4-fermi theory!!

Since we know 'integrating out' UV physics makes these 'effective', higher-point vertices, we can write down all the ways this can happen...

$$\mathcal{L}_{\text{EFT}} = \mathcal{L}_{\text{low-energy}} + \sum_{d=4} \frac{c_i}{\Lambda^{d-4}} Q_i^{(d)}$$

↑ Wilson coefficients

Operators made of SM fields.
(obeying SM symmetry group)



SMEFT

Linearly realised ew symmetry

$$SU(3)_c \times SU(2)_L \times U(1)_Y$$

$$H = \frac{1}{\sqrt{2}} \begin{pmatrix} \phi_1 + i\phi_2 \\ \phi_4 + i\phi_3 \end{pmatrix} \sim \frac{(h+V)}{\sqrt{2}} \begin{pmatrix} \text{radial} \\ \text{'angular'} \end{pmatrix}$$

$U(4)$

V_{SMEFT} is analytic function of $H^\dagger H$

HEFT

Non-linearly realised ew symmetry

$$SU(3)_c \times U(1)_{\text{em}}$$

most general EFT parameterisation of low energy physics involving SM DoF.

$$h \quad \text{and} \quad \vec{\Phi} = \begin{pmatrix} \varphi_1 \\ \varphi_2 \\ \varphi_3 \\ \sqrt{V - \vec{\varphi} \cdot \vec{\varphi}} \end{pmatrix}$$

\downarrow

$h \rightarrow h$ and $\vec{\Phi} \rightarrow 0 \vec{\Phi}$

$$V_{\text{HEFT}} = V(h)$$

$$\int_{\text{HEFT}} = \frac{1}{2} (\partial h)^2 + \frac{v^2 f(h)^2}{2} (\partial \vec{\psi})^2$$

flare function $f(h) = 1 + ah + bh^2 + \dots$

$$+ \frac{v^2}{8} f(h)^2 \left[2g^2 W_\mu^+ W_\mu^- + (g^2 + g'^2) Z_\mu Z^\mu \right]$$

$$- \frac{m_h^2}{2} h^2 - \frac{m_h \sqrt{\lambda}}{2} \gamma_4(1-\epsilon) h^3 - \frac{\lambda}{8} \gamma_4^2 h^4 + \dots$$

$\underbrace{\qquad\qquad\qquad}_{V(h)}$

+ higher derivatives which we will neglect

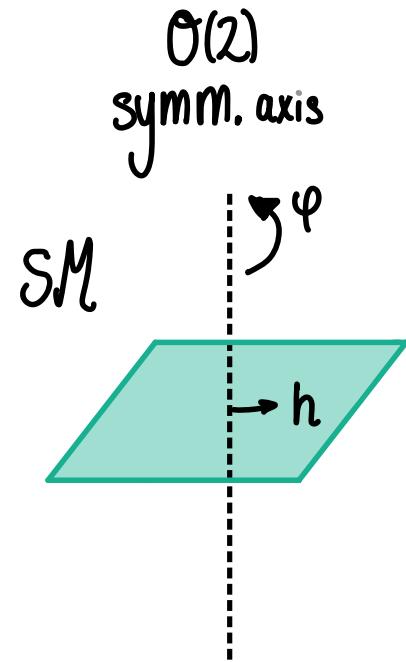
NB for the SM, $f(h) = 1 + \frac{h}{v}$ and $h \in [0, \infty]$.

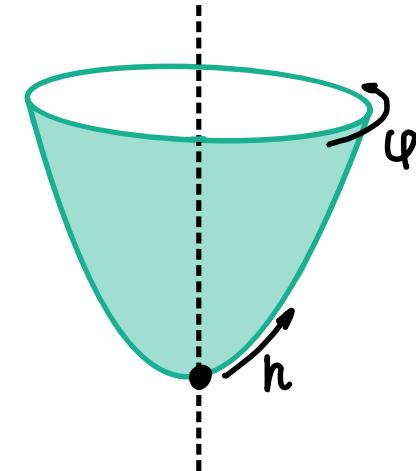
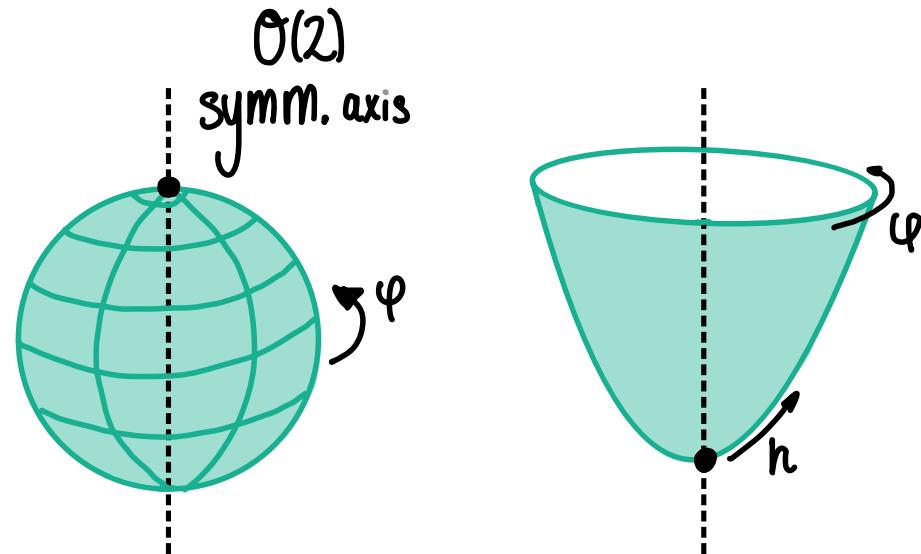
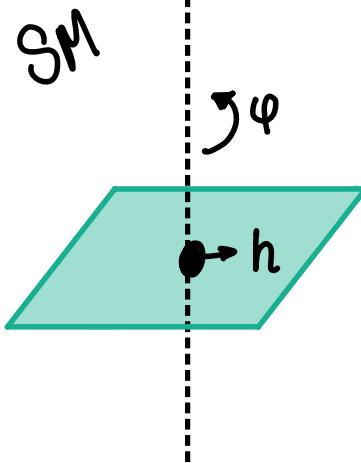
ttHET theories as Manifolds: [arXiv: 1605.03602]

1: Symmetric axis around which $\vec{\Psi} \rightarrow 0\vec{\Psi}$

↳ Goldstone bosons act as angular modes

2: Higgs boson in the 'radial' direction





Examples of manifolds you can draw!

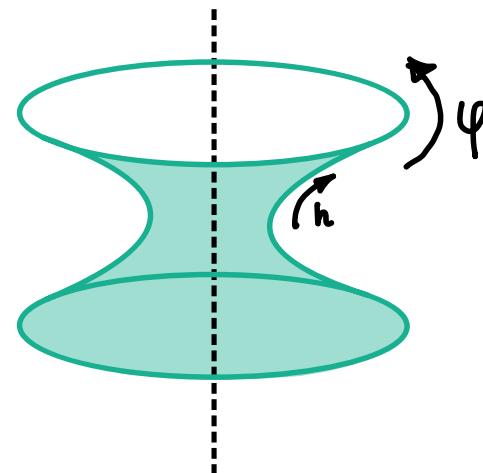
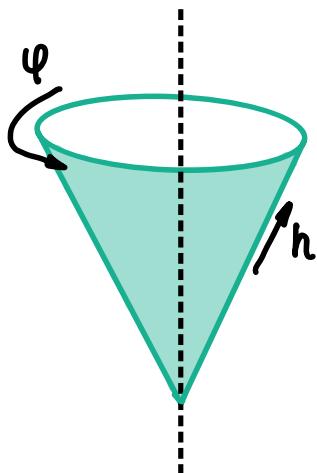
See eg:

[arXiv: 2109.1329]

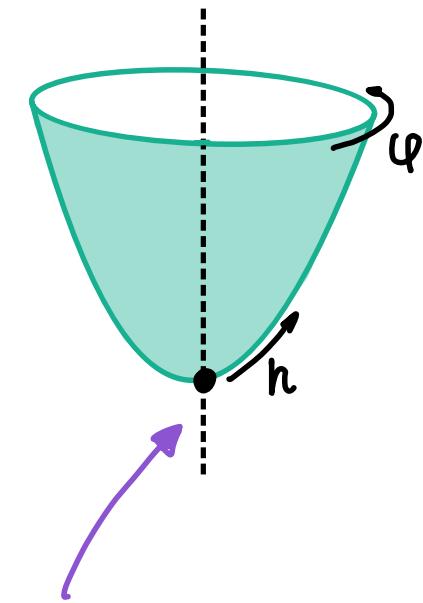
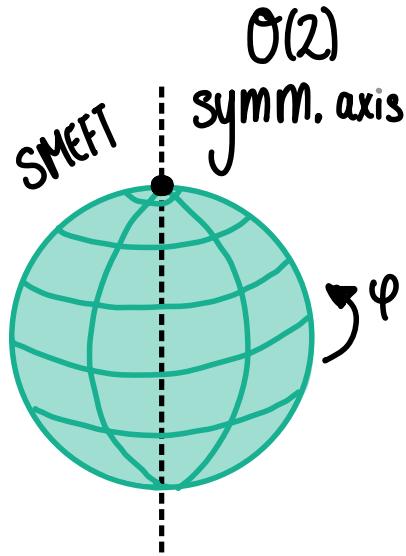
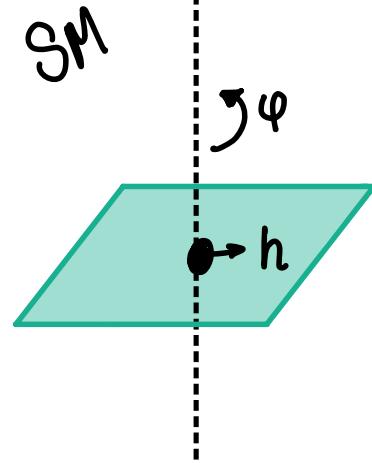
R. Alonso & MW]

[arXiv: 2008.08597; T. Cohen,

N. Craig, X. Lu, D. Sutherland]

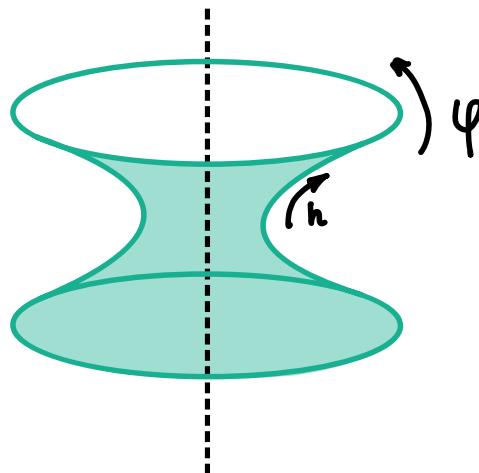
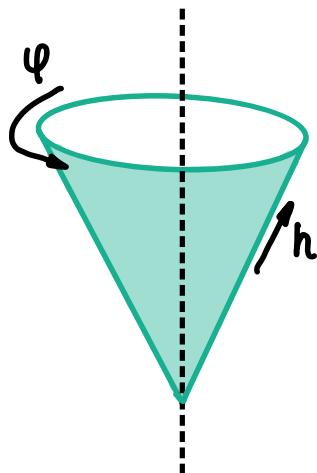


SM

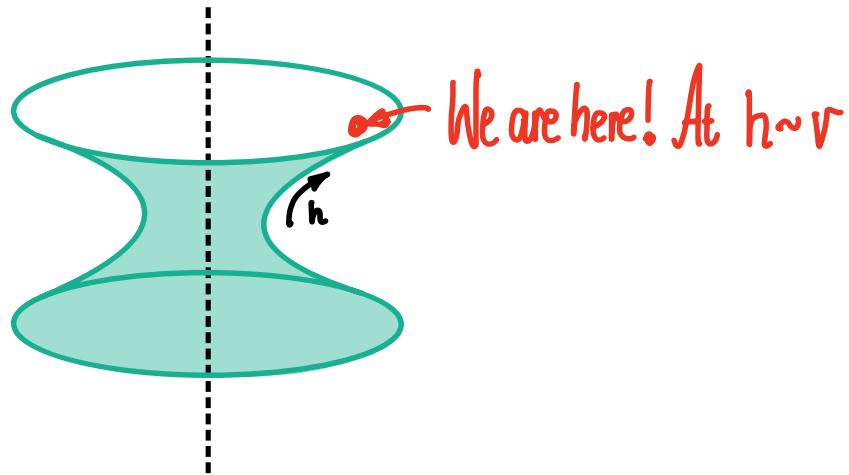


Does your manifold have an $O(4)$ invariant fixed point or not?

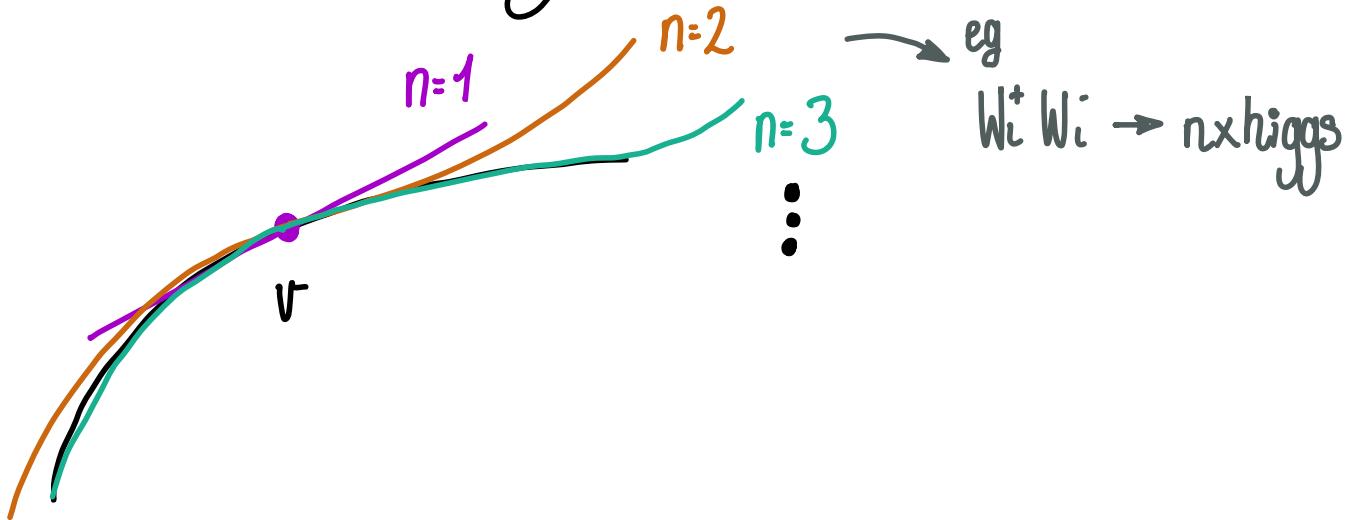
HEFT/SMEFT
(Quotient)



But how can we look for a hole or singularity?
THIS TALK [S. Kanemura, R. Nagai, M. Tanaka,
arXiv:2202.12774]



Scattering experiments look locally around the vacuum.



What do we do?

a) Panic and/or give up?

Too stubborn

b) Hope that an answer miraculously descends to us from the sky.

Sounds about right!

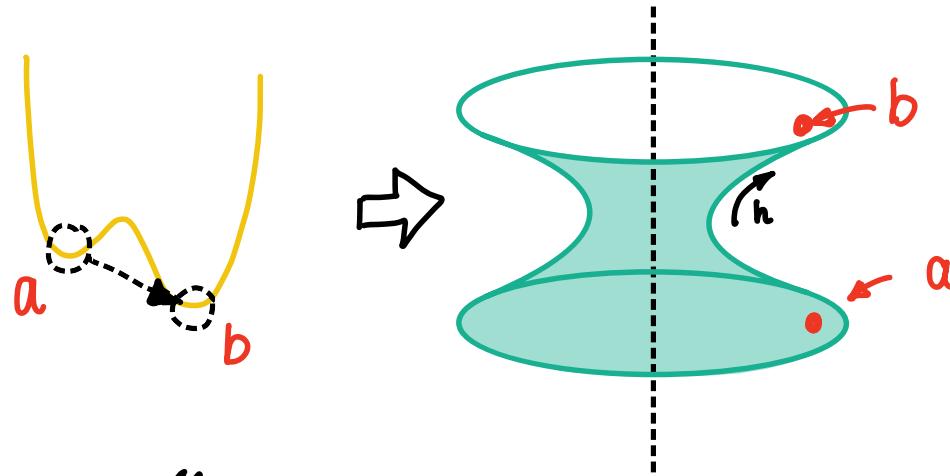
Look at Cosmology for a global approach

1: Sphaleron Tunnelling

2: Vacuum Evolution

↳ Walls, Bubbles & Doom.

We're interested in whether the universe may have tunneled between vacua as the universe cooled...



We need to build ourselves an effective potential

↳ Tree level (quartic) potential.

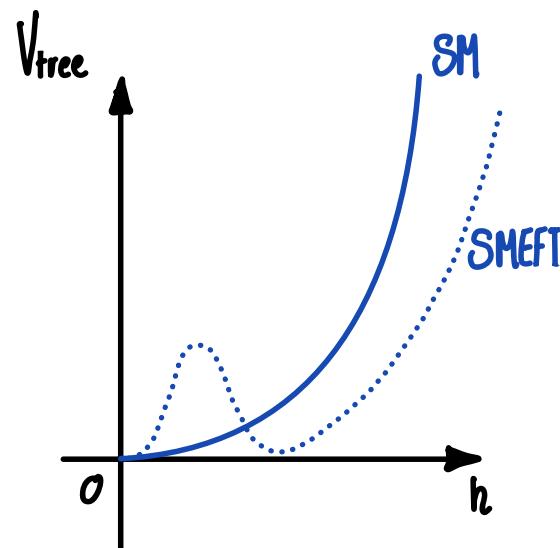
$$V_{\text{eff}} = V_{\text{tree}} + V_{\text{thermal}} + V_{\text{loops}}$$

↳ Temperature in the early universe modifies couplings.

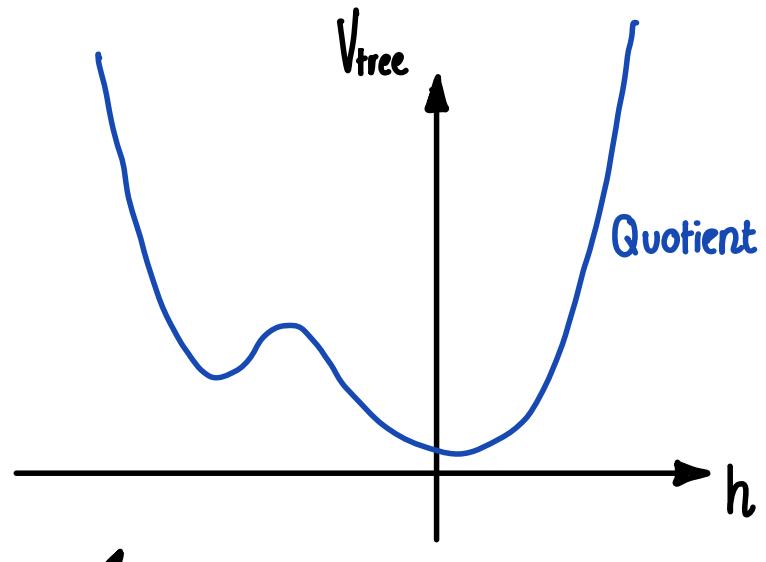
↳ Loop corrections modifies couplings

~~EFT~~ Potential at Tree Level

SM and SMEFT



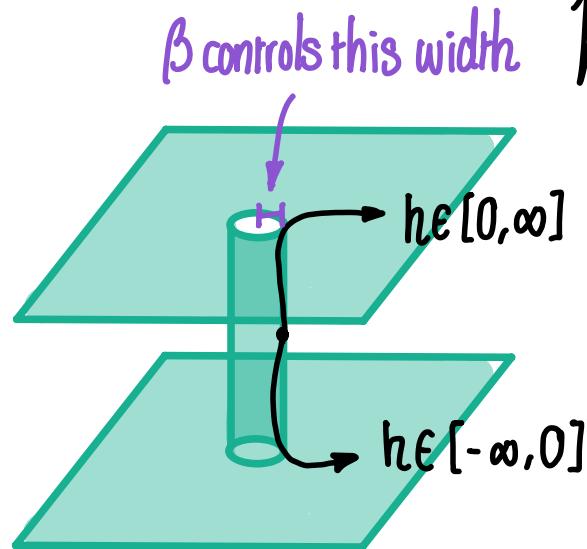
Quotient Theory



Quotient theories can allow for $h \in [-\infty, \infty]$, revealing possible new minima

Quotient Effective Potential

We would like to open up the domain of the Higgs h to reveal new minima...



We've done this by gluing together two SM manifolds..

$$h \in [-\infty, \infty]$$

$$F(h) = \sqrt{\sin^2 \beta + \cos^2 \beta (1 + \gamma a h/v)^2}$$

NB// Keep in mind that we've restricted quotient space!

Higgs - Dependent Masses

$$\tilde{\mathcal{D}} \sim p^2 - m^2(h)$$

- effective potential: $V_{\text{eff}}(\phi_c) = V_0(\phi_c) + i \int \frac{d^4 h}{(2\pi)^4} \text{Tr} \log i \tilde{\mathcal{D}}^{-1}(\phi_c, h) + \dots$

acts over all Lorentz & gauge indices.

$$\int_{\text{HEFT}} = \frac{v^2 F(h)^2}{8} \left[2g^2 W_\mu^+ W_\mu^- + (g^2 + g'^2) Z_\mu Z^\mu \right] + V(h) + \dots$$

$m_{W/Z}(h) = m_{W/Z} F^2(h)$

$m_h^2 = V''(h)$

& for simplicity also $m_t(h) = m_t f^2(h)$

If only it were that simple!

1710.06848 : R.Alonso, K.Kanshin, S.Saa
1511.00724 : R.Alonso, E.Jenkins,
A.Manohar

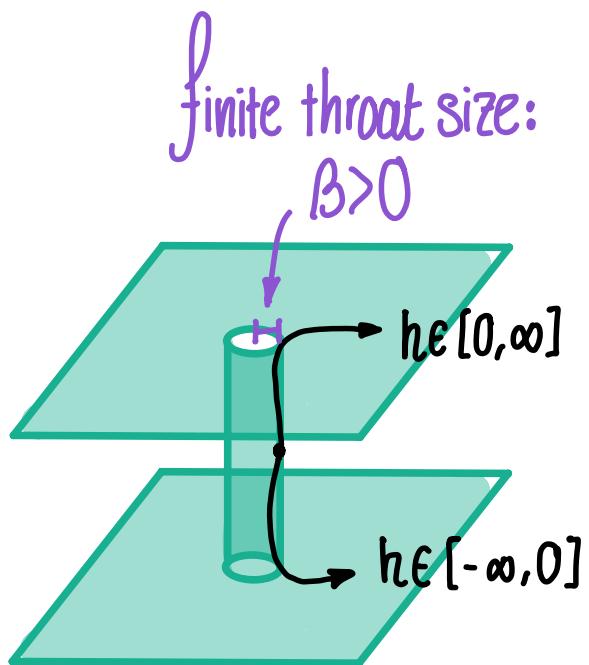
- Want to compute in *on-shell gauge*
- W_L^\pm and Z_L get a contribution from the Goldstones.
 $\rightarrow m_h$ behaves nicely!

Higgs-dependent masses:

$$m_h^2(h) = V_{\text{tree}}''(h)$$

$$m_t^2(h) = m_t^2(0) F^2(h)$$

$$m_{\Sigma_W, Z_3}(h) = (\log F)' V + m_{\Sigma_W, Z_3}(0) \quad m_{\Sigma_W, Z_3}^2(h) = m_{\Sigma_W, Z_3}^2(0) F^2(h)$$



Bear in mind for $\beta > 0$,

$$F(h) = \sqrt{\sin^2 \beta + \cos^2 \beta (1 + \gamma_0 h/v)^2}$$

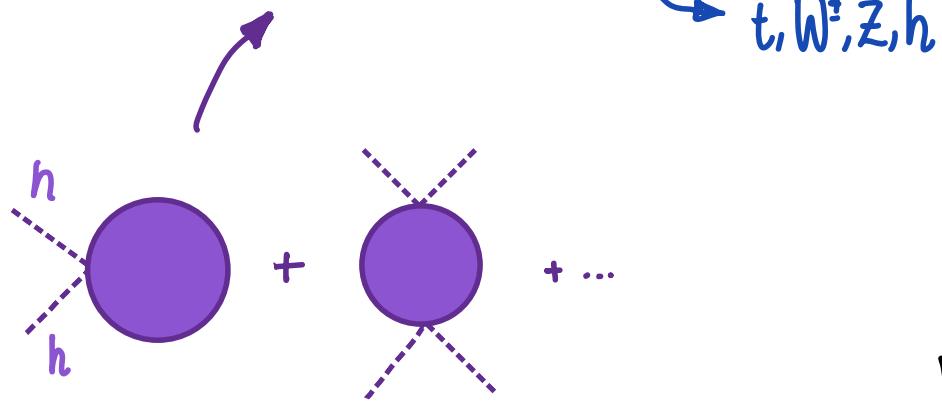


$$\min(F(h)) = \sin \beta$$

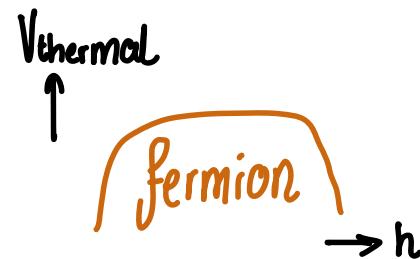
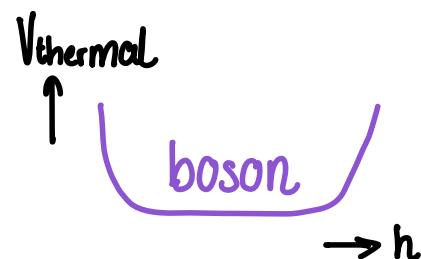
$$\min(M_{w/z}(h)) = M_{w/z} \sin^2 \beta$$

Loop Corrections in on-shell scheme...

$$V_{\text{1-loop}} = \frac{1}{64\pi^2} \sum_i \left\{ m_i^4(h) \left(\log \frac{m_i^2(h)}{m_i^2(0)} - \frac{3}{2} \right) + 2m_i^2(h)m_i^2(0) \right\}$$



$$V_{\text{1-loop}} \sim$$



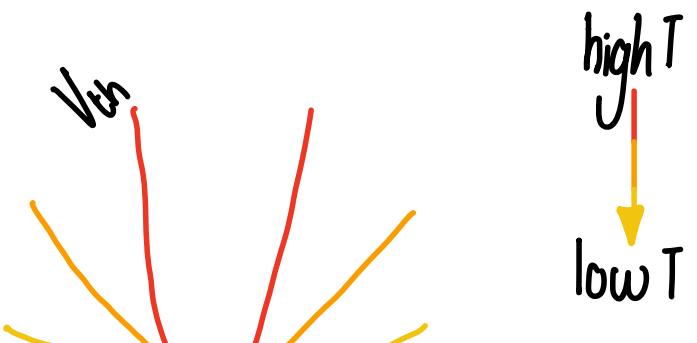
Thermal Corrections

arXiv: 9901312: M. Quiros

$$V_{\text{thermal}} = \frac{T^4}{2\pi^2} \sum_i n_i J_{B/F} \left(\frac{m_i^2(h)}{T^2} \right) \sim$$

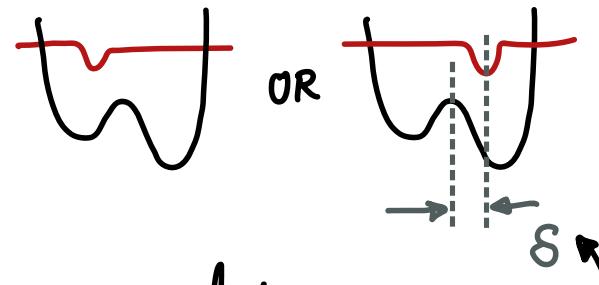
W^\pm, Z, h, t

$$J_{B/F} = \pm \int_0^\infty dx x^2 \log [1 \mp e^{-\sqrt{x^2 + m^2/T^2}}]$$



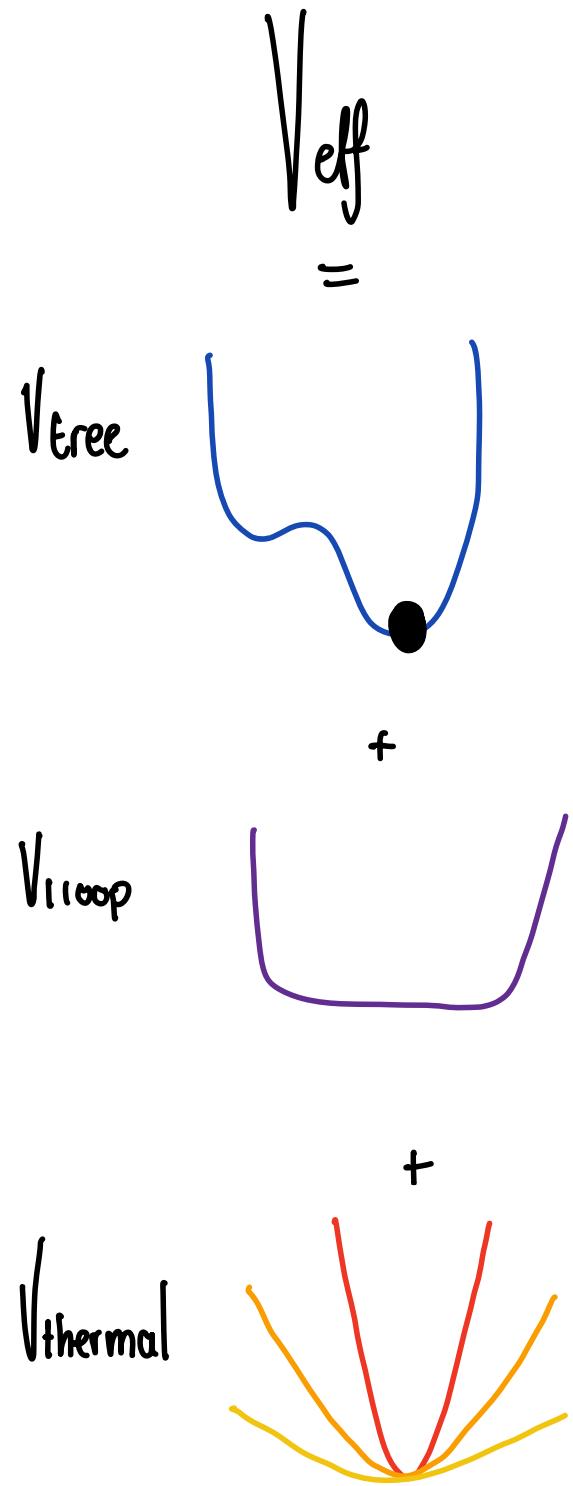
Changing γ_a changes where the thermal corrections are wrt V_{tree} :

$$f(h) = \sqrt{\sin^2 \beta + \cos^2 \beta \left(1 + \frac{\gamma_a h}{\sqrt{v}}\right)^2}$$



the sign - delta tells us which side of V_{tree} thermal corrections go at tree level.

$$\delta = \begin{cases} \gamma_4^{-1} \gamma_\epsilon - \gamma_a^{-1} & \epsilon \leq \sqrt{9/8 - 1} \\ \gamma_4^{-1} \sqrt{2} - \gamma_a^{-1} & \epsilon > \sqrt{9/8 - 1} \end{cases}$$



Validity of V_{thermal}

Linde's Infrared Problem

A. Linde, 1980

2009.10080: D. Croon, G. Gould, et al

At high temperatures, IR bosonic modes are highly occupied.

$$\epsilon_{IR} = \frac{g^2 T}{\pi M_{ew}(h)}$$

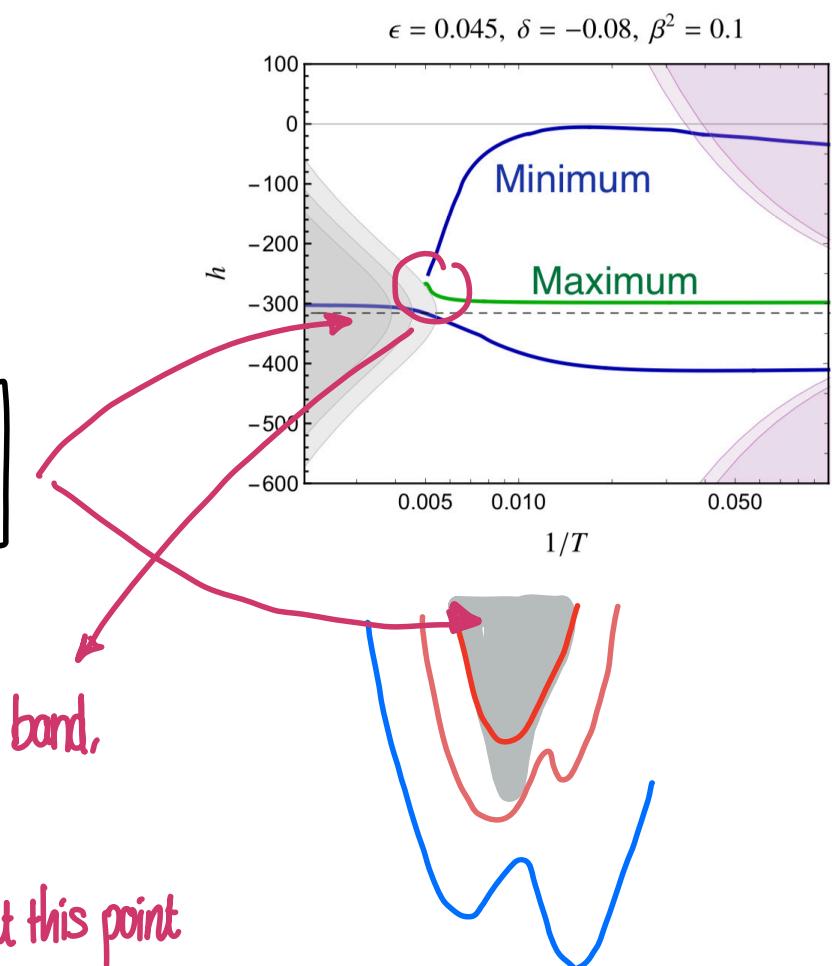
↑
effective expansion
parameter

Define IR-unsafe band

$$\left(\frac{h + V/\delta_a}{T} \right)^2 > \frac{1}{\gamma_a} \left[\left(\frac{2g}{C\pi \cos \beta} \right)^2 - \tan^2 \beta \frac{V_h^2}{T} \right]$$

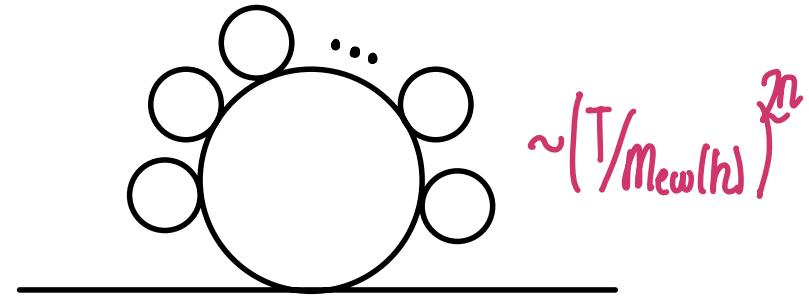
If these appear OUTSIDE band,
we're OK.

If not, we say nothing about this point

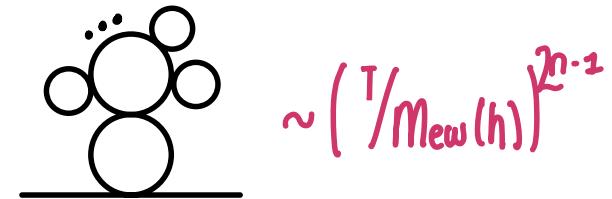


What advantages does this have over daisy corrections?

ie we resum the most IR divergent diagrams



1// Only elevates the problem
Super daisies diverge next



2// Maintain a consistent loop expansion

[D. Curtin, P. Meade, H. Ramani
arXiv: 1612.00466]

3// It turns out that in our quotient theories (since $M_w(h) / M_z(h)$ cannot get too small), most interesting phenomenology is outside the band.

And dimensional reduction? 2009.10080: D. Croon, G. Gould, et al

↳ Takes advantage of a hierarchy of scales to dimensionally reduce to a 3D effective theory.

If's much simpler.

- UV cutoff.

Because we're working with an EFT:

- $|h| < 4\pi V_{ke} \sim 4\pi v$
- $T < 4\pi V_{ke} \sim 4\pi v$

\rightarrow T. Cohen, N. Craig, X. Lu, D. Sutherland
[arXiv: 2108.03240]

- Boundedness from below before $h = 4\pi v \sim 3 \text{ TeV}$

$V_{1\text{loop}, \text{top}} \sim$

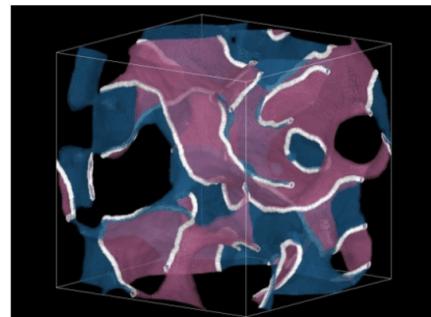
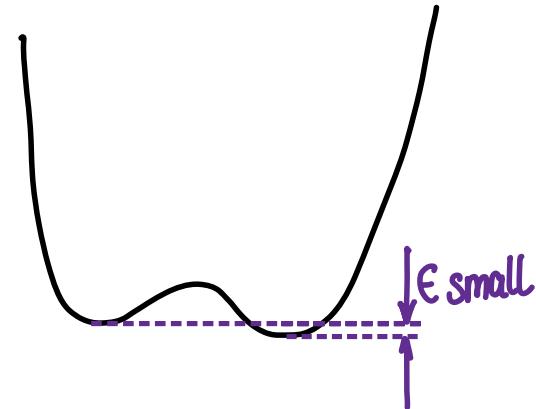


Minima Evolution.

Domain Walls

[G.Gelmini, M.Gleiser, E.Kolb 1989]

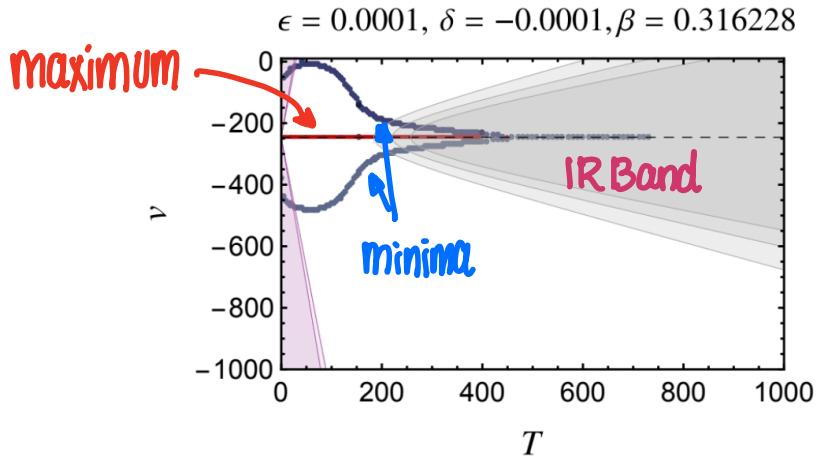
- If $V(h)$ is nearly symmetric around $(h + V/f_0) \rightarrow - (h + V/f_0)$, domain walls form.
- A small bias ($|\epsilon| > 0$) in the potential will let the walls evaporate, leaving behind a GW signal.



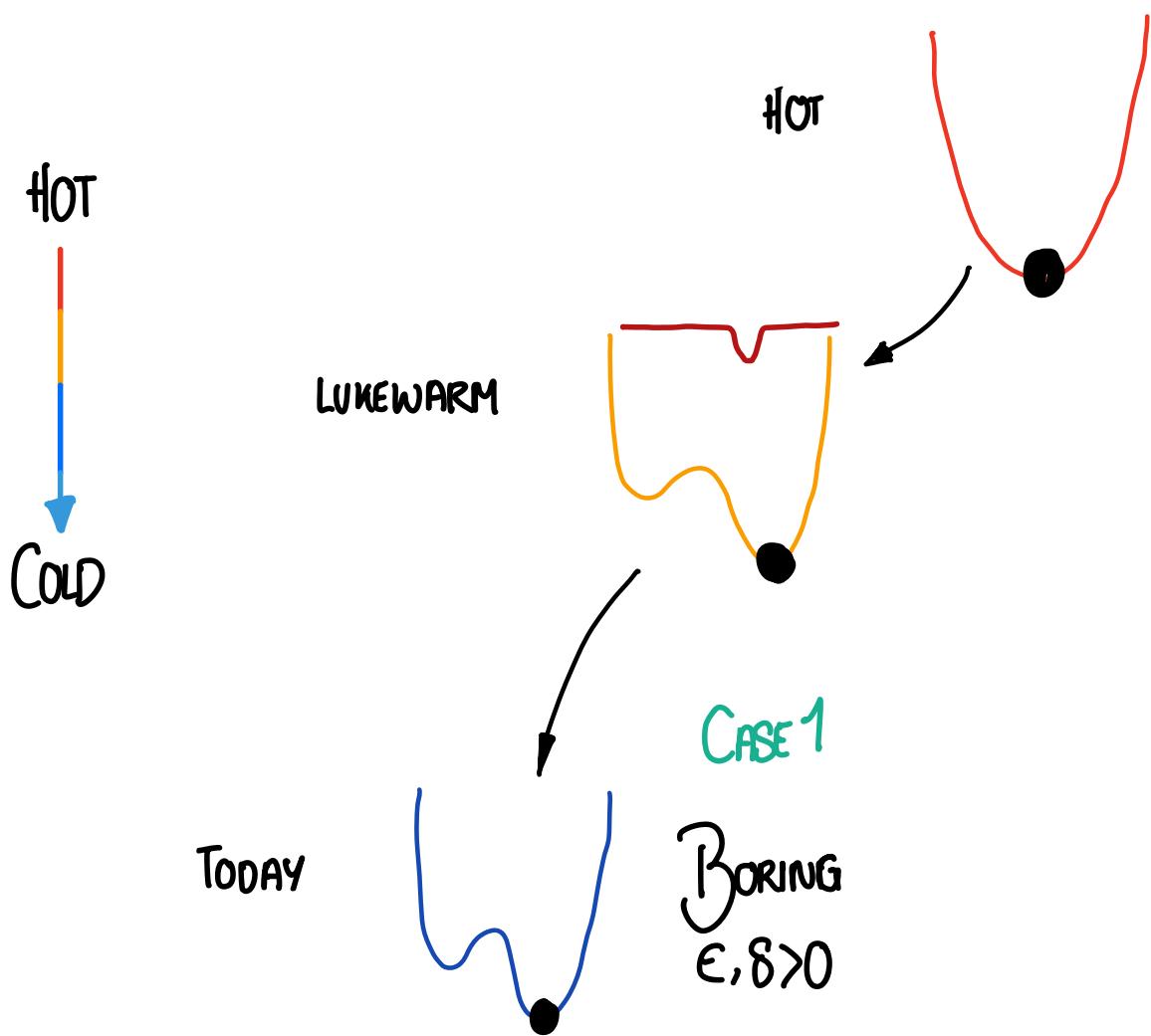
2009.01903 : G. Gelmini, S. Pascoli, E. Vitagliano,
Y-L. Zhou

Image lifted from Hiramatsu, Kawasaki,
Saikawa, Sekiguchi, arXiv:1207.3166

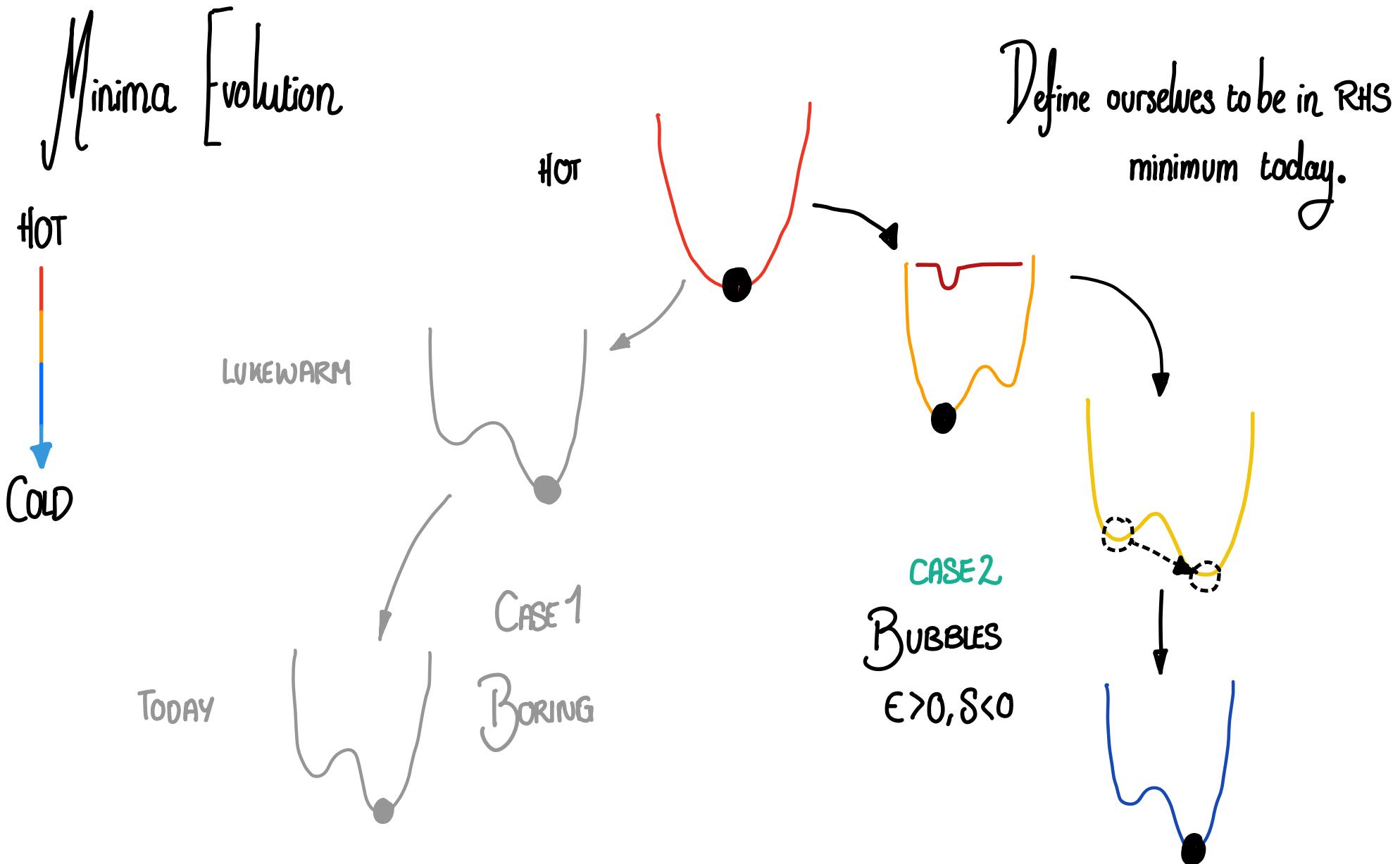
- Note that walls do not have the same IR criterion.
- Due to approximate symmetry of $V(h)$.



Bubbles & Doom



Define ourselves to be in RHS
minimum today.



Minima Evolution

HOT

COLD

LUKEWARM

TODAY

FUTURE?

HOT

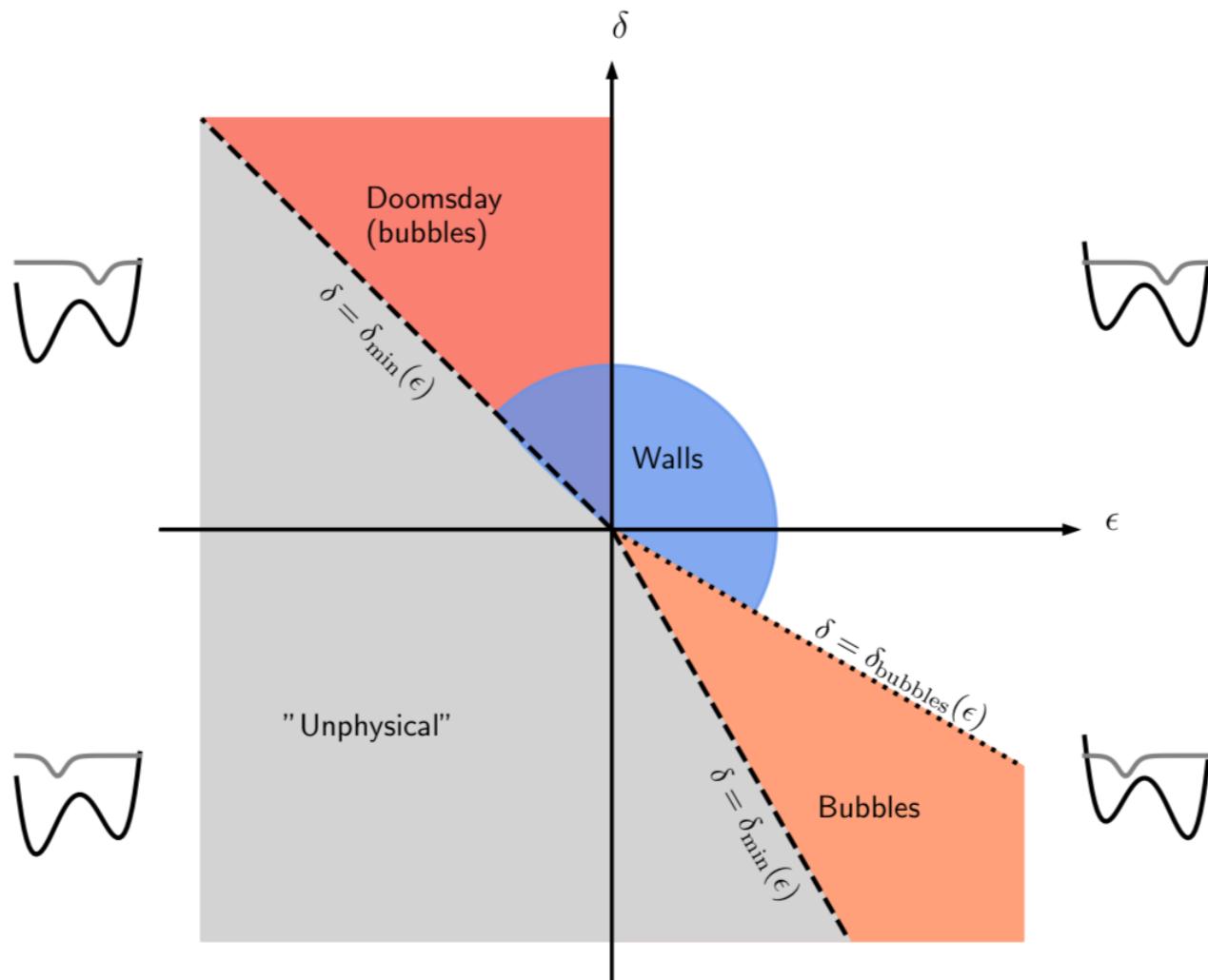
CASE 3

Doom (Non SM effect)

$\epsilon < 0, S > 0$

Define ourselves to be in RHS minimum today.

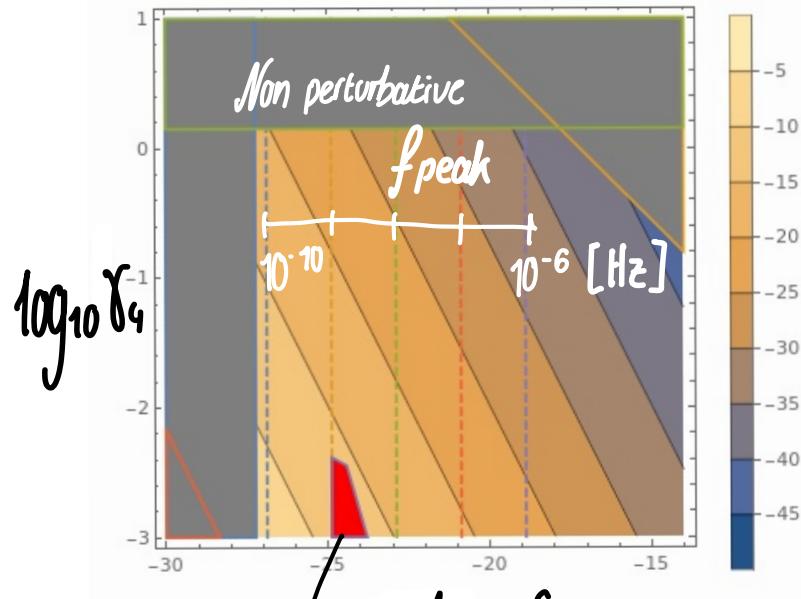
Schematic Summary...



Phenomenology

1 walls

[Preliminary]

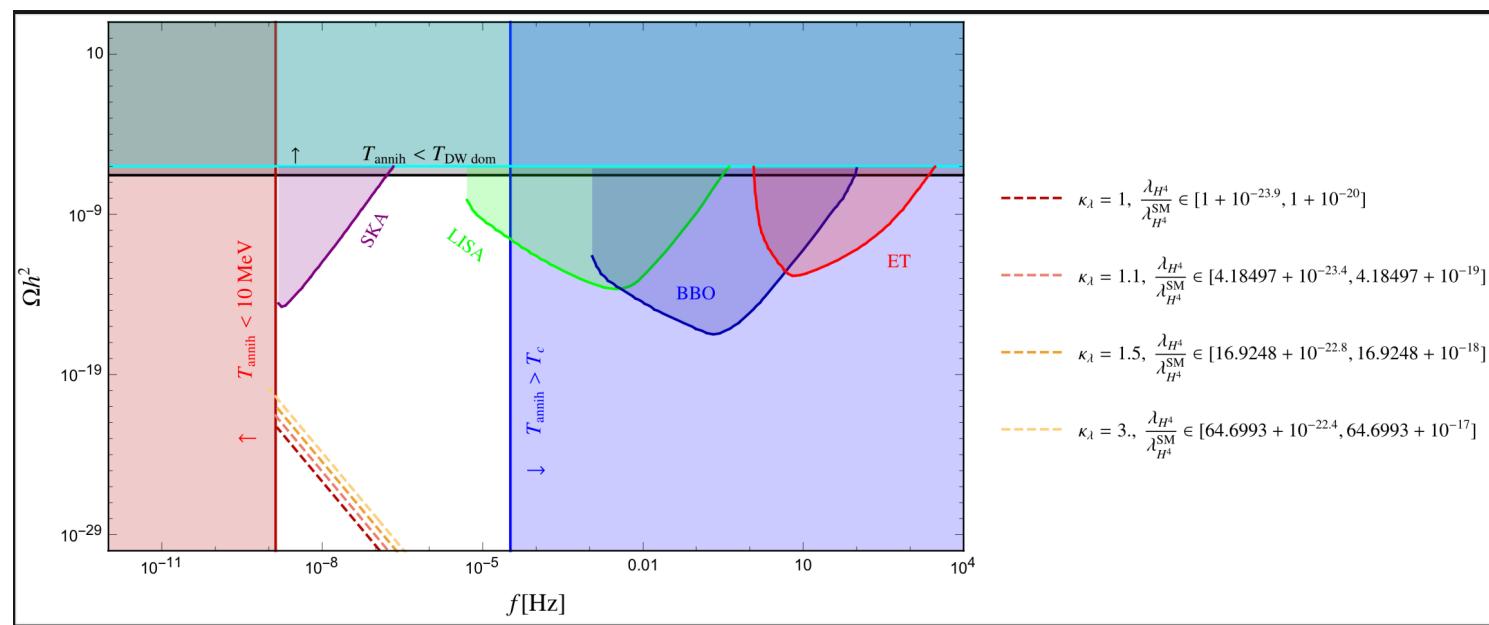


Reduced hubble Not Higgs

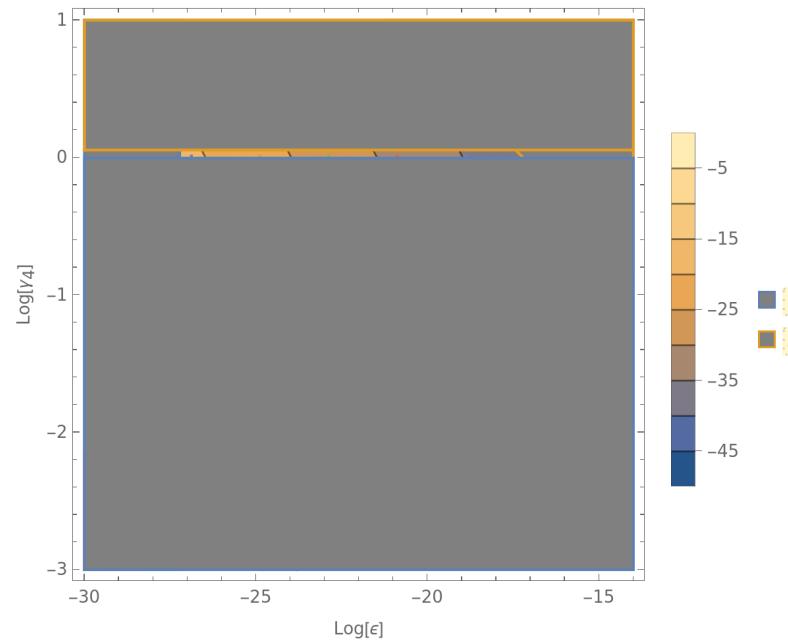
$\log_{10} h^2 \rho_{GW}$

$\log_{10} \epsilon$

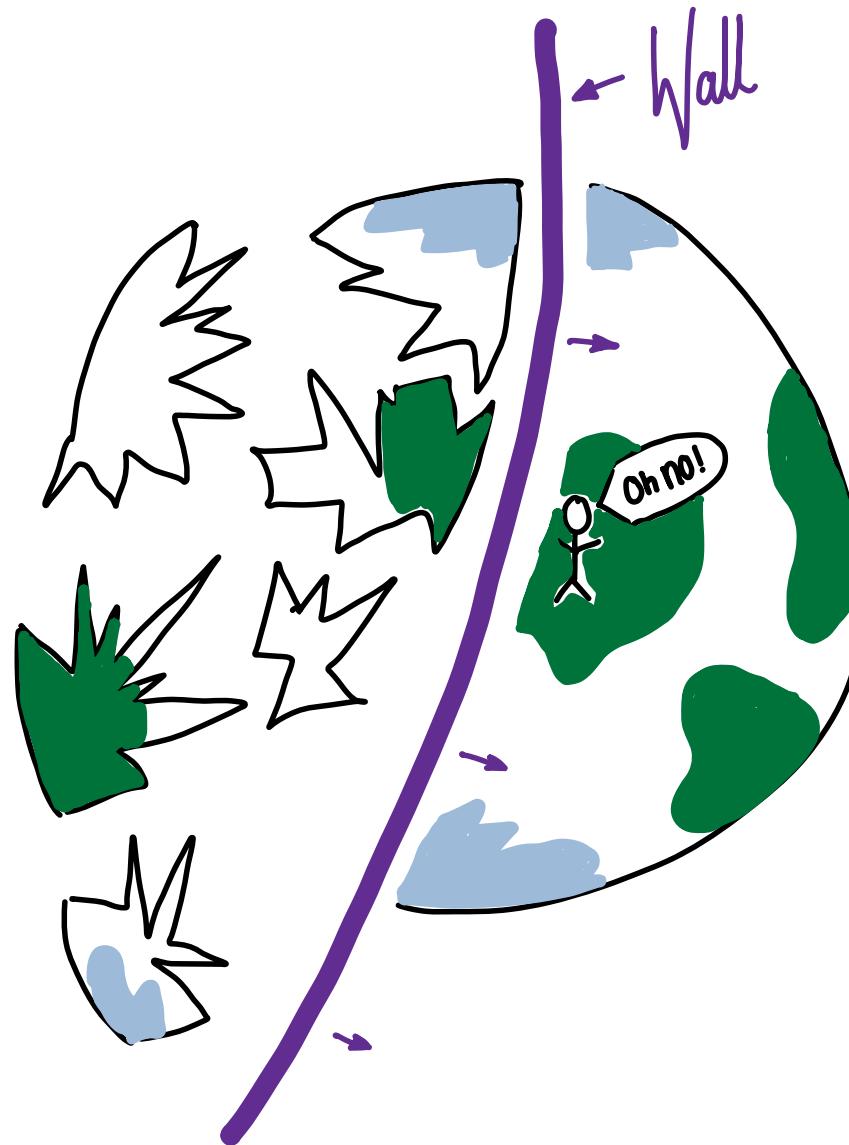
Visible at SKA



Adding in LfC bounds...

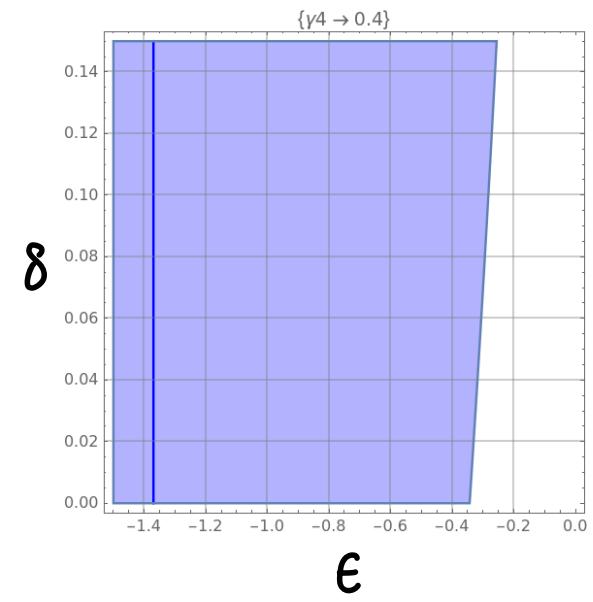
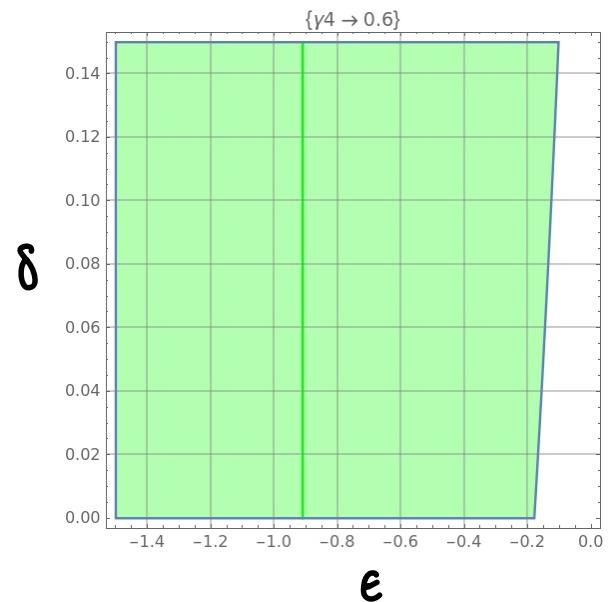
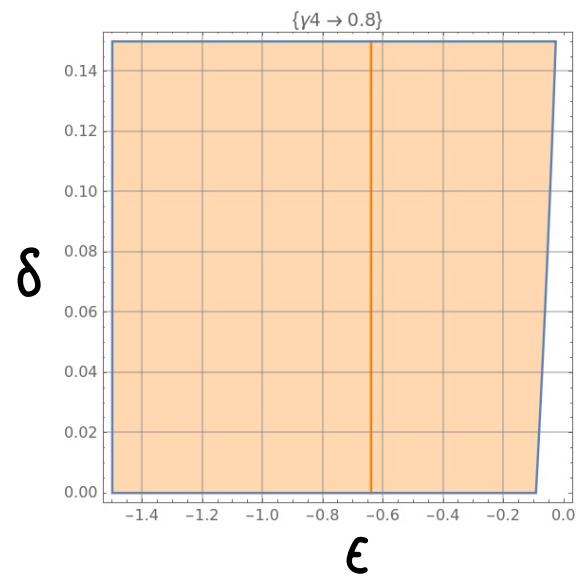
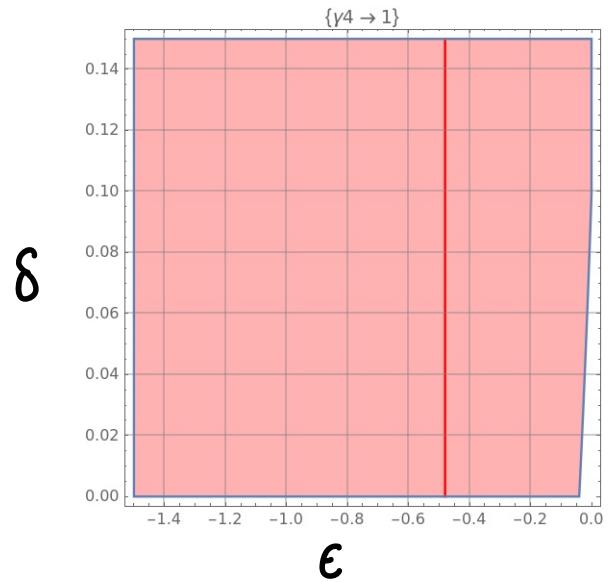


Doom

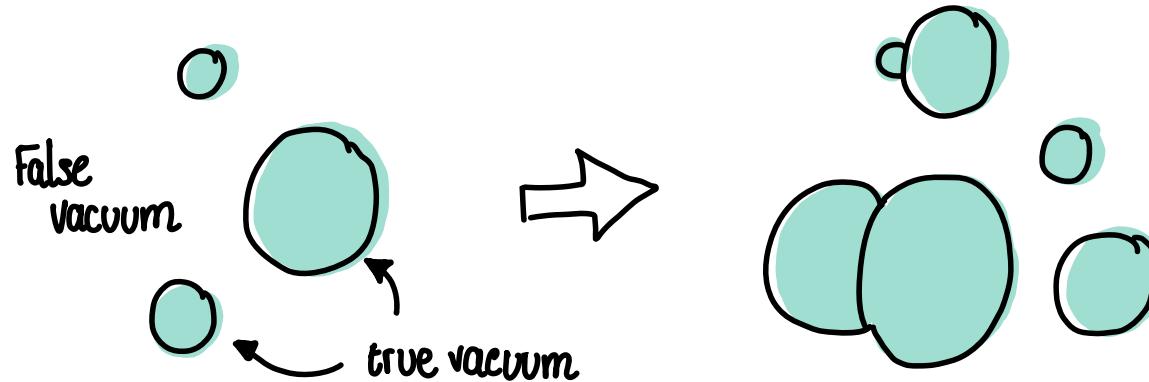
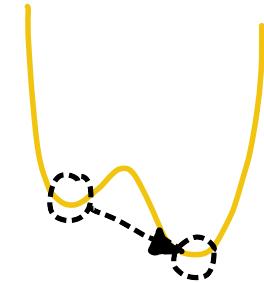


Luckily This Isn't Likely!!

- Shaded is excluded by perturbativity
- Solid line means $P(\text{decay today}) = 1.$



Bubble nucleation in the Early Universe.



3 gravitational wave production mechanisms:

1/1 Bubble collisions

2/1 Sound waves from motion of plasma w/ wall.

3/w Turbulence from shocks in the plasma

T Transition Parameters

- $S_3 = \int d^3x \left[\frac{1}{2} (\partial_\mu h)^2 + V(h) \right]$
- Assuming spherically symmetric bubbles, eom is $\frac{d^2h}{dp^2} + \frac{2}{\rho} \frac{dh}{dp} = \nabla V(h)$
- Boundary conditions are false & true minima & non-singularity at $\rho=0$.

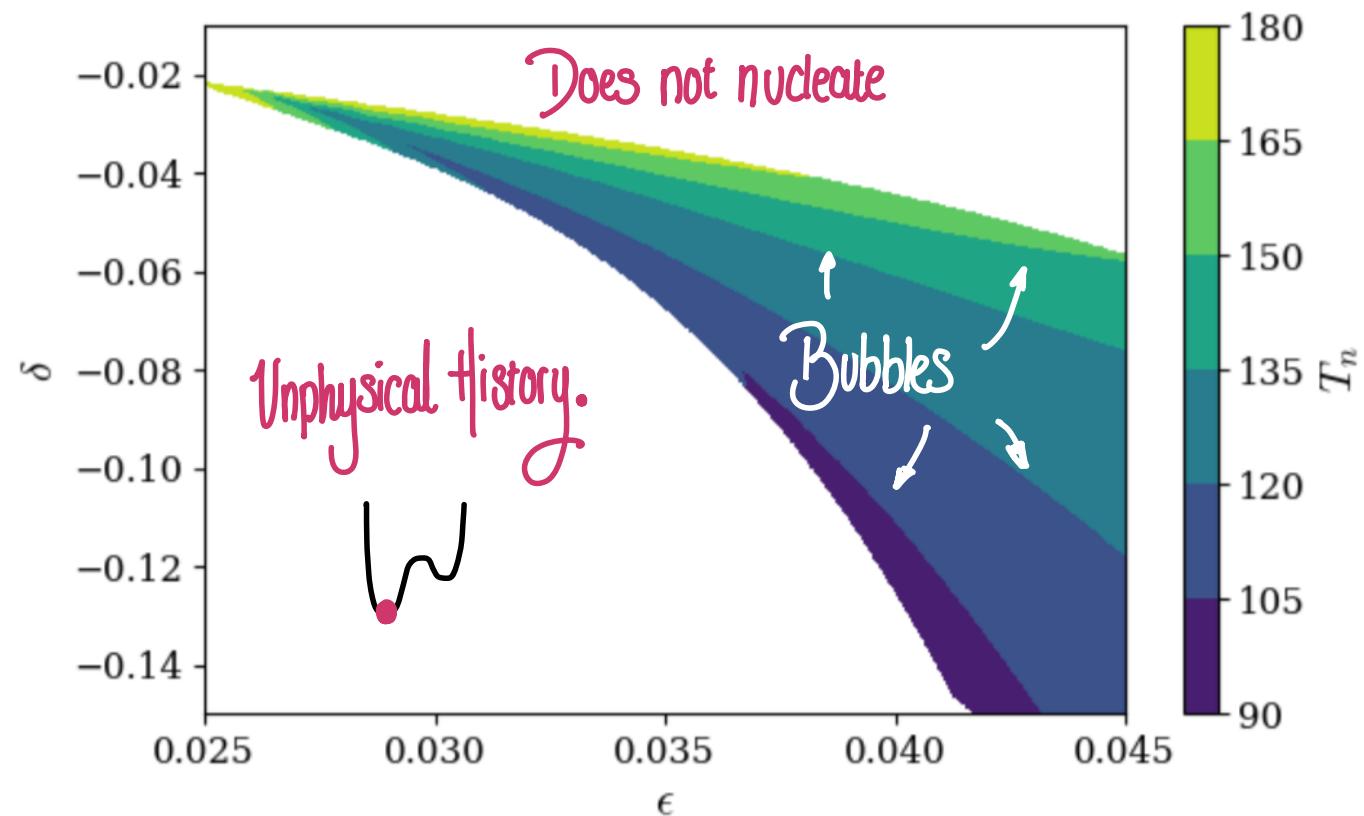
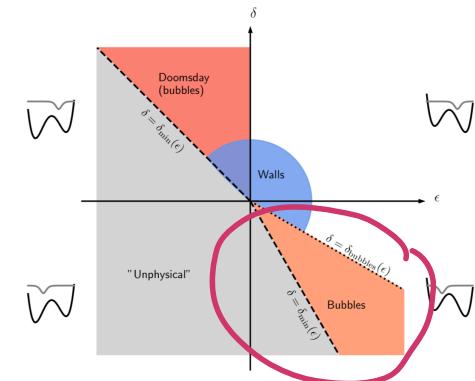
Calculated with:

- CosmoTransitions [C. Wainwright
arXiv: 1109.4189]
- AnyBubble [A. Masoumi et.al.
arXiv: 1610.06594]

Nucleation occurs when $\frac{S_3}{T_W} \sim 1/40$.

1 bubble per Hubble Volume

Bottom right quadrant



GW Spectra for Bubbles.

The gravitational wave signal can be estimated using thermal parameters calculated from the tunelling action.

Latent heat

$$\alpha = \frac{1}{\rho_{\text{ad}}} \left(\Delta V - T \frac{d(\Delta V)}{dT} \right) \Big|_{T=T_N}$$

Duration

$$\frac{\beta}{H_K} = T \frac{d(S_3/T)}{dT} \Big|_{T=T_N}$$

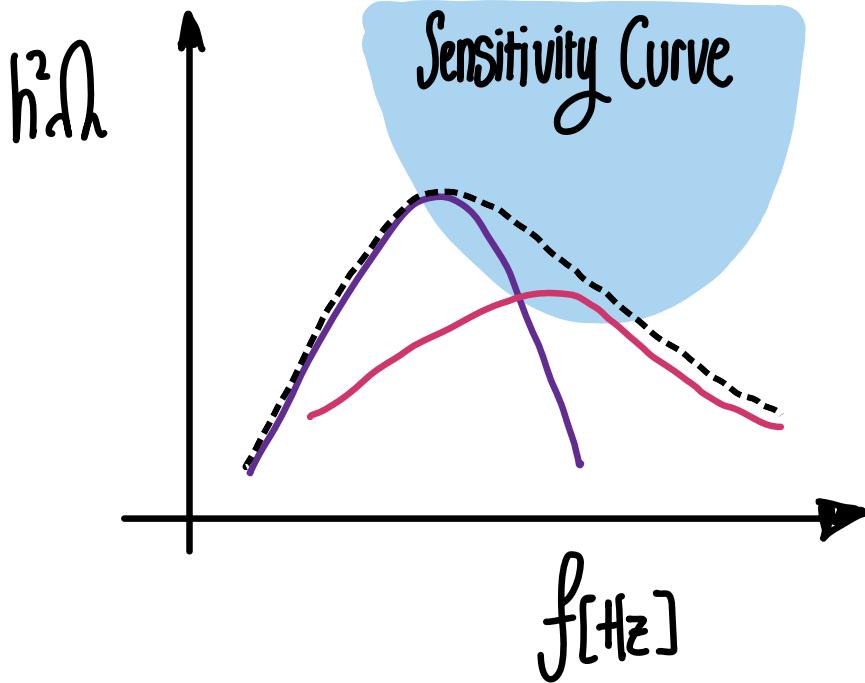
[M. Lewicki, M. Merchand, M. Zych
arXiv: 2111.02393]

Wall Velocity

$$v_w = \begin{cases} \sqrt{\Delta V / \alpha \rho_r} & \sqrt{\Delta V / \alpha \rho_r} < v_j(\alpha) \\ 1 & \sqrt{\Delta V / \alpha \rho_r} \geq v_j(\alpha) \end{cases}$$

What Do the Spectra Look Like?

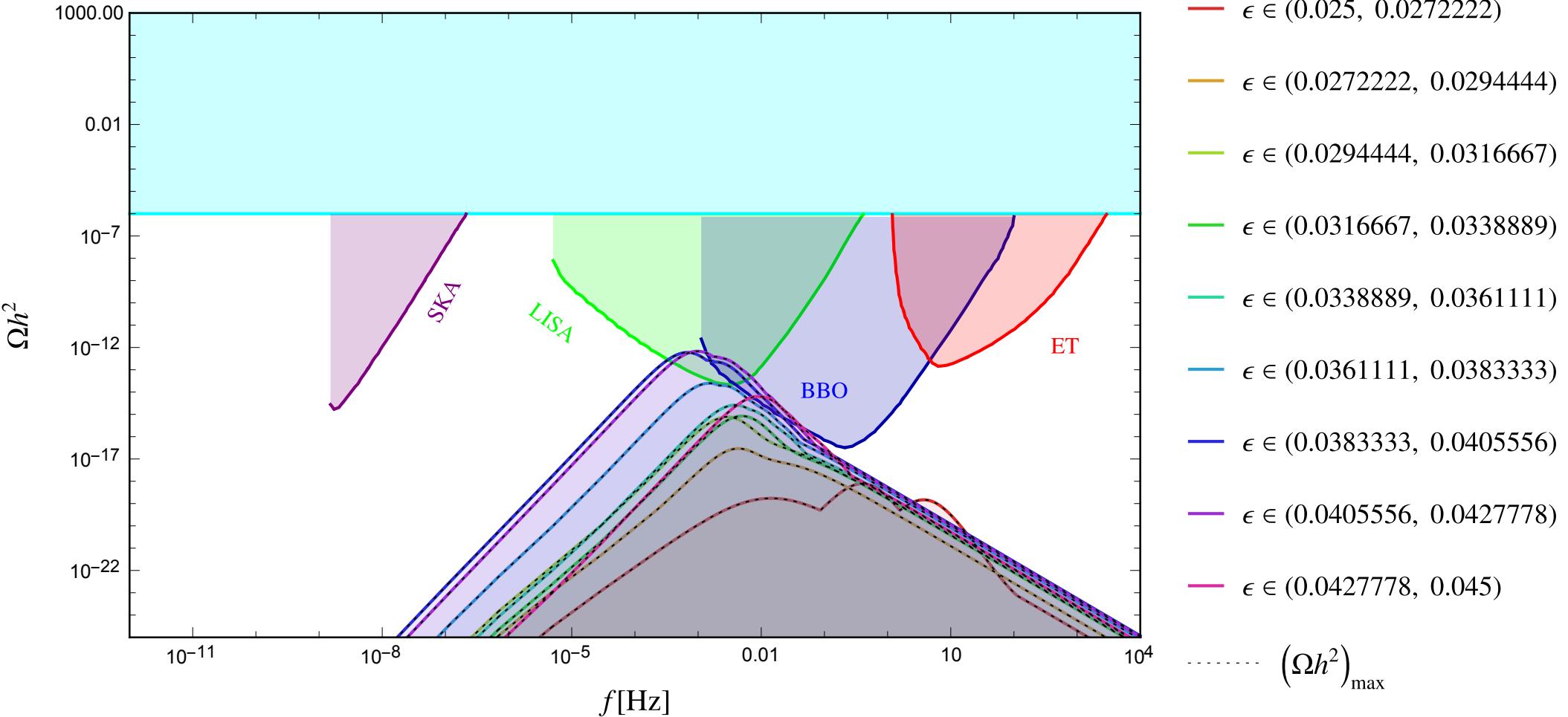
$$\tilde{h}_{\text{GW}} = \underbrace{\tilde{h}_{\text{Sound Waves}}}_{\text{neglect}} + \underbrace{\tilde{h}_{\text{Turbulence}}}_{\text{neglect}} + \underbrace{\tilde{h}_{\text{collisions}}}_{\text{neglect}}$$



$$h^2 \Omega_{\text{sw}}(f) = 2.59 \times 10^{-6} h^2 \left(\frac{100}{g_*} \right)^{1/3} \Gamma^2 \bar{U}_{fl}^4 \left(\frac{H_*}{\beta} \right) v_w \text{Max}[H_* \tau_{\text{sh}}, 1] S_{\text{sw}}(f),$$

$$h^2 \Omega_{\text{tu}}(f) = 3.35 \times 10^{-4} \left(\frac{H_*}{\beta} \right) \left(\frac{\kappa_{\text{tu}} \alpha_{T_*}}{1 + \alpha_{T_*}} \right)^{3/2} \left(\frac{100}{g_*} \right)^{1/3} v_w S_{\text{tu}}(f),$$

$$\gamma_4 = 1.4, \beta^2 = 0.1$$



Complimentarity With LHC

[arXiv: 2301.03212 , ATLAS]

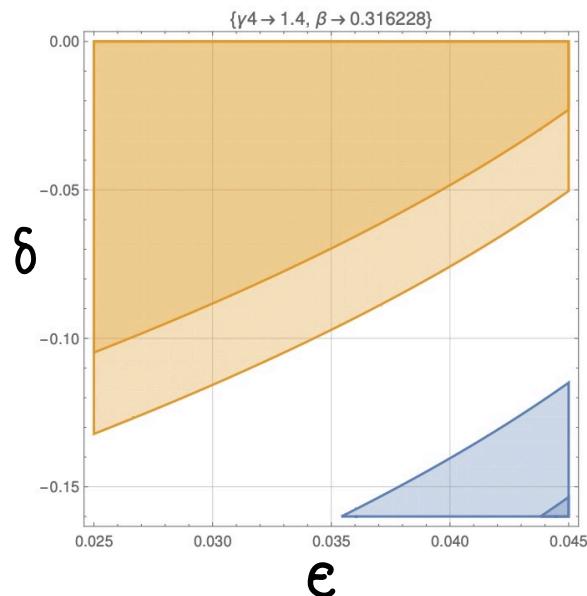
[ATLAS CONF-NOTE 2021-053]

$$\mathcal{A}(W_L^+ W_L^- \rightarrow W_L^+ W_L^+) = 1 - C_\beta^4 \gamma_a^2 \\ = -.080^{+1.0}_{-1.1}$$

$$\mathcal{A}(W_L^+ W_L^- \rightarrow hh) = -\gamma_a^2 S_\beta^2 C_\beta^2 \\ = .080^{+0.12}_{-0.13}$$

i.e we get a restriction on the ranges of δ, ϵ for a given γ_4

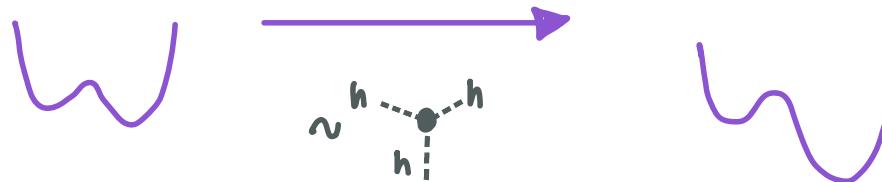
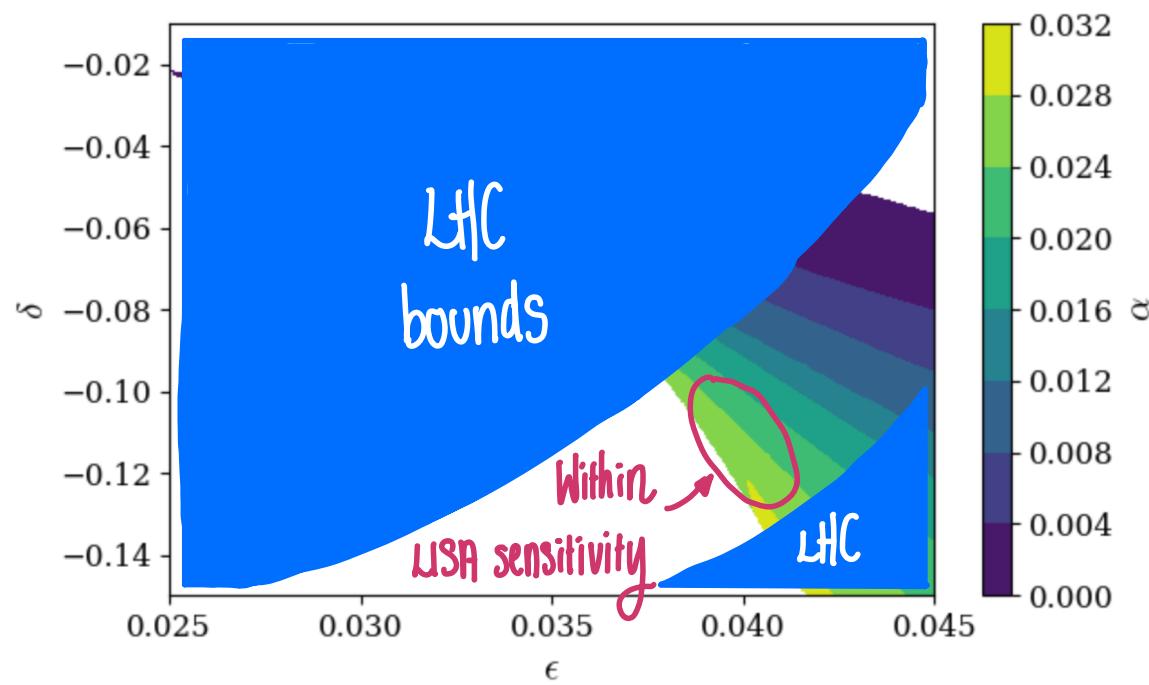
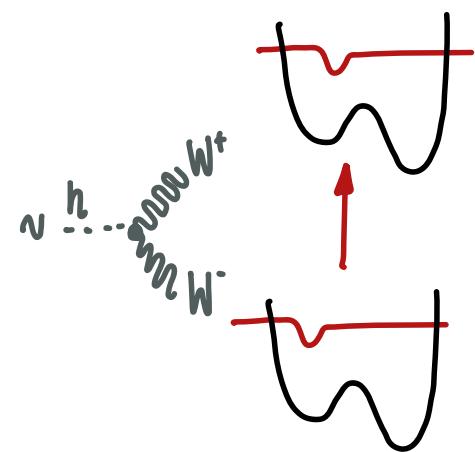
$$\delta = \begin{cases} \gamma_4^{-1} \gamma_\epsilon - \gamma_a^{-1} & \epsilon < \sqrt{9/8 - 1} \\ \gamma_4^{-1} \sqrt{2} - \gamma_a^{-1} & \epsilon > \sqrt{9/8 - 1} \end{cases}$$



Summary Plot !

$\gamma_4 = 1.4 ; \beta = \sqrt{0.1}$

(Preliminary X Schematic)



Conclusions

- Increasingly pertinent question: Which EFT should we use to describe nature?
- It is possible that cosmology could tell us about the global properties of quotient theories.
- Bubbles seem to be a promising phenomenological avenue.

Outlook / Interesting Questions...

- What happens when you change β ?



- What sort of UV physics could give a wormhole manifold?
- How about more complicated V_{tree} 's? Multi-step transitions?
Different $f(h)$'s?
- What happens when $f_{\text{boson}}(h) \neq f_{\text{fermion}}(h)$?

B
ackups

UII Physics : SM + singlet S which mix.

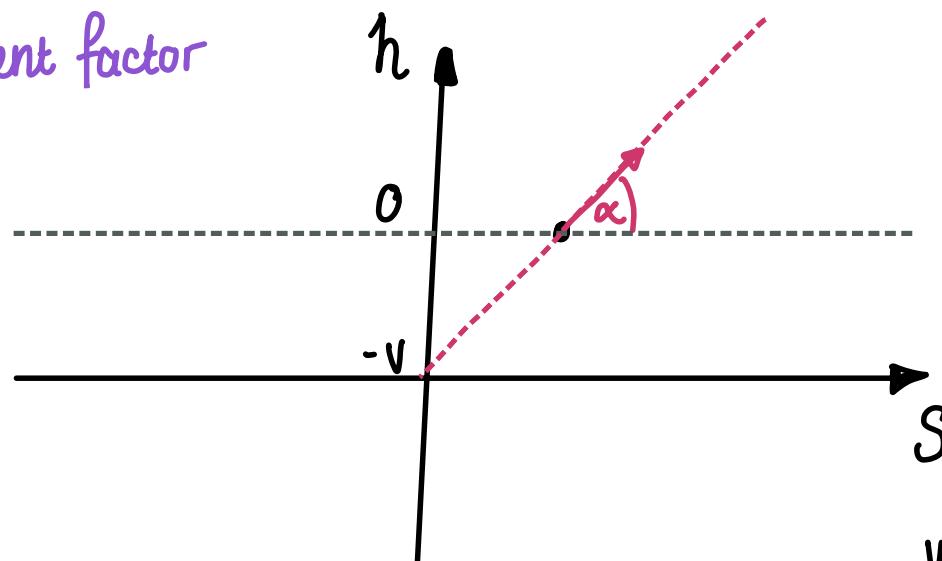
$$\mathcal{L} = \frac{1}{2} (\partial S)^2 + \frac{1}{2} D_\mu \underline{h}^T D^\mu \underline{h} - V(\underline{h}, S)$$

$$\underline{h} = \underline{h}' U(\varphi)$$

$$V^2 = V_h^2 + V_S^2 C_S$$

representation ↑

dependent factor



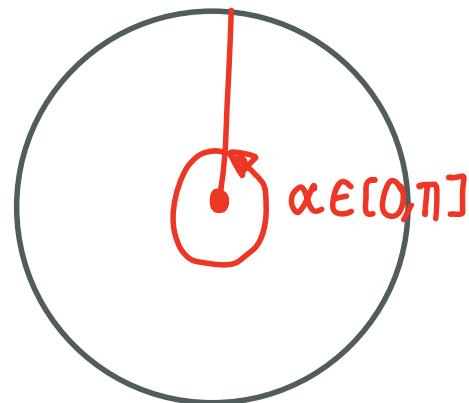
h & S mix with angle α .

$$\begin{pmatrix} h \\ S \end{pmatrix} = \begin{pmatrix} \cos\alpha & -\sin\alpha \\ \sin\alpha & \cos\alpha \end{pmatrix} \begin{pmatrix} \underline{h}' \\ \underline{h} \end{pmatrix}$$

$$V^2 F(h) \sim v^2 + 2C_B^2 \gamma_a v h + C_B^2 \gamma_a^2 h^2 + \dots$$

Sphalerons

- Static tunnelling solution to the equations of motion, between topologically distinct vacua.
- Maximum energy point of a minimum-energy non-contractible loop through scalar & gauge boson field space.



Why am I drawing pictures?

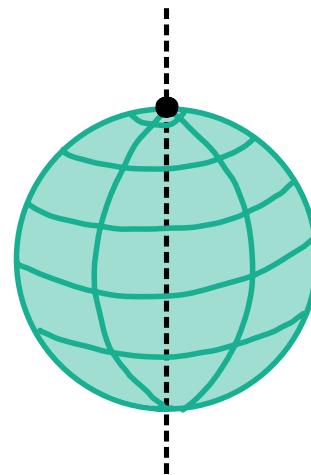
- Recall the LSZ formula where amplitudes are invariant under field-space redefinitions.
 - ↳ think coordinate redefinition invariance
 - ↳ treating the EFT as one would GR saves a lot of trouble!
- Eg 'best basis' where (ignoring $V(h)$), no 3pt $W^+ W^- h$ vertex.
 - ↗ Riemann normal coordinates
- Amplitudes relate to curvature invariants.

'Linearly realised' and 'non-linearly realised' ew symmetry is a field-space redefinition-dependent statement.

→ Think Cartesian & Polar coordinates

$$\begin{pmatrix} x \\ y \\ z \end{pmatrix} \xrightarrow{R} \begin{pmatrix} x \\ y \\ z \end{pmatrix}$$

$$\begin{aligned} r &= \sqrt{x^2 + y^2 + z^2} \rightarrow r \\ (\phi) &\rightarrow (a\phi) \\ (\theta) &\rightarrow (b\theta) \end{aligned}$$



SMEFT theories admit a linear representation.