Large Neutrino Masses in Cosmology:
DARK SECTOR TO THE RESCUE

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22nd May 2025, KIAS HEP Seminar

Based on:
JCAP 04 (2025) 054 [Cristina Benso, Thomas Schwetz, **DV**]
arXiv: 2410.23926

Dark Matter in our Universe

Properties

Mass and number density unknown

DM

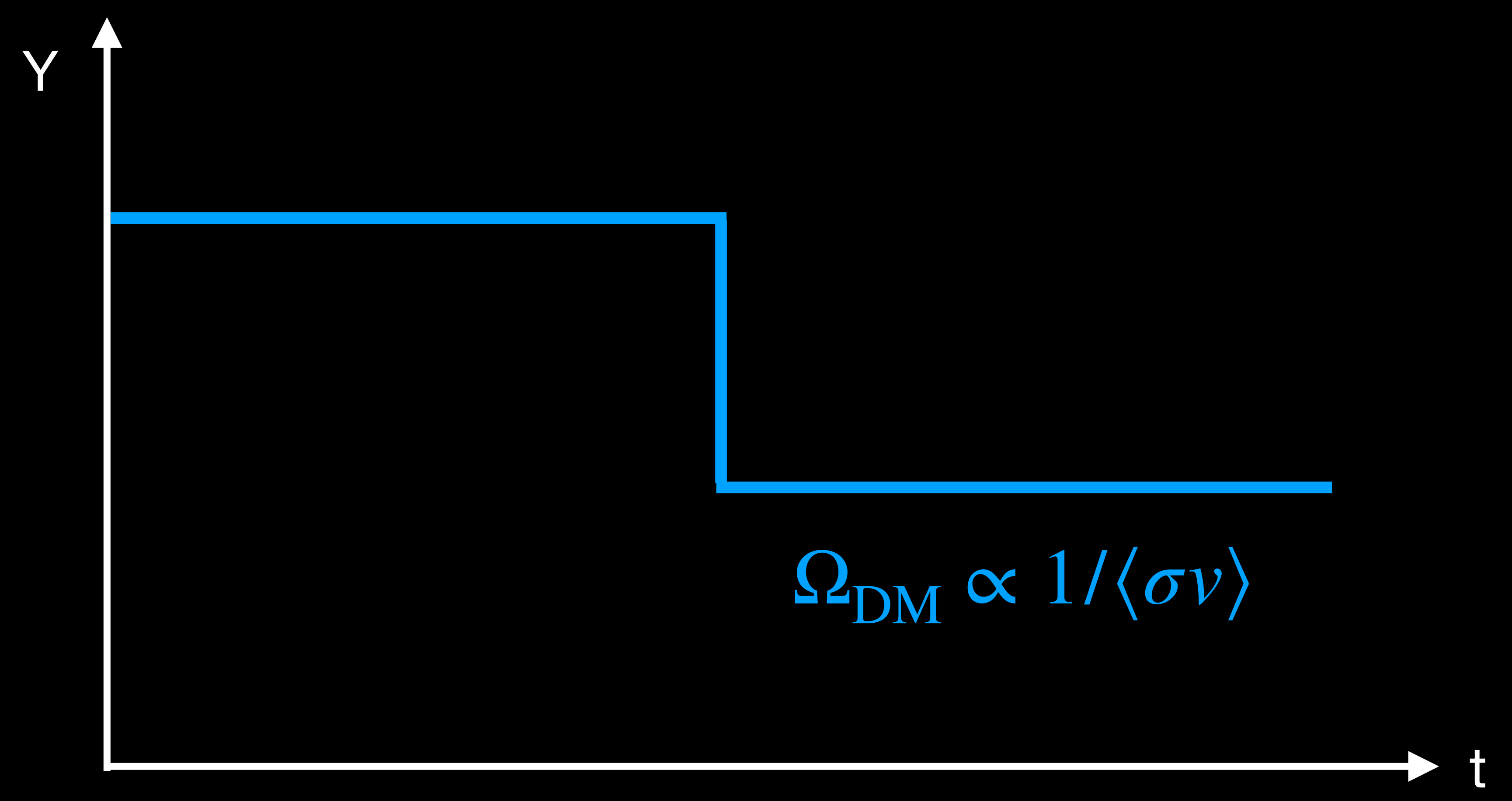
Theoretical limits:
Fermions - > keV
Bosons - 10^{-22} eV

- **Stable** — Long lived to produce the observed relic
- **Neutral** — Does not interact with photons
- **Cold** — Non relativistic (at least around T ~ 1 eV)
- **Non-baryonic** — Observations from BBN and CMB
- **Collisionless** — Very weak self-interactions

Dark Matter in our Universe

Thermal relics

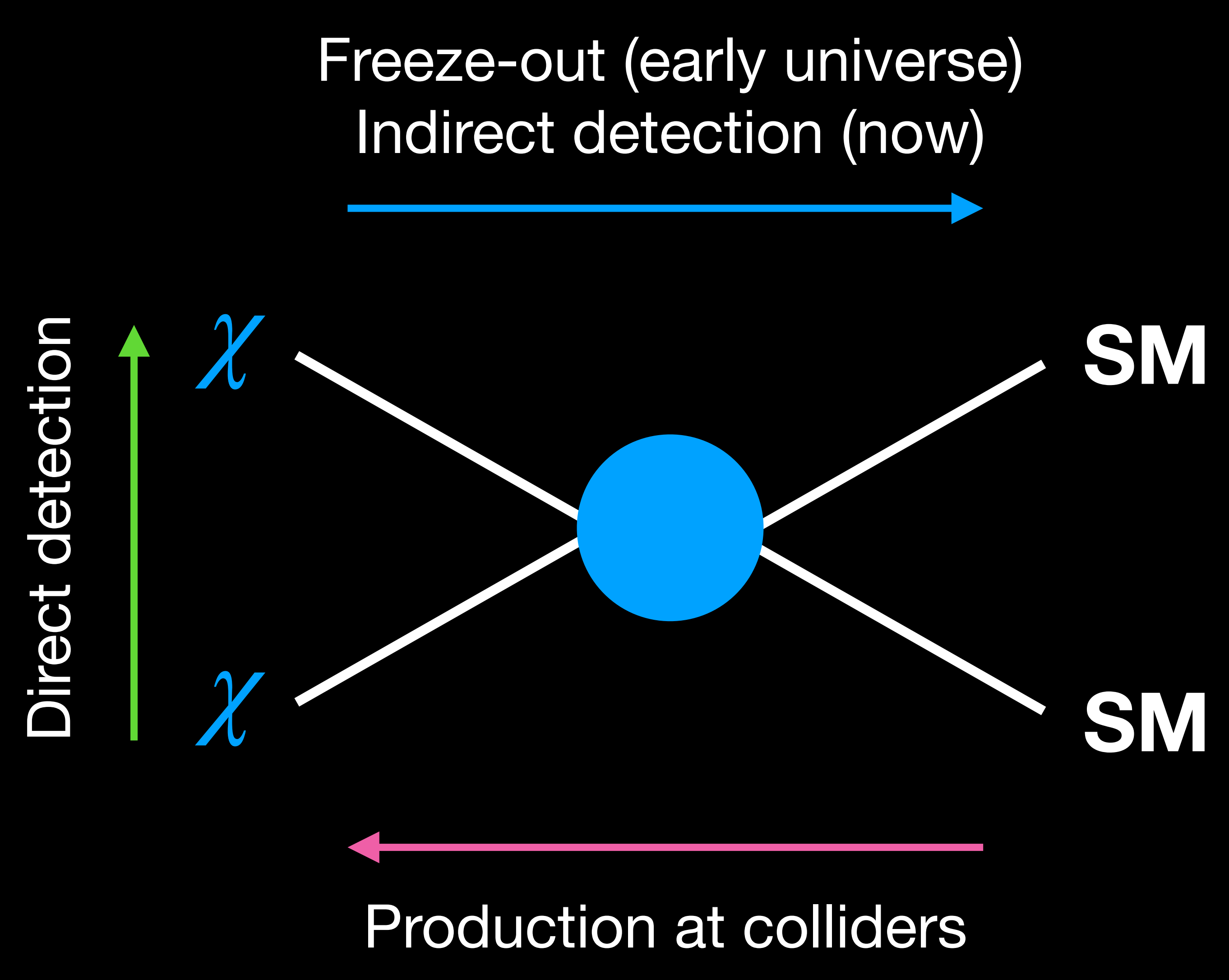
Thermal freeze-out of Weakly Interacting Massive Particles



$\langle\sigma v\rangle \sim 2.2 \cdot 10^{-26}\,\mathrm{cm^3/s}$

Produces the correct relic abundance

WIMPS



Dark Sectors?

DM vs. Visible Sector

Particle DM

Only **one** DM state makes up the entire dark sector

DM χ $\quad \mathbf{Z}_2$ / $U(1)_{dark}$

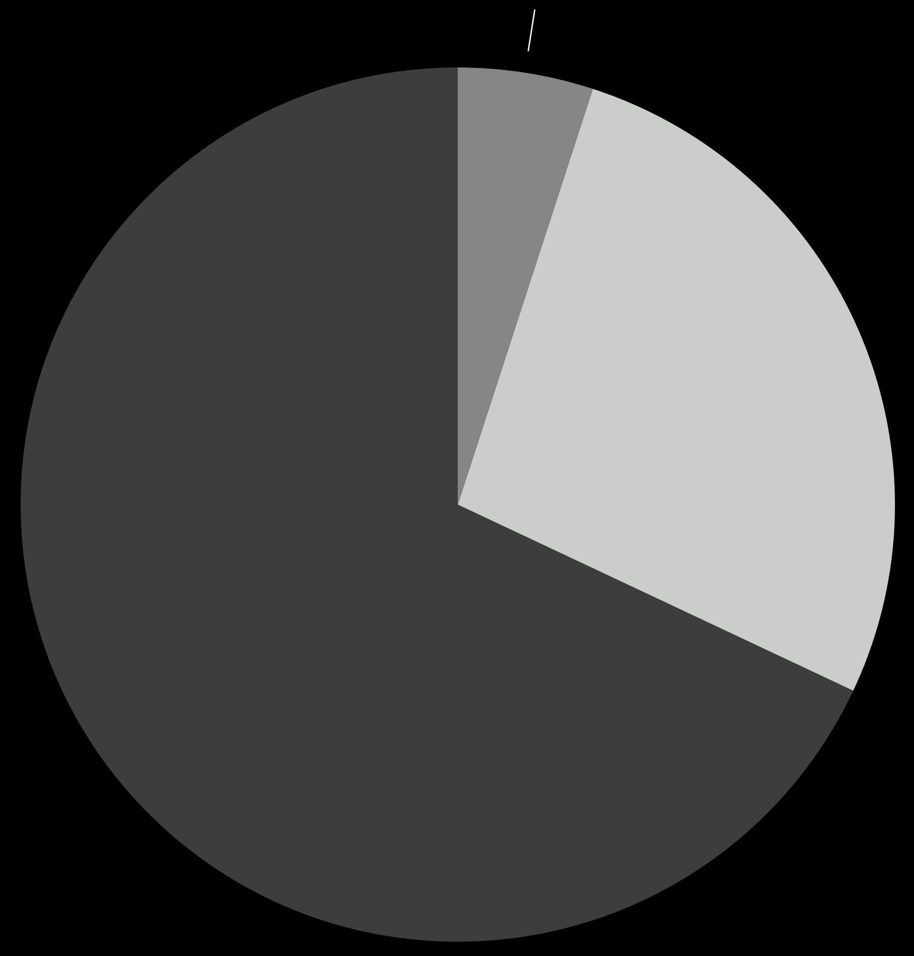
DM states are **symmetric,** abundance set by freeze-out/freeze-in

Visible Sector

Several stable components - stabilising symmetries

Electrons	Electric charge
Protons	Baryon number
Neutrinos	Spin
Photons	Poincare

Abundance set by the **baryon asymmetry** of the universe $\eta_B \sim 10^{-10}$



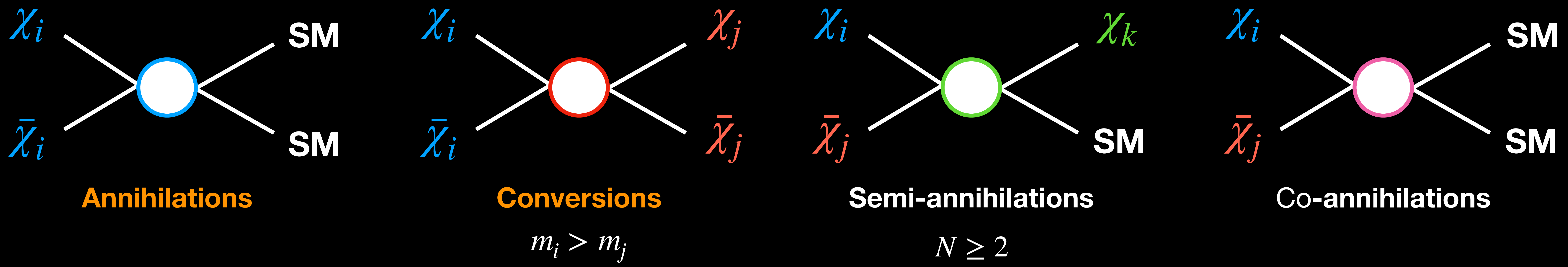
$\rho_{DM} = 5\rho_B$

Why should there be just one particle on the dark side?

Multi-component Dark Sectors

Generic features

Many degrees of freedom → e.g. ***N***-component dark sector, may have the following interactions:



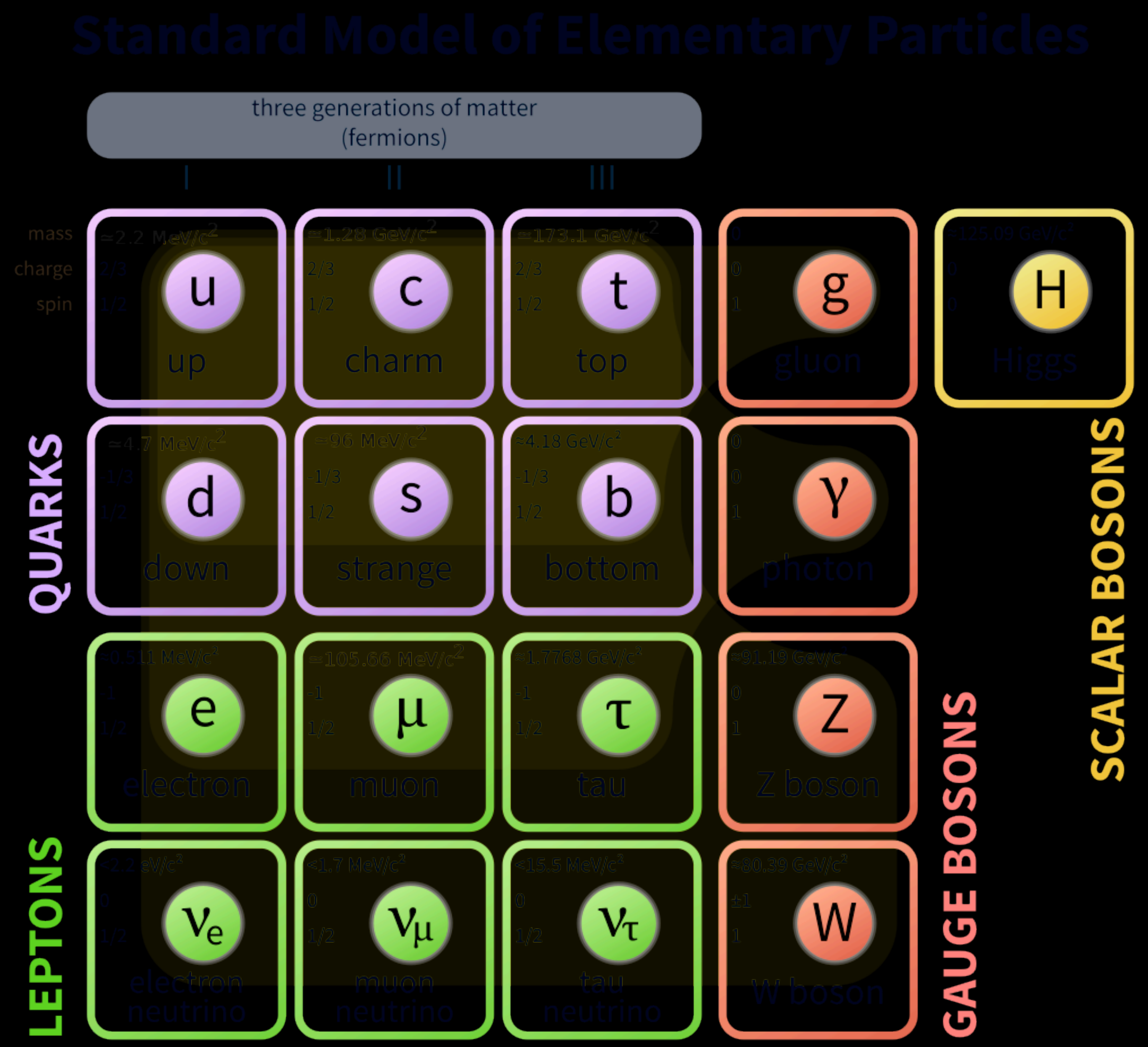
Complex system → behaviour characterised by power laws and exponentials

A. Bas, J. Herrero-Garcia, **DV**
JHEP 10 (2022) 075

Relaxation of existing bounds on direct/indirect detection, collider searches, etc.

Neutrinos

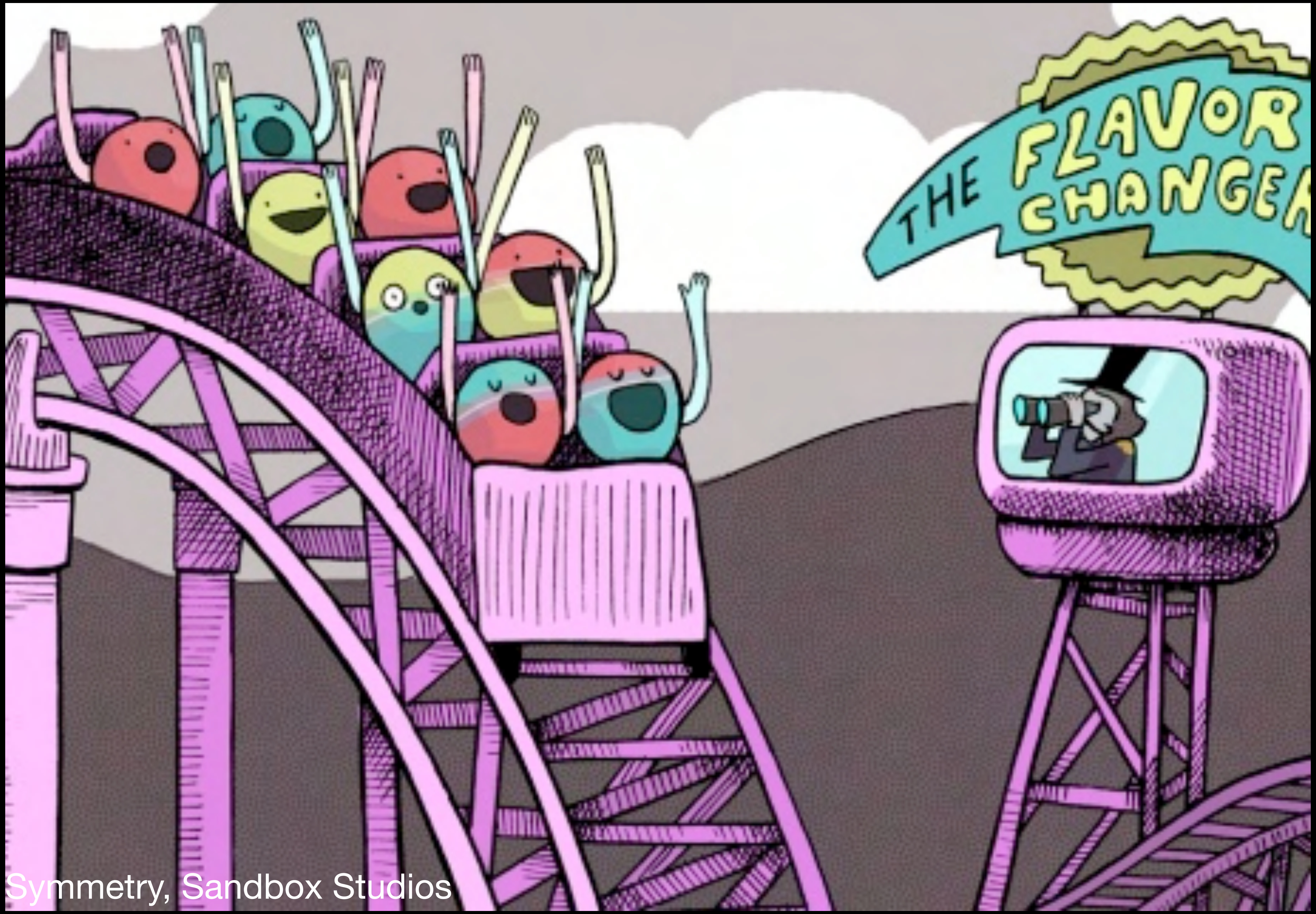
In the SM



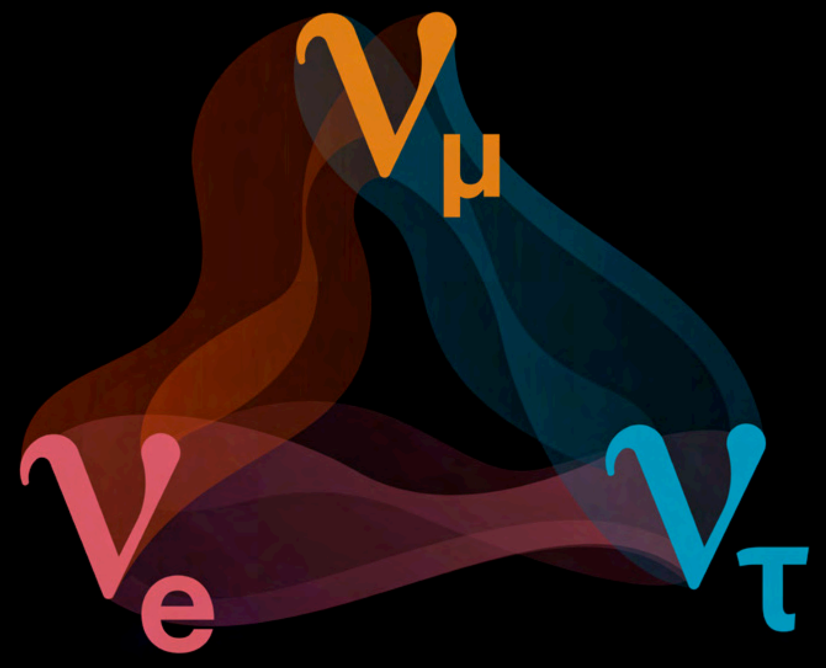
ν

- Right-handed partners missing
- Masses absent
- Electrically neutral
- Interact only via weak interactions
- Very difficult to detect

Neutrinos
Oscillations



Symmetry, Sandbox Studios

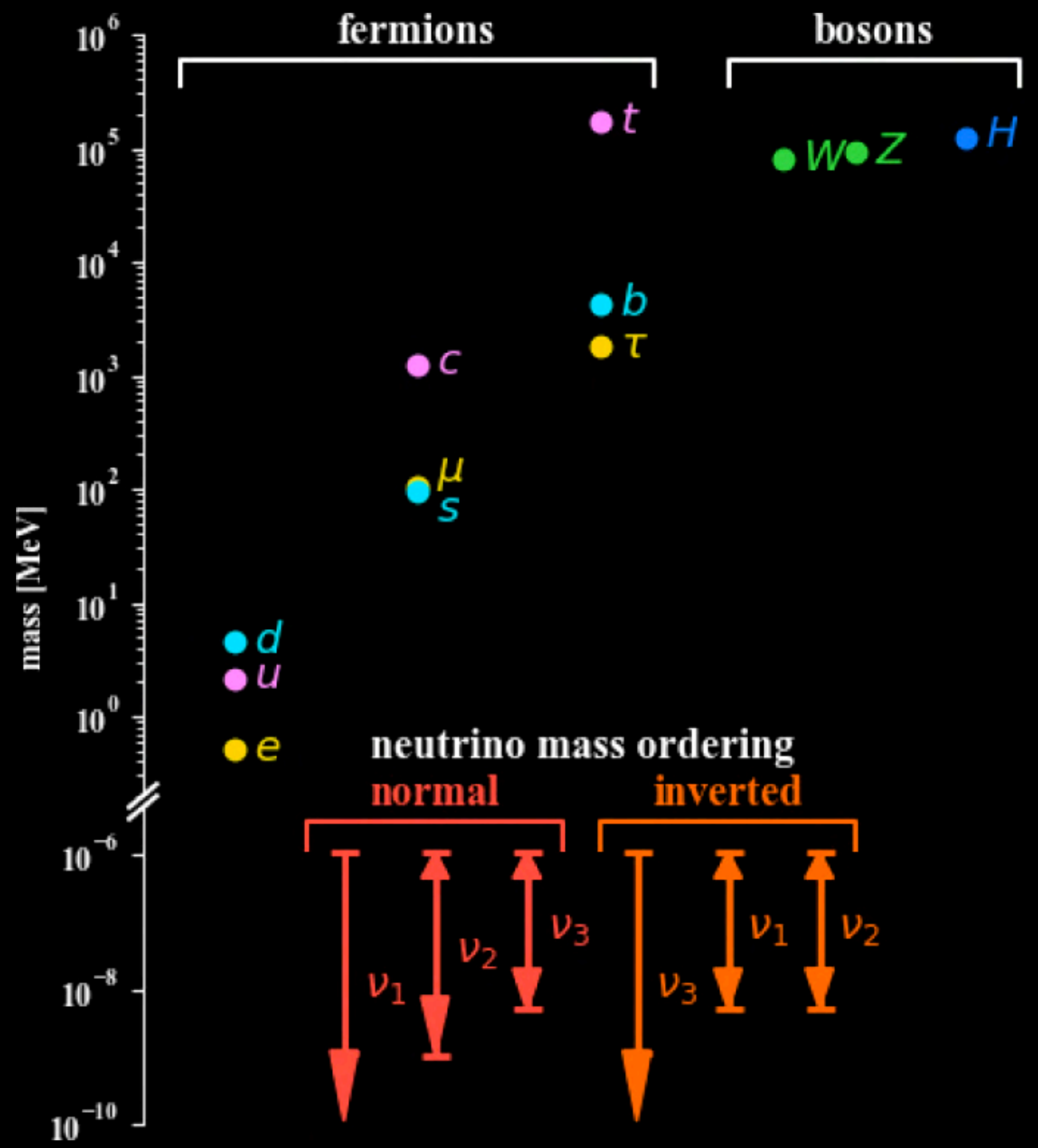


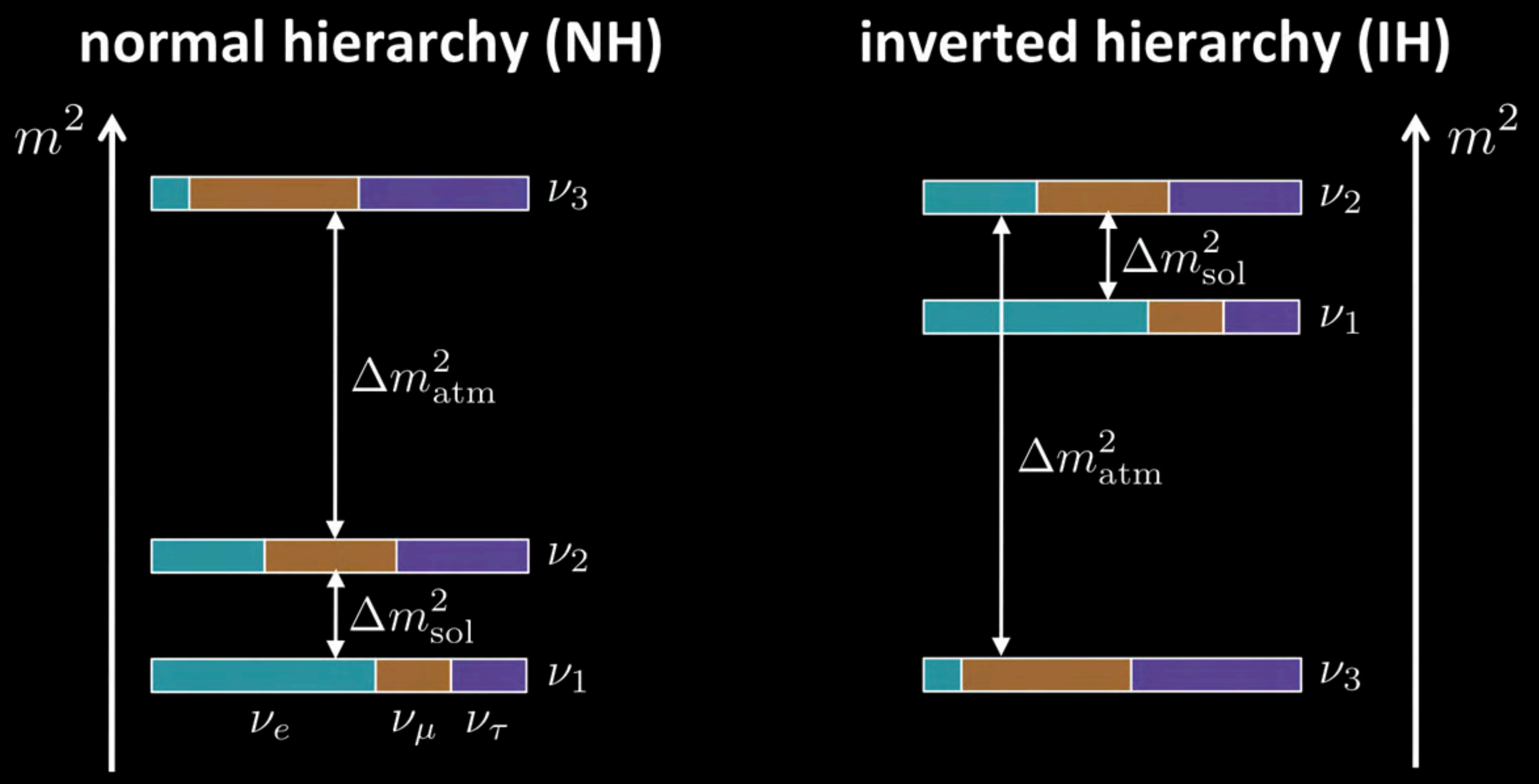
Neutrinos can change their flavour during propagation

Neutrino Oscillations → Neutrinos have a tiny mass

Neutrinos

Mass spectrum





Origin of mass unknown!

Absolute mass unknown!

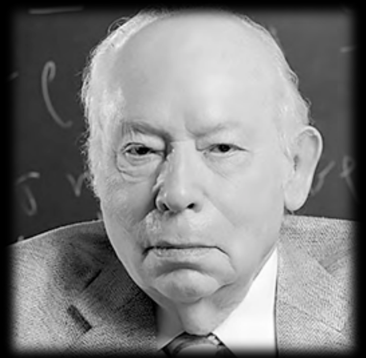
Neutrino Masses

Effective Field Theory

Higher-dimensional operators respecting SM gauge symmetries

$$\mathscr{L} = \mathscr{L}_{\text{SM}} + \frac{c}{\Lambda_{\text{NP}}}\mathcal{O}^{d=5} + \frac{c'}{\Lambda^2_{\text{NP}}}\mathcal{O}^{d=6} + \frac{c''}{\Lambda^3_{\text{NP}}}\mathcal{O}^{d=7} + \dots$$

New Physics energy scale



Unique operator at $d = 5$

Weinberg: PRL 43 (1979)

$$\mathscr{L}_5 = \frac{c_5}{\Lambda_{NP}} LLHH$$

EWSSB

$\langle H \rangle = v$

$$m_\nu \sim \frac{c_5}{\Lambda_{NP}} v^2 \gtrsim 0.05 \times 10^{-9} \text{ GeV}$$

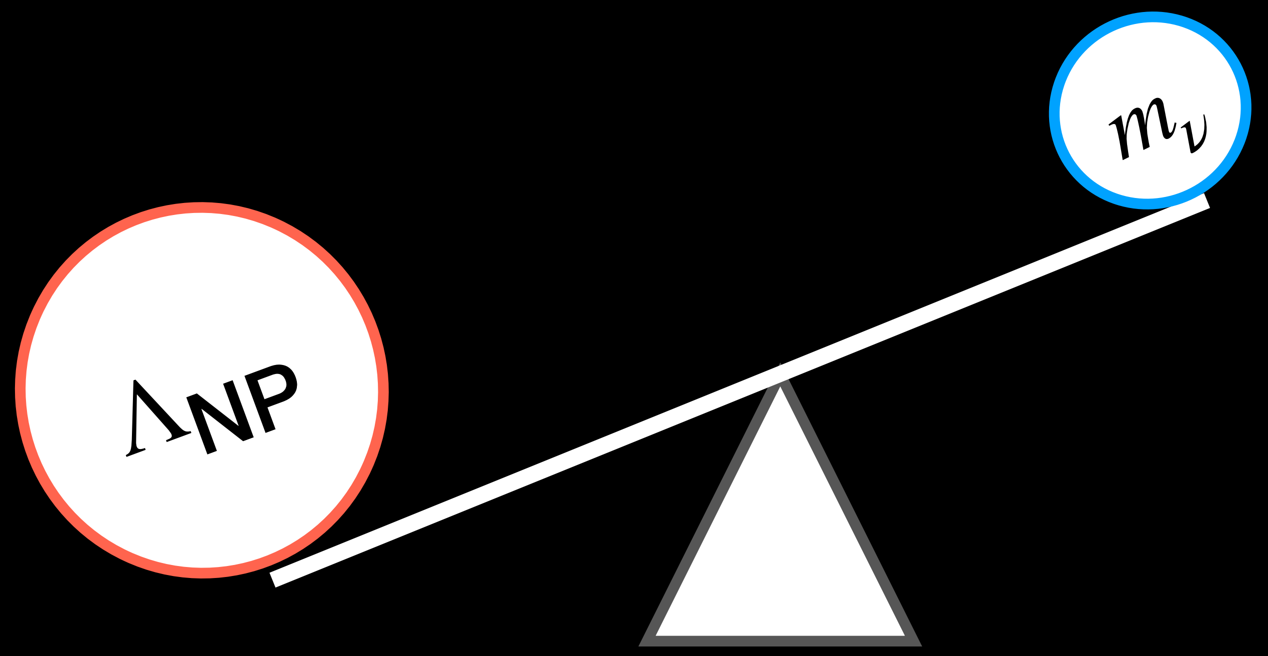
174 GeV

$\lesssim 10^{14}$ GeV

Close to the GUT scale

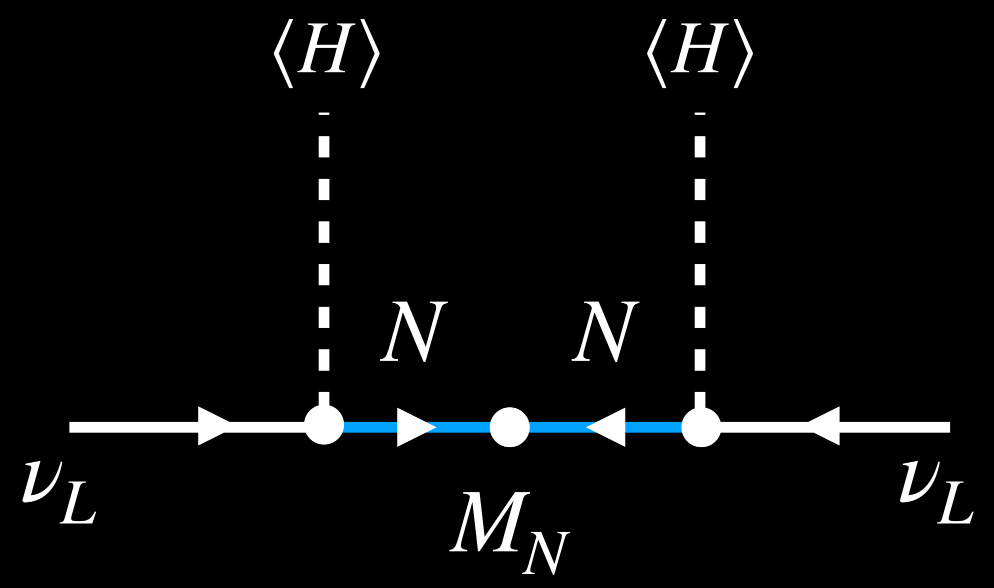
Neutrino Masses

Seesaw mechanism



UV completions of the Weinberg operator at the tree level → Usual Seesaws

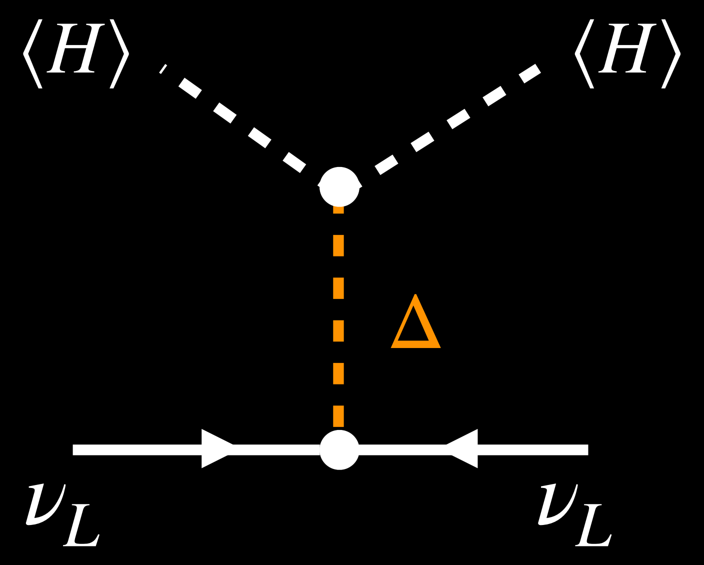
Fermion singlet: N



I

Minkowski (1977); Yanagida (1980); Gell-Mann, Raymond, Slansky (1979), Mohapatra, Senjanovic (1980)

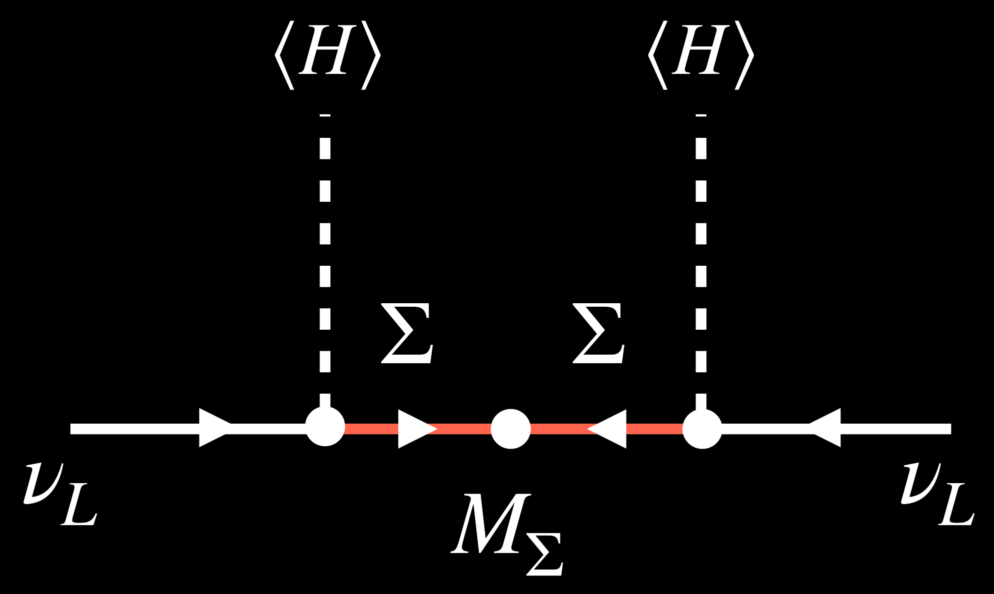
Scalar triplet: Δ



II

Schechter, Valle (1980); Lazarides, Shafi, Wetterich (1981); Mohapatra, Senjanovic (1981)

Fermion triplet: Σ



III

Foot, Lew, He, Joshi (1989)

Neutrino Mass Bounds

Oscillations

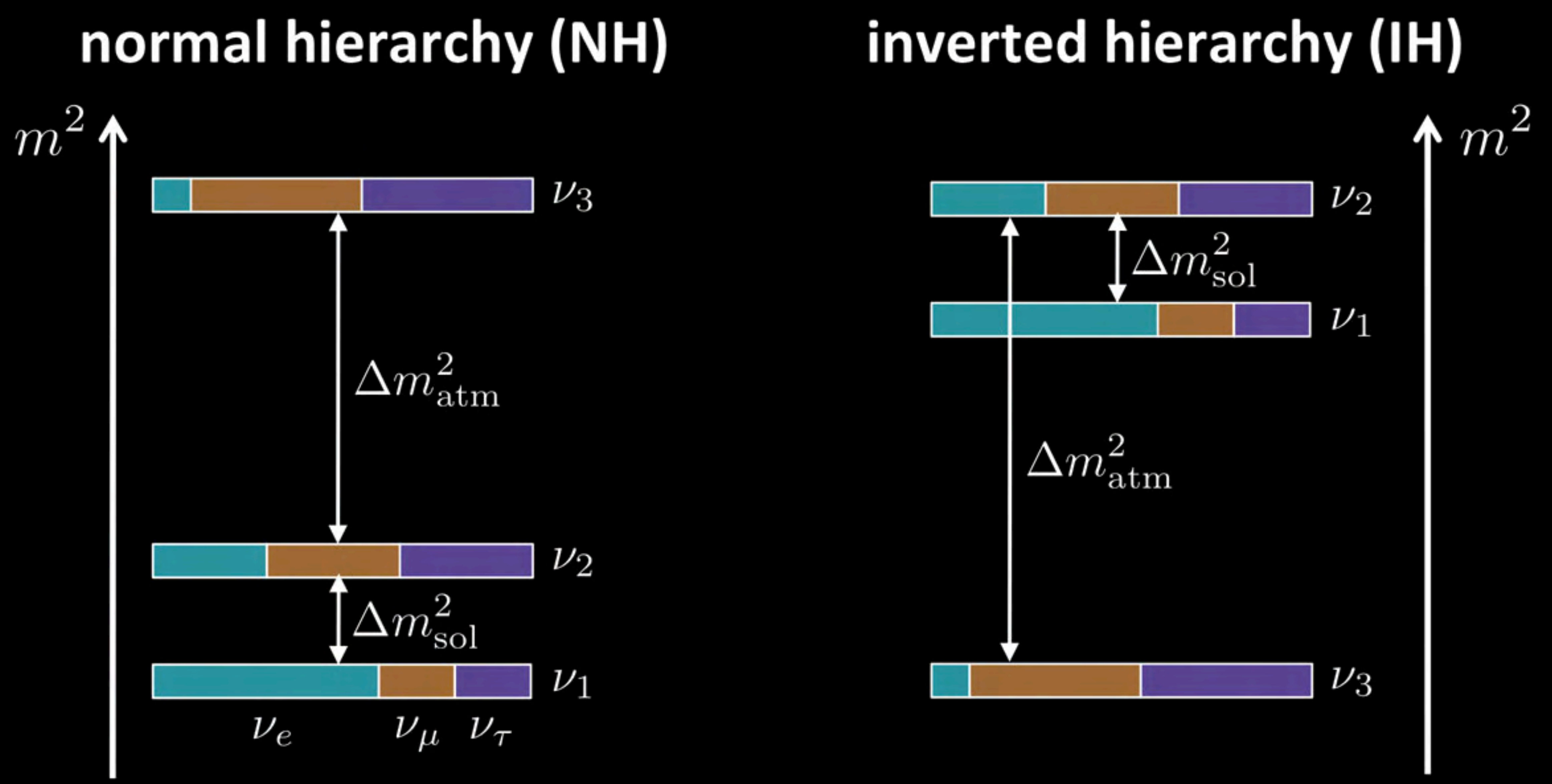
$$\sum_{i=1}^{3} m_i = \begin{cases} m_0 + \sqrt{\Delta m_{21}^2 + m_0^2} + \sqrt{\Delta m_{31}^2 + m_0^2} & \text{(NO)} \\ m_0 + \sqrt{|\Delta m_{32}^2| + m_0^2} + \sqrt{|\Delta m_{32}^2| - \Delta m_{21}^2 + m_0^2} & \text{(IO)} \end{cases}$$

$$m_0 \to 0$$

NuFit Collaboration

$$\sum m_\nu > \begin{cases} 0.058 \text{ eV} & \text{Normal ordering} \\ 0.098 \text{ eV} & \text{Inverted ordering} \end{cases}$$

95% CL



Neutrino Mass Bounds

Terrestrial Experiments

Experiments looking for absolute neutrino mass scale and $0\nu\beta\beta$

KATRIN Experiment

$\sum m_\nu \lesssim 1.35$ eV

KamLAND-ZEN Experiment

$m_{\text{lightest}} < 0.084 - 0.353$ eV

Neutrinos

In cosmology

$$n_\nu = \frac{g}{(2\pi)^3}\int d^3p\, f_\nu(E, T_\nu)$$

$$\rho_\nu = \frac{g}{(2\pi)^3}\int d^3p\, E\, f_\nu(E, T_\nu)$$

$T \gg \mathrm{MeV}$

Weak interactions and ν's in equilibrium

$n + \nu_e \leftrightarrow p + e^-$

$$f_\nu(E, T_\nu) = \frac{1}{1 + \exp[(E-\mu)/T_\nu]}$$

Fermi-Dirac distribution

$T \sim \mathrm{MeV}$

Weak interactions drop out: $\Gamma \sim G_F^2 T^5 < H$

ν's decouple while relativistic at $T \simeq 2$ MeV

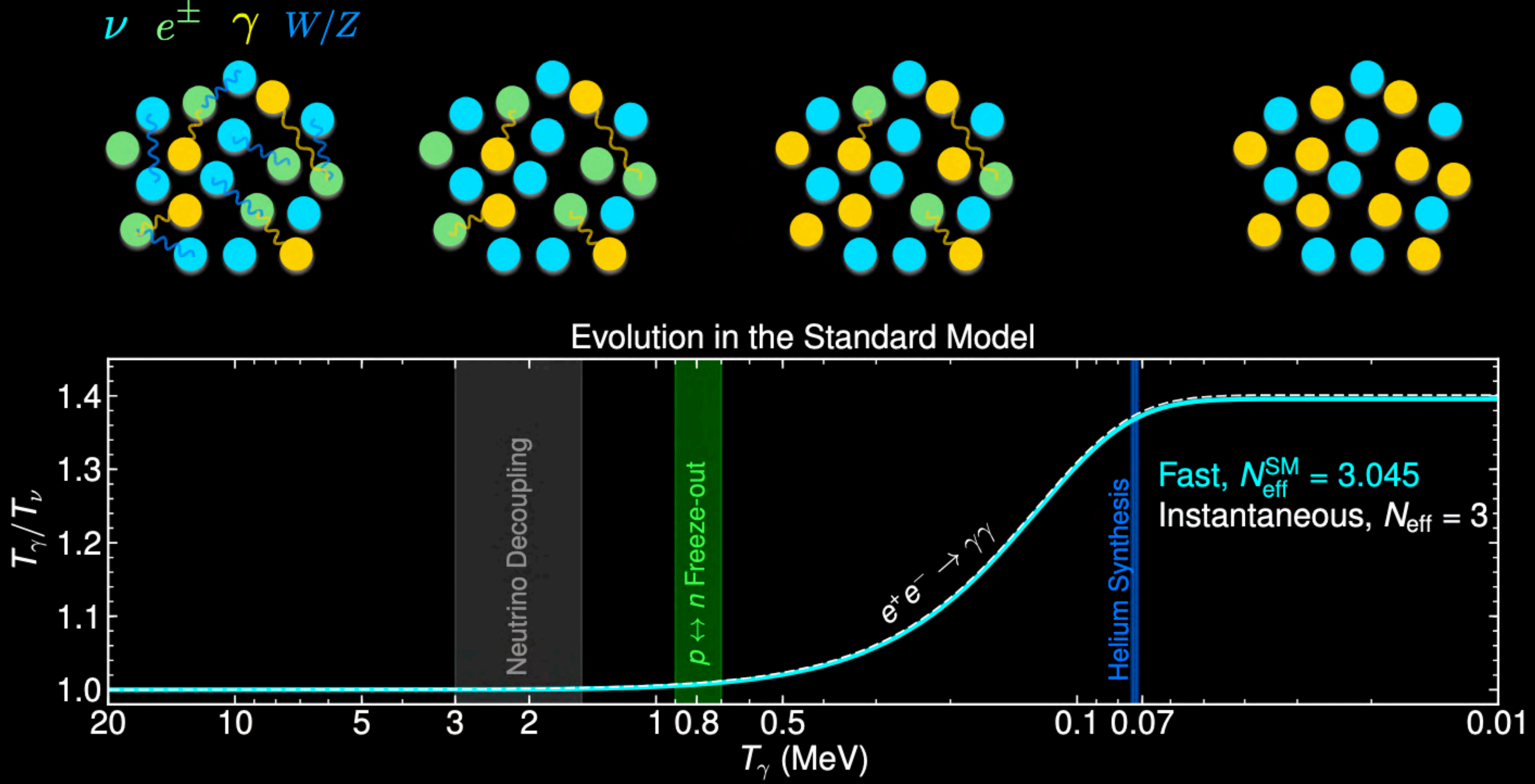
$$f_\nu(E, T_\nu) = \frac{1}{1 + \exp[p/T_\nu]}$$

Well-approximated

$$\sum_i \rho_{\nu,i} \equiv N_{\mathrm{eff}}\rho_{\nu,0}$$

Neutrinos

In cosmology



$n_{\nu,0} = 56\ \mathrm{cm}^{-3}$

$$\frac{T_\nu}{T_\gamma} = \left(\frac{4}{11}\right)^{1/3}$$

Escudero: 2001.04466

Neutrinos

In cosmology

$$\Omega_\nu = \frac{\rho_\nu}{\rho_{\text{critical}}} \propto \sum \langle E_\nu \rangle n_\nu$$

$T_{\nu,0} \gg m_\nu$

$\langle E_\nu \rangle \simeq \langle \rho_\nu \rangle$, energy density characterised by $N_{\text{eff}} \propto \langle p_\nu \rangle n_\nu$

$$N_{\text{eff}} \equiv \frac{8}{7}\left(\frac{11}{4}\right)^{4/3}\left(\frac{\rho_{\text{rad}} - \rho_\gamma}{\rho_\gamma}\right)$$

$3.044(1)$
SM prediction

2.99 ± 0.17
PLANCK 2018

$T_{\nu,0} \ll m_\nu$

$\langle \rho_\nu \rangle = m_\nu,\ \rho_\nu = \sum m_\nu n_\nu$

ν's contribute to expansion rate as DM

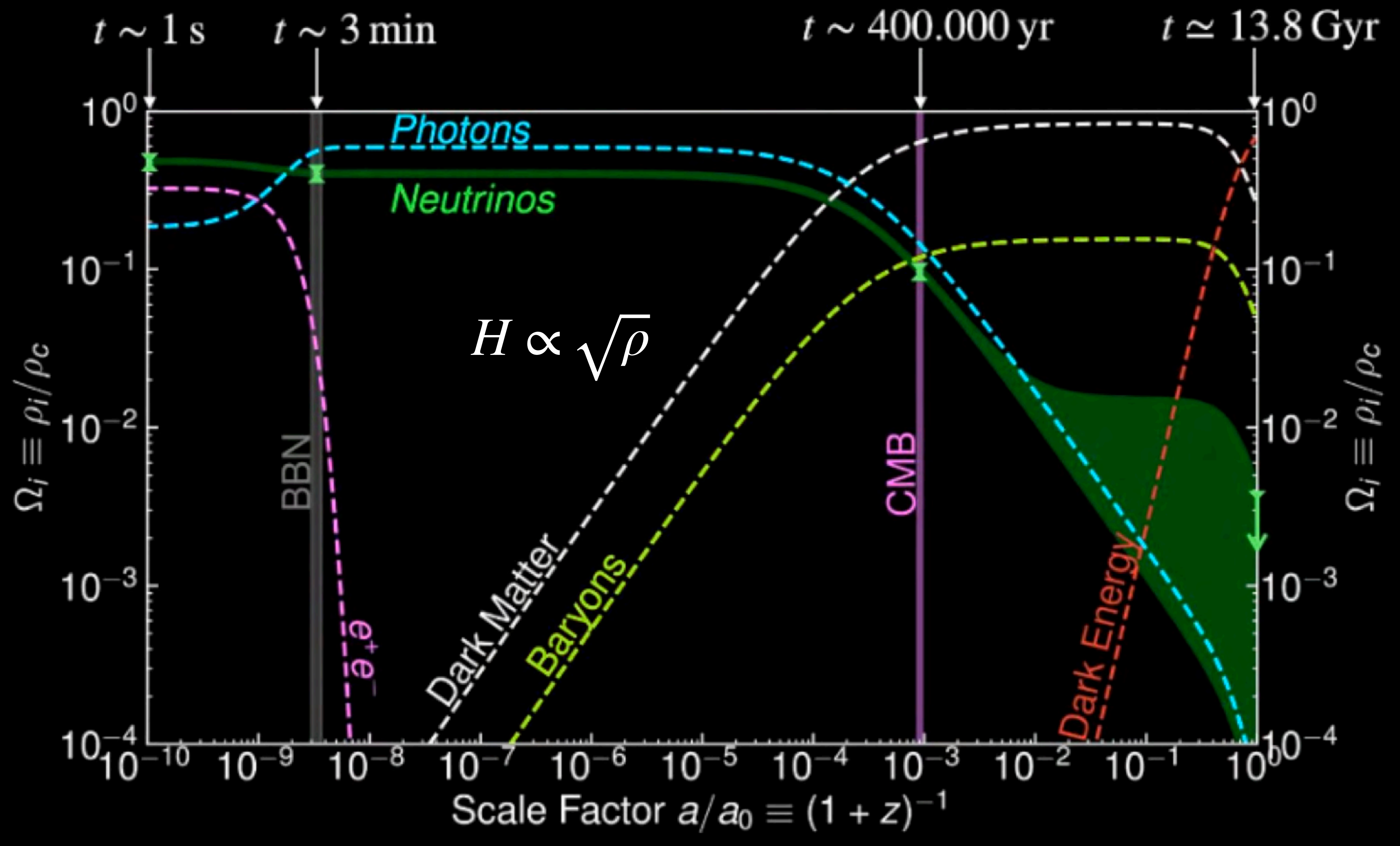
$$\Omega_\nu h^2 \equiv \frac{\sum m_\nu n_\nu^0 h^2}{\rho_{\text{critical}}} < 1.3 \times 10^{-3}\ (95\ \%\ \text{CL})$$

PLANCK 2018

Neutrino Masses

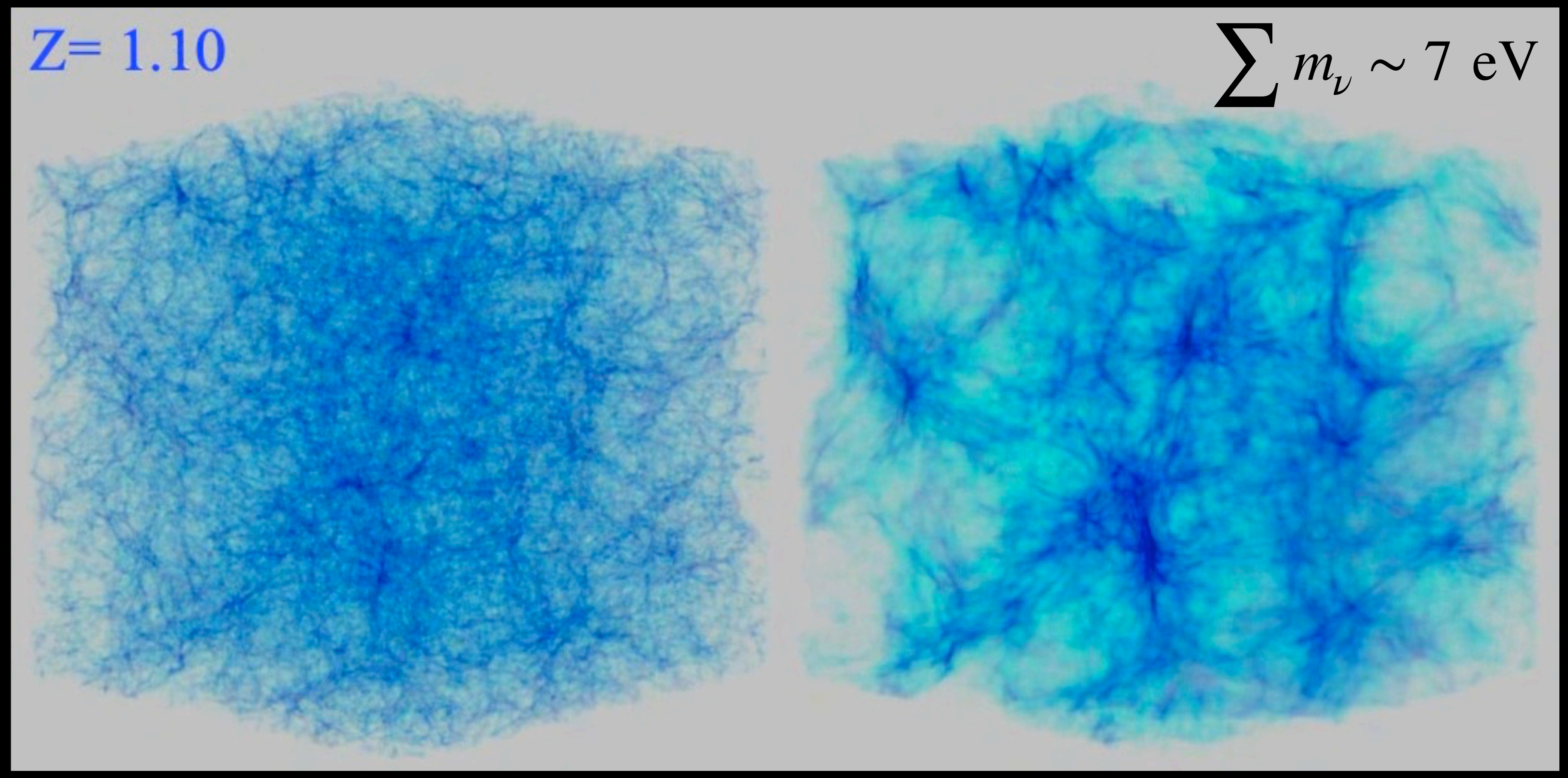
Cosmology

Alter the expansion history of the universe near matter-radiation equality epoch



Credit: Miguel Escudero

Free-streaming affects the growth of structures at late times

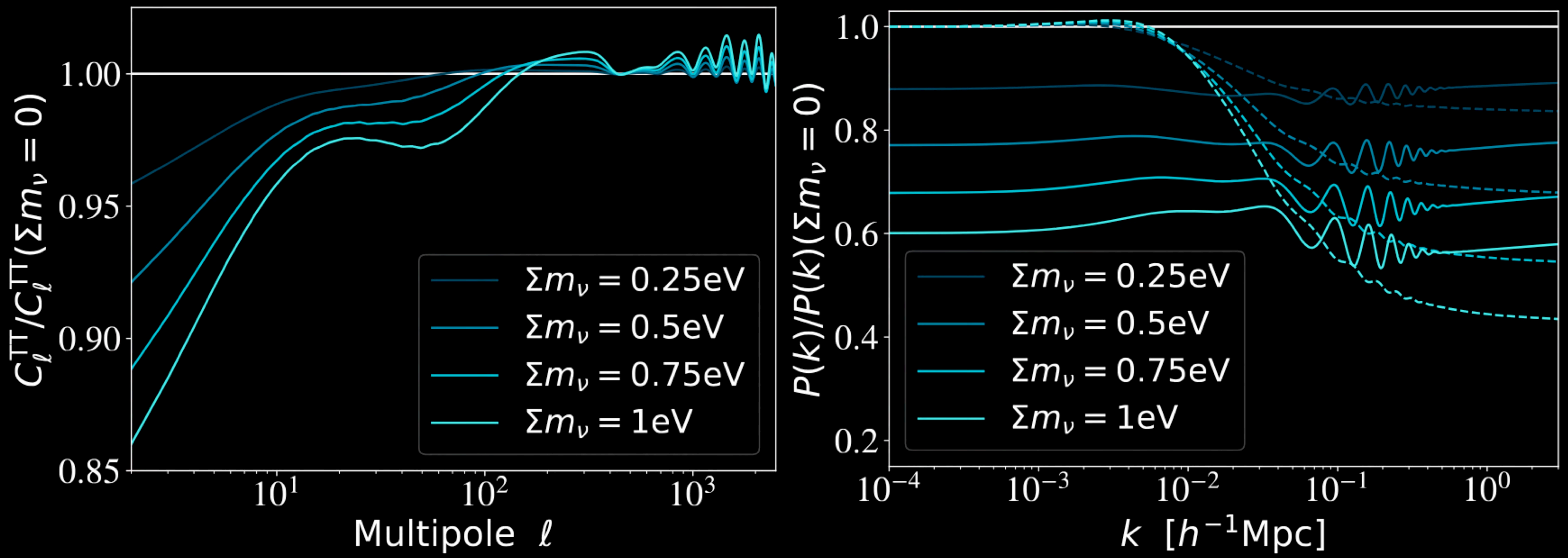


Credit: Troels Haugbølle

Cosmological observables → CMB + LSS

Neutrino Masses

Cosmology



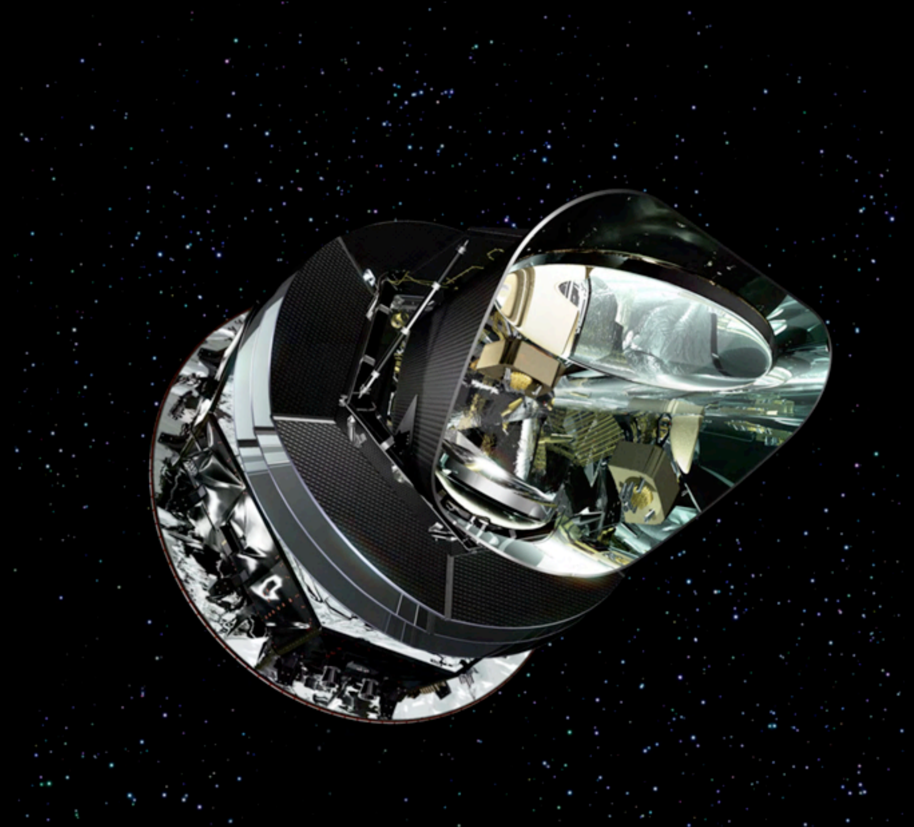
Credit: PDG 2022; Lesgourges, Verde

Neutrino Mass Bounds

Cosmology

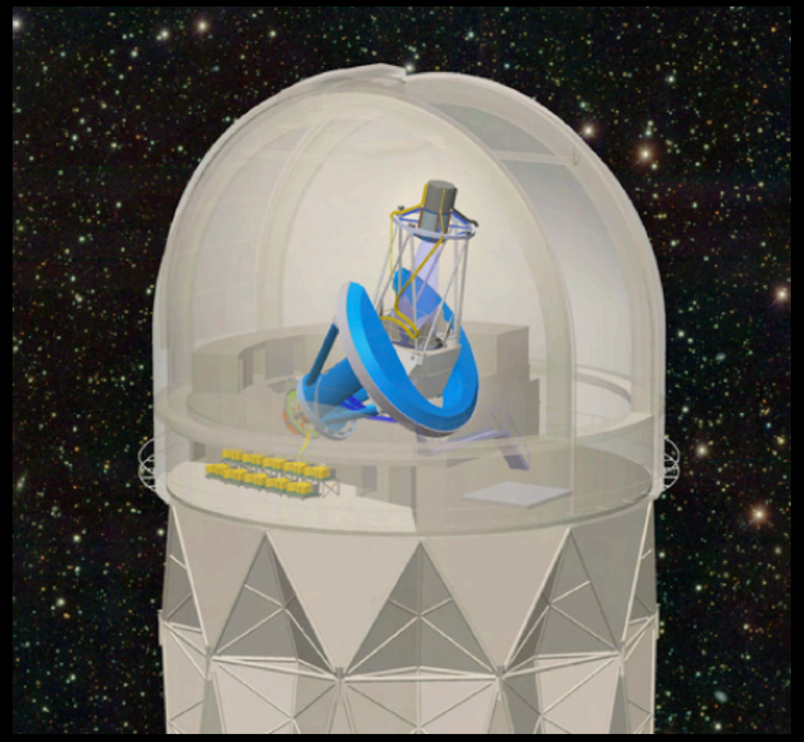
$m_\nu \neq 0 \rightarrow$ Cosmological Implications $\rightarrow$ Suppression of growth of small scale structures; Affect on CMB anisotropies

$\sum m_\nu \equiv \sum_{i=1}^{3} m_i$ can be constrained from cosmological surveys



PLANCK CMB+BAO (2018)

$$\sum m_\nu < 0.12 \text{ eV}$$



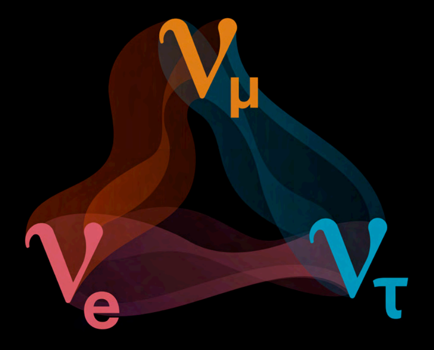
DESI 2025 + CMB

$$\sum m_\nu < 0.064 \text{ eV}$$

Neutrino mass bounds from cosmology keep getting stronger

Neutrino Mass Bounds

Oscillations vs. Cosmology



Oscillations

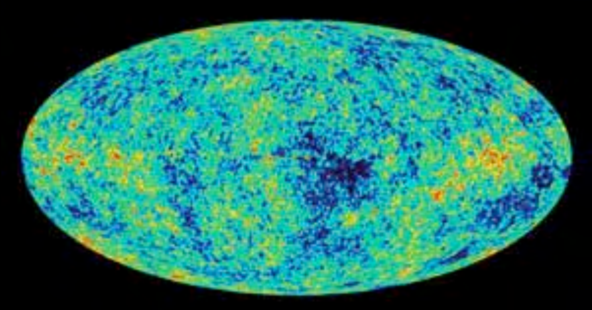
NuFit Collaboration

$$\sum m_\nu > \begin{cases} 0.058 \text{ eV} & \text{Normal ordering} \\ 0.098 \text{ eV} & \text{Inverted ordering} \end{cases}$$

95% CL

$$\sum m_\nu < \begin{cases} 0.12 \text{ eV} & \text{PLANCK CMB+BAO (2018)} \\ 0.064 \text{ eV} & \text{DESI 2025 + CMB} \end{cases}$$

95% CL



Cosmology

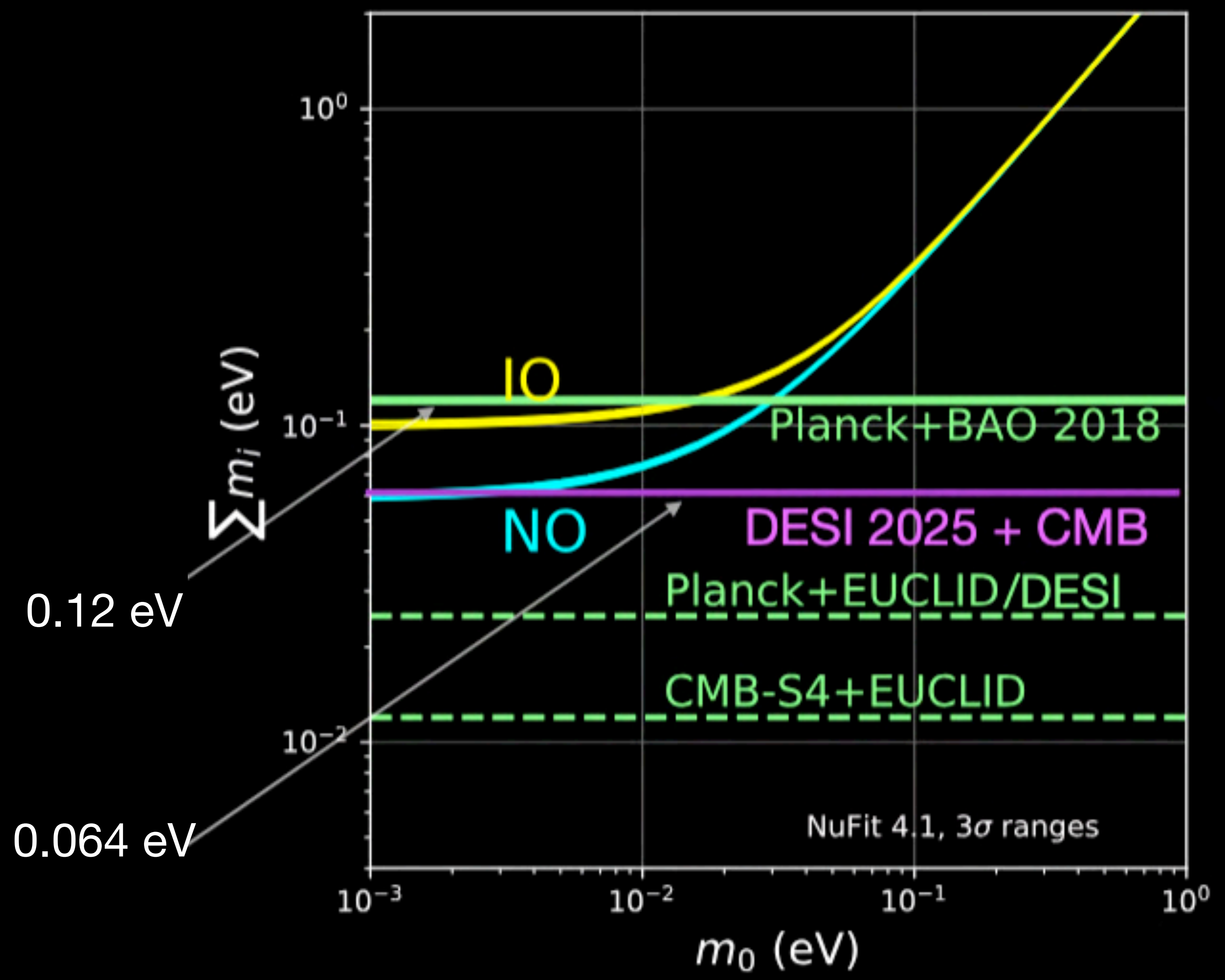
DESI bound is in significant tension with IO, very close to NO

2-zero neutrino mass textures → $\sum m_\nu > 0.12$ eV

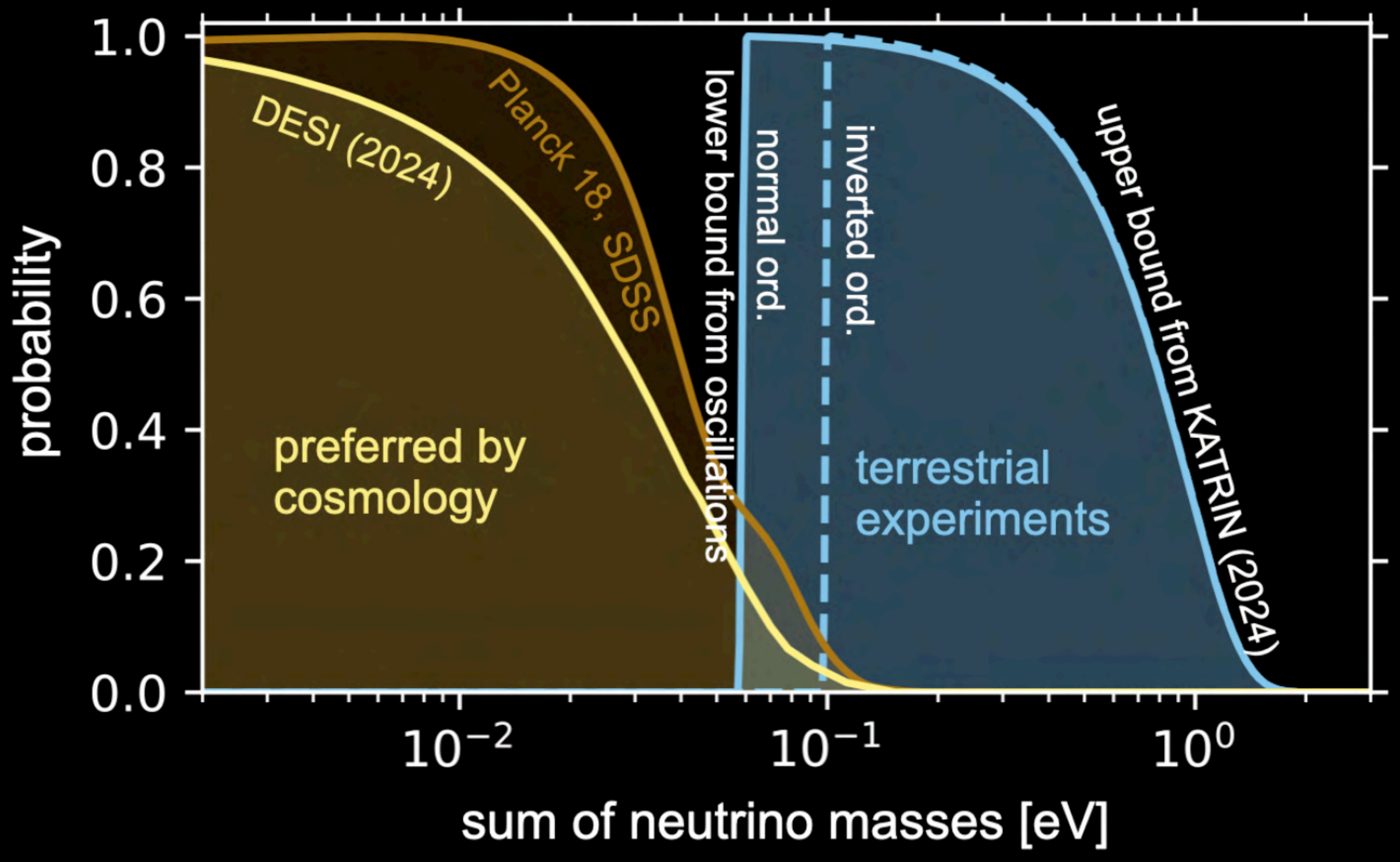
Alcaide et al: 1806.06785;
Lattanzi et al: 2007.01650

Neutrino Mass Bounds

Laboratory vs. Cosmology



Courtesy: Thomas Schwetz (Durham 2025); updated from 2302.14159



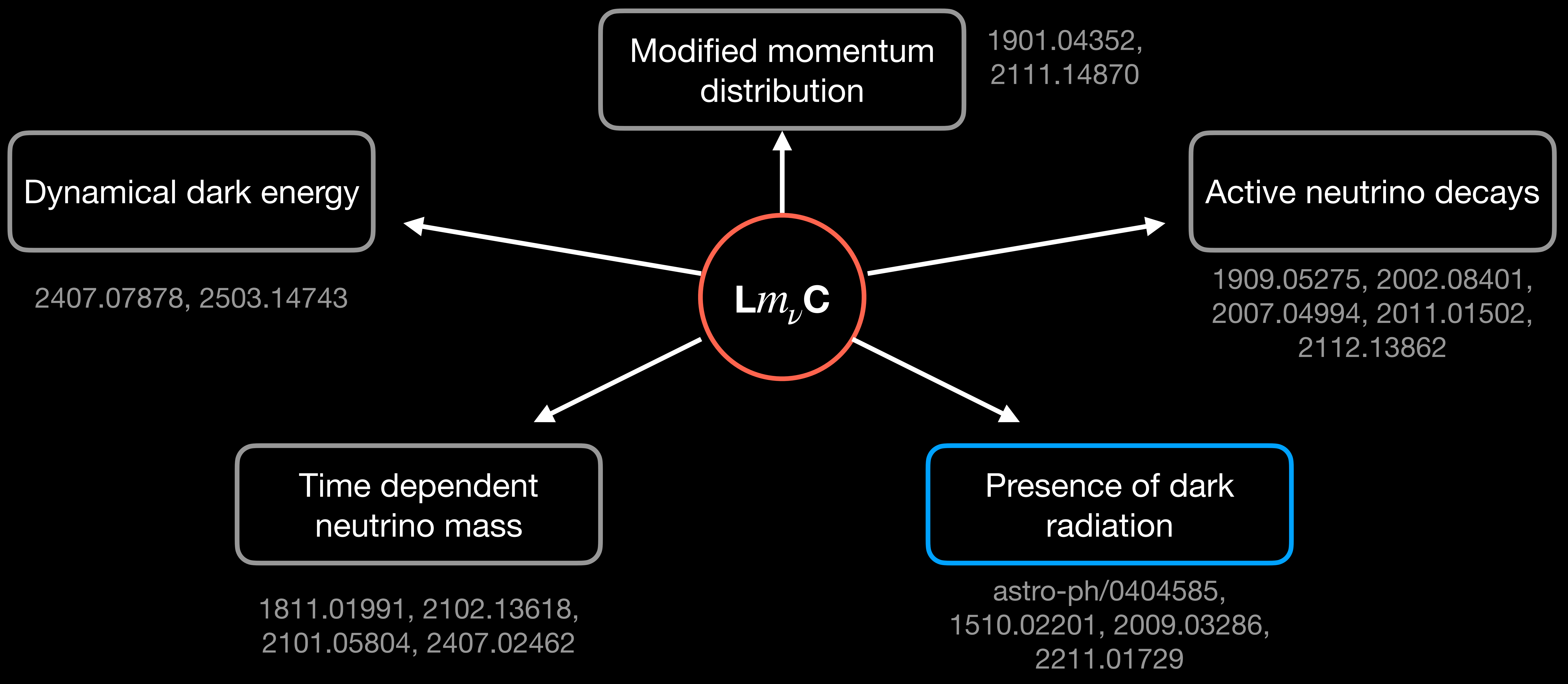
Standard cosmological scenario → We may not observe finite absolute neutrino mass in the laboratory

Can the two be reconciled? Can cosmological bounds be relaxed?

Relaxing the Cosmological ν mass bound

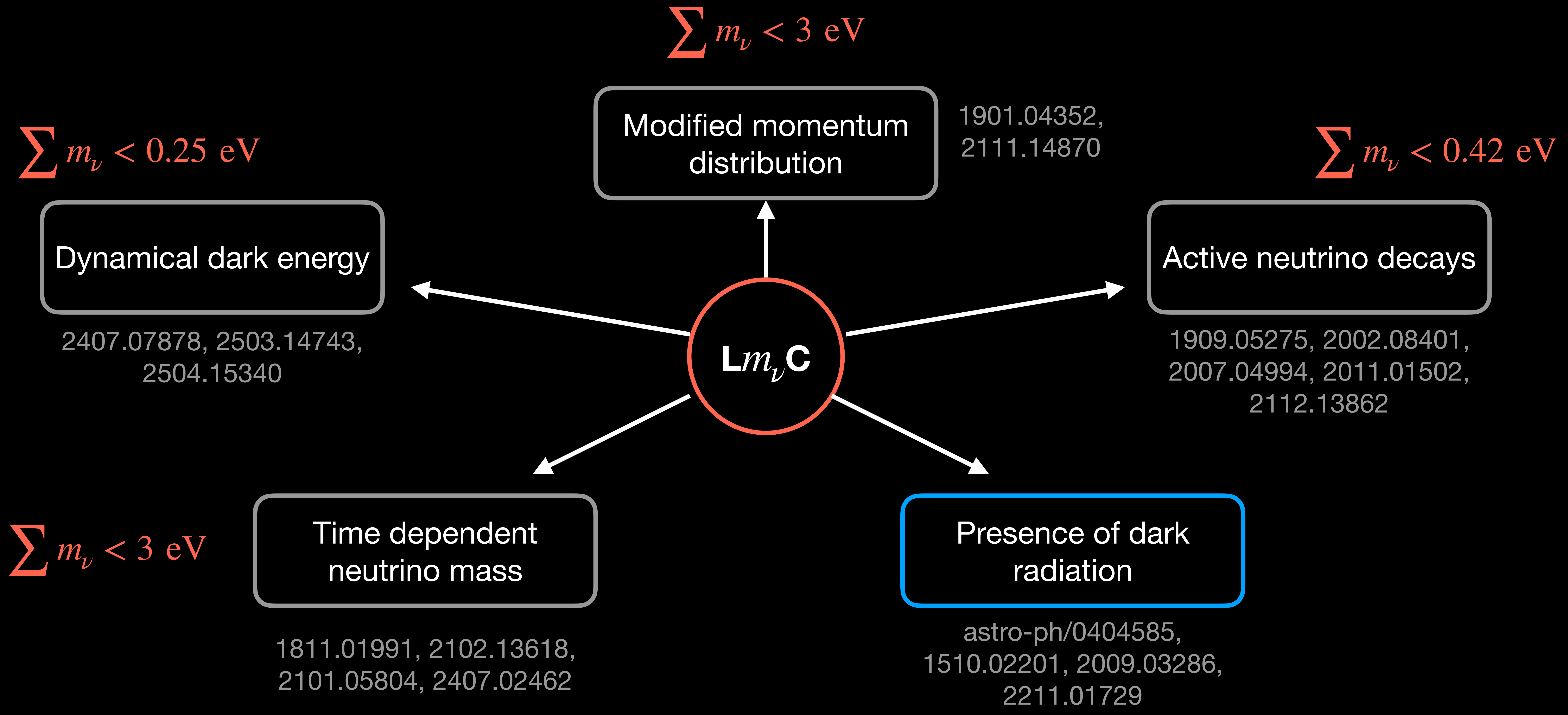
Large ν mass cosmology

Relaxing the cosmological bound requires non-standard scenarios → Large ν mass cosmologies



Relaxing the Cosmological ν mass bound

Large ν mass cosmology



Large m_ν Cosmology
Presence of dark radiation

Cosmological bounds are sensitive to neutrino energy density

$$\Omega_\nu h^2 \equiv \frac{\sum m_\nu n_\nu^0 h^2}{\rho_{\rm critical}} < 1.3 \times 10^{-3}\ (95\,\%\ \mathrm{CL}) \longrightarrow \sum m_\nu \times \left(\frac{n_\nu^0}{56\ \mathrm{cm}^{-3}}\right) < 0.12\ \mathrm{eV}\ (95\,\%\ \mathrm{CL})$$ PLANCK 2018

Reduce number density of neutrinos → Mass bound can be relaxed

At earlier times for ultra-relativistic νs: Energy density characterised by $N_{\rm eff} \propto \langle p_\nu \rangle n_\nu$
- 2.99 ± 0.17 PLANCK 2018
- $3.044(1)$ SM prediction

$$N_{\rm eff} \equiv \frac{8}{7}\left(\frac{11}{4}\right)^{4/3}\left(\frac{\rho_{\rm rad} - \rho_\gamma}{\rho_\gamma}\right)$$

Compensate decrease in n_ν : Add new light/massless d.o.f → Dark radiation

Large m_ν Cosmology
Presence of dark radiation

BBN ⟷ CMB

Introduce a mechanism to reduce n_ν

Add BSM states: X, χs that thermalise with νs

$100\ \mathrm{keV} \gtrsim T_\gamma \gtrsim 10\ \mathrm{eV}$

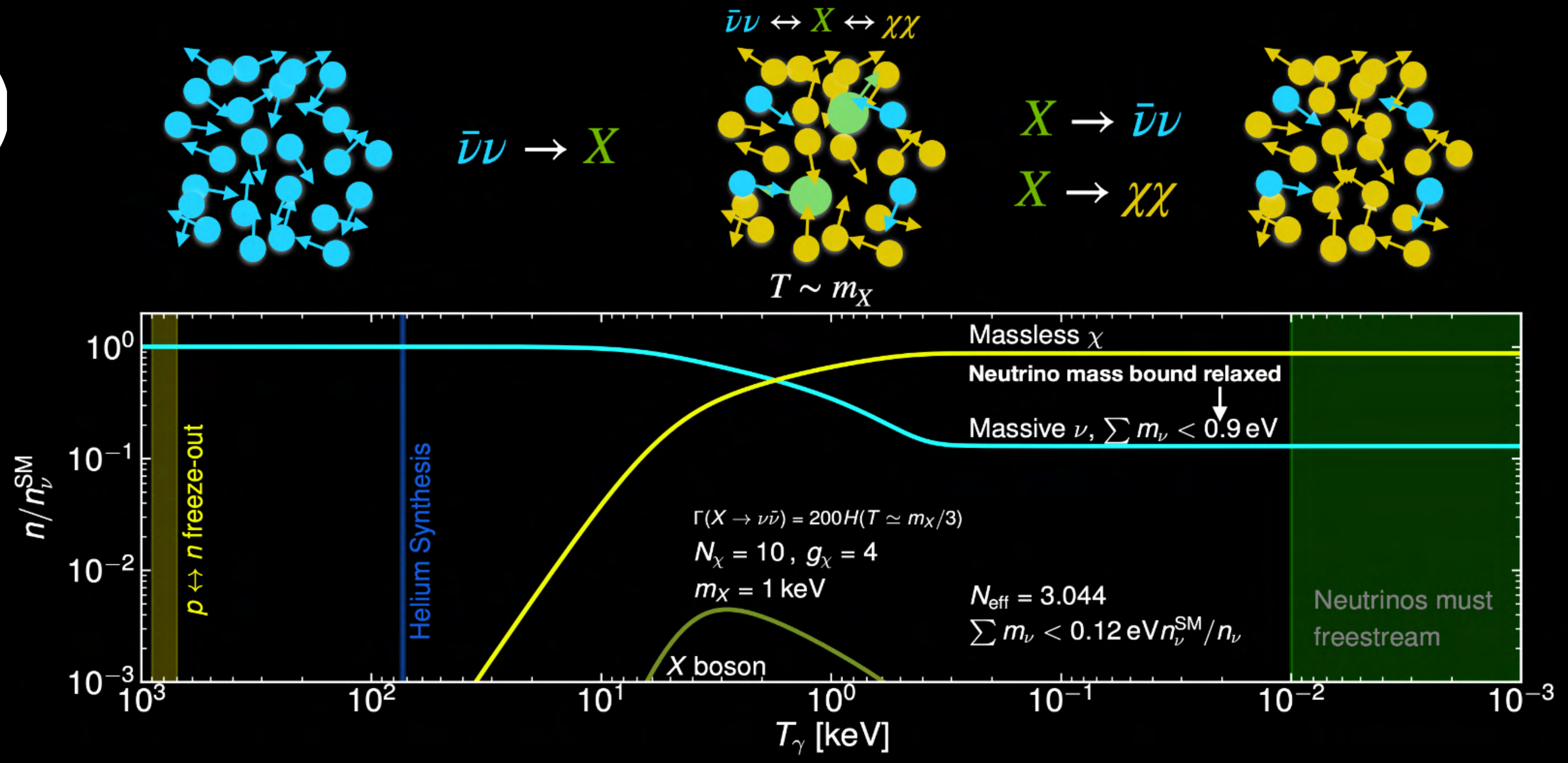
$N_{\rm eff}$ unchanged at earlier times

Post $\nu-$decoupling ($T_\gamma \sim 2\ \mathrm{MeV}$):

Neutrinos cannot be produced anymore:
Production of new states at their expense
→ n_ν reduced

$$\left[\sum m_\nu\right]_{\rm eff} = \sum m_\nu \frac{n_\nu}{n_\nu^{\rm SM}}$$

Depends on the dark d.o.f



Escudero, Schwetz, Terol-Calvo: 2211.01729

Large m_ν Cosmology

Presence of Dark Matter?

The dark sector can be enlarged to contain a light keV fermionic DM candidate along with the dark radiation → Multi-component DS

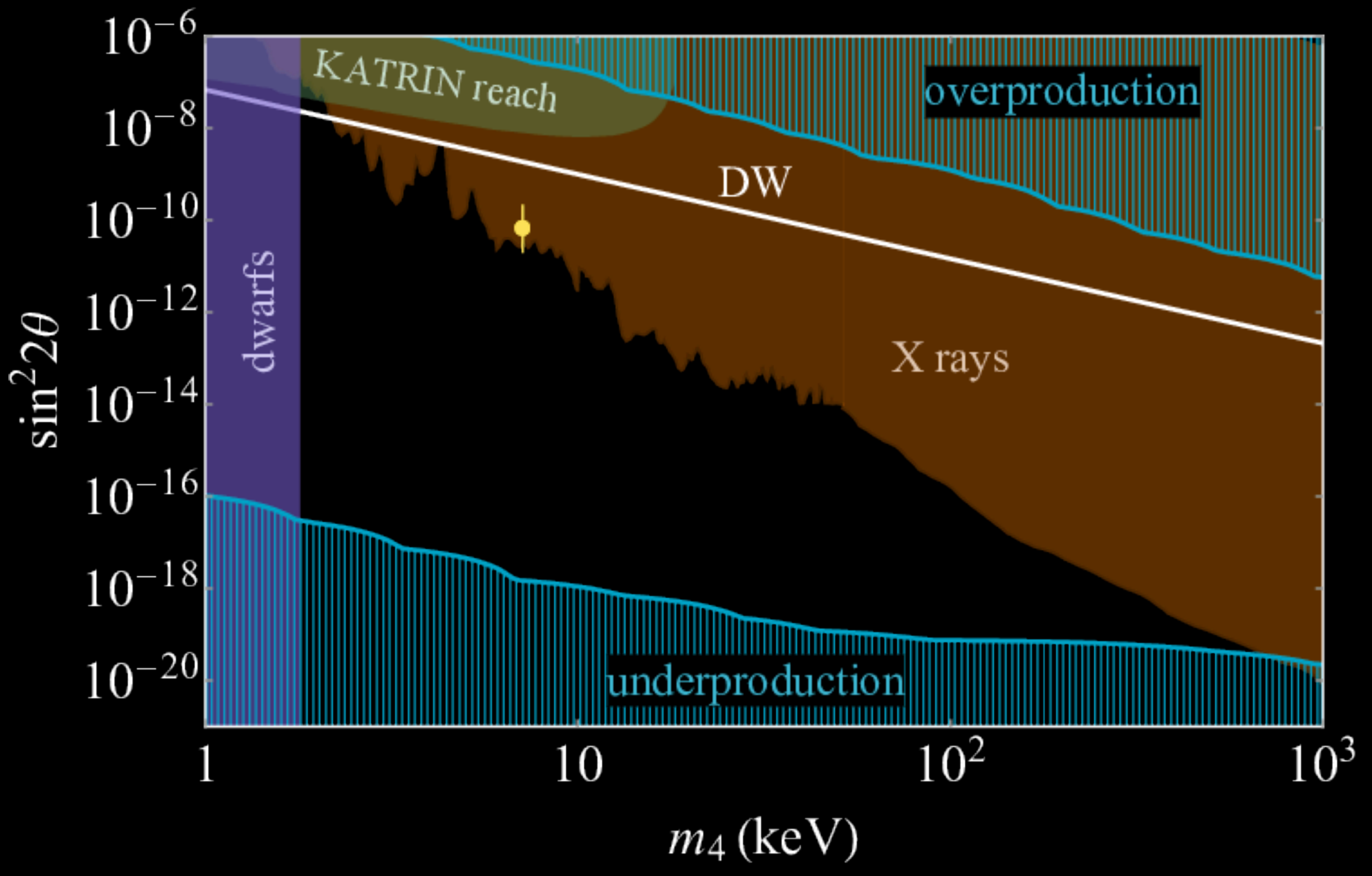
Problem keV scale sterile neutrino (ν_s) DM from oscillations: Dodelson Widrow is severely constrained

Solution Non-standard neutrino interactions open up parameter space

Thermal DM below MeV possible if DM comes into thermal equilibrium post ν-decoupling

Berlin, Blinov: 1706.07046

de Gouvêa, Sen, Tangarife, Zhang: 1901.04901

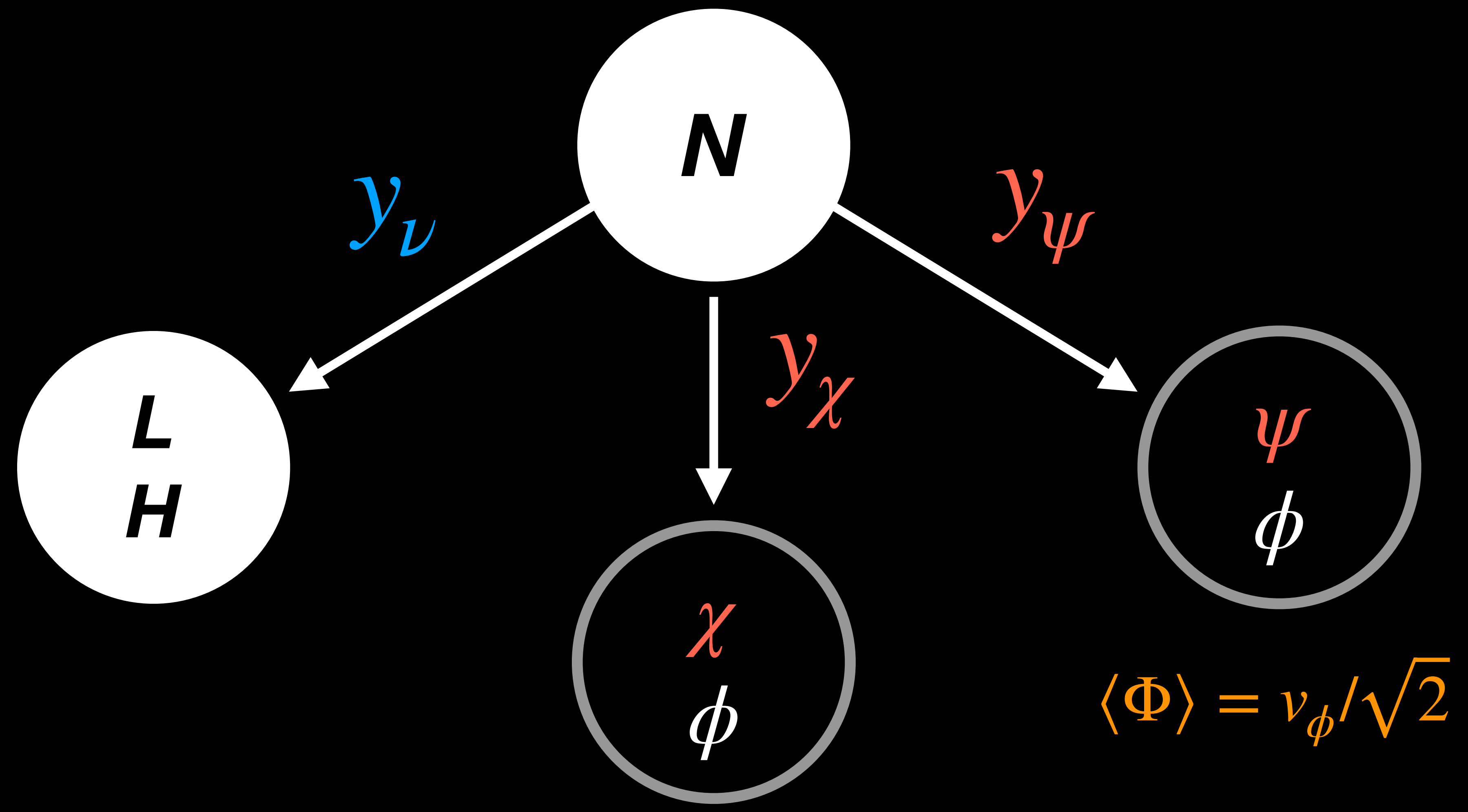


The Model

Minimally Extended Type-I Seesaw with U(1)$_X$

Similar to $\nu\Lambda$MDM: Ko and Tang: 1404.0236

	Field	Species	$U(1)_X$
Scalar	Φ	1	$+1$
Fermions	χ	N_X	-1
	ψ	1	-1
	N	3	0
	N'	1	0



+N′

$$-\mathscr{L}_{\text{new}} = Y_\nu \bar{N} l_L \tilde{H}^\dagger + Y_\chi \bar{N}\chi_L \Phi + Y_\psi \bar{N}\psi_L\Phi + Y'_\nu \bar{N}' l_L \tilde{H}^\dagger + Y'_\chi \bar{N}'\chi_L\Phi + Y'_\psi \bar{N}'\psi_L\Phi + \frac{1}{2}M\bar{N}N^c + \frac{1}{2}M'\bar{N}'N'^c + \text{H.c.}$$

$$V(H,\phi) = \mu_H^2 H^\dagger H + \lambda_H (H^\dagger H)^2 + \mu_\phi^2 |\Phi|^2 + \lambda_\phi |\Phi|^4 + \lambda_{H\phi}|\Phi|^2 H^\dagger H$$

Gauge interaction $\mathscr{L}_{\text{int}} = \sum_f Q_f g Z'_\mu \bar{f}\gamma^\mu f$ $\quad g \equiv m_{Z'}/v_\phi$

The Model

Masses & Mixings

Neutral fermion mixing matrix:

$(\chi_L^c, \nu_L^c, \psi_L^c, N', N)$ basis

$$\mathcal{M}_n = \begin{pmatrix} 0 & 0 & 0 & \Lambda' & \Lambda \\ 0 & 0 & 0 & m_D{}' & m_D \\ 0 & 0 & 0 & \kappa' & \kappa \\ \Lambda'^T & m_D^{T'} & \kappa'^T & M' & 0 \\ \Lambda^T & m_D^T & \kappa^T & 0 & M \end{pmatrix}$$

$Y_\chi v_\phi/\sqrt{2}$ $Y_\nu v_{\rm EW}/\sqrt{2}$ $Y'_\psi v_\phi/\sqrt{2}$

$M \gg M' \gg m_D \gg \kappa', \Lambda \gg m'_D, \Lambda', \kappa$

Massive: $2N_{\rm heavy}$

Massless: $(3 + N_{\rm light} - N_{\rm heavy})$

$m_N \approx M\,, m_{N'} \approx M'$

$m_\chi = 0$ N_χ massless fermions

$m_\nu \approx m_D M^{-1} m_D^T$

$m_\psi \approx \kappa' M'^{-1} \kappa'^T$

Seesaw induced Majorana masses

$$m_\psi \sim 10\ {\rm keV} \left(\frac{\kappa'}{10^4\ {\rm keV}}\right)^2 \left(\frac{10\ {\rm GeV}}{M'}\right)$$

Small mixing $\theta_{\nu\chi} = \frac{\Lambda}{m_D}$

$\theta_{\nu\psi} = \frac{m_D{}'}{\kappa'}$ Suppressed mixing

ψ becomes the DM candidate in the model

DM freeze-out in DS

Production & Depletion

$$\nu\nu \leftrightarrow Z' \leftrightarrow \chi\chi$$

DM can be produced by $Z' \leftrightarrow \psi\psi \ (m_{Z'} > 2m_\psi)$ or $Z'Z' \leftrightarrow \psi\psi$ and $\chi\chi \leftrightarrow \psi\psi \ (m_\psi > m_{Z'})$

Annihilations $\psi\psi \to \chi\chi$ and $\psi\psi \to Z'Z'$ freeze-out at $T_{\rm dark} < m_\psi$

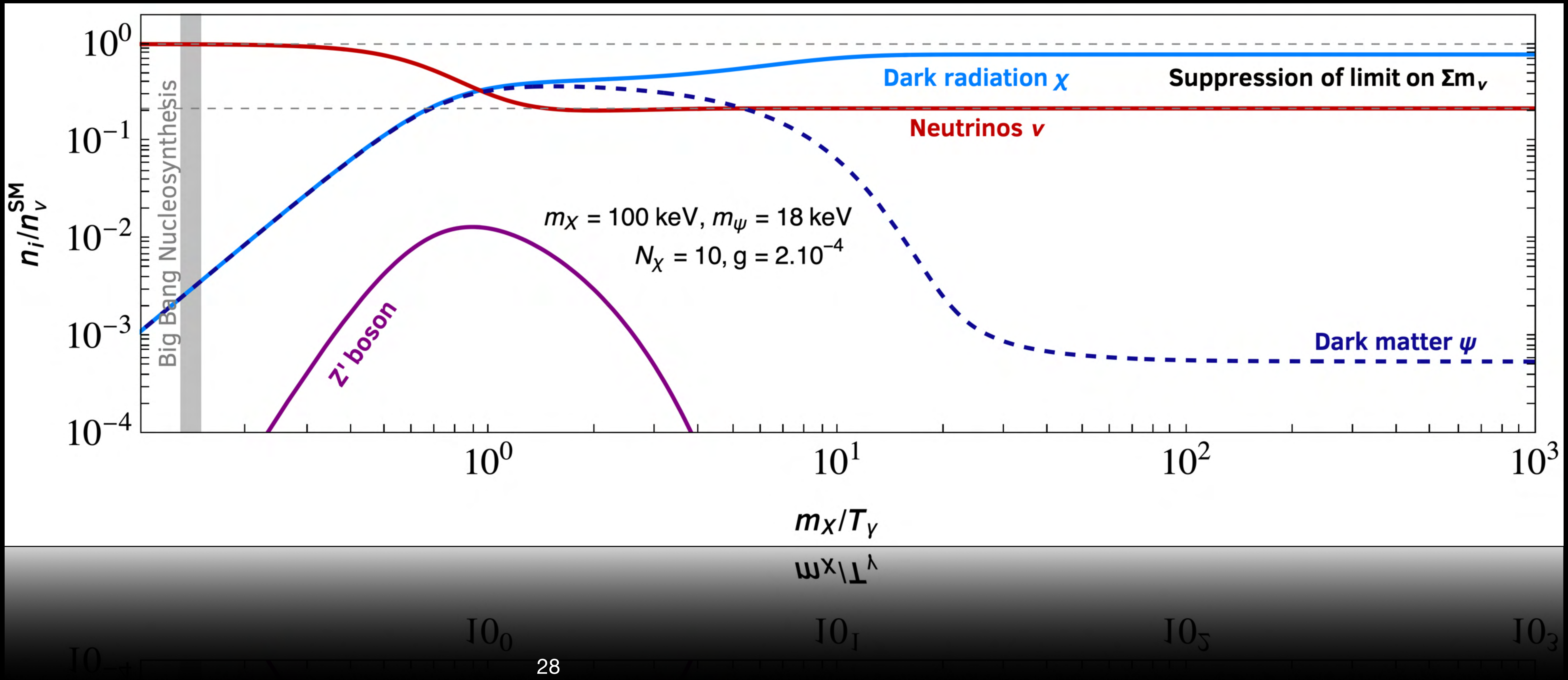
$$\mathscr{L}_{\rm int} = \sum_f Q_f g Z'_\mu \bar{f}\gamma^\mu f$$

$$\lambda_{Z'}^{\nu\nu} \simeq \frac{m_{Z'}}{v_\phi}\theta_{\nu\chi}^2$$

$$\lambda_{Z'}^{\nu\chi} = \frac{m_{Z'}}{v_\phi}\theta_{\nu\chi}$$

$$\lambda_{Z'}^{\psi\psi} = \lambda_{Z'}^{\chi\chi} = \frac{m_{Z'}}{v_\phi}$$

$$\lambda_{Z'}^{\nu\psi} = \frac{m_{Z'}}{v_\phi}\theta_{\nu\psi}$$



DM freeze-out in DS

Relic abundance

ψ comes into thermal equilibrium with the DS and finally freezes out

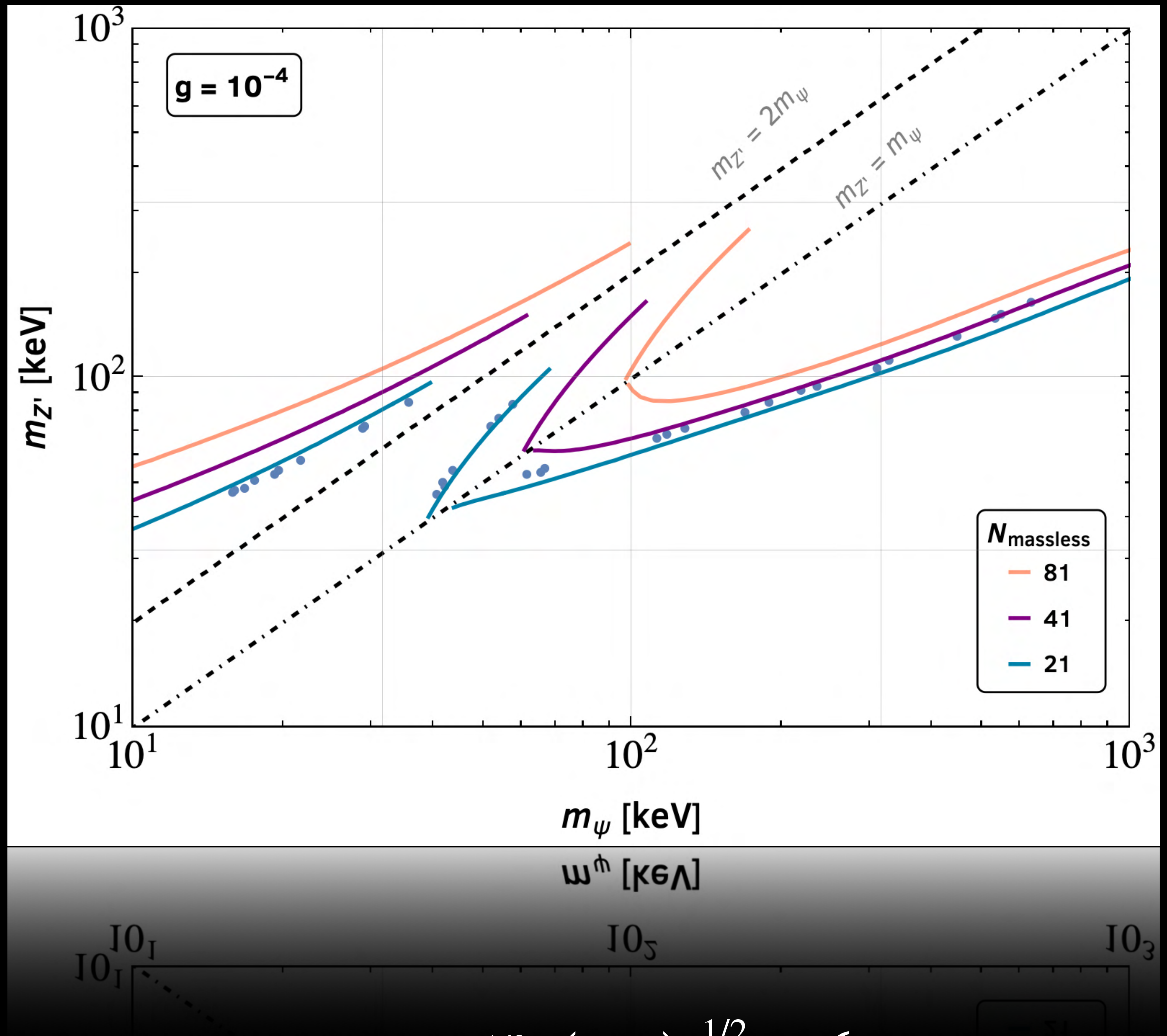
$$\Omega_\psi h^2 \simeq x_f \frac{10^{-10}\,\mathrm{GeV}^{-2}}{\langle \sigma v \rangle_{\mathrm{tot}}}$$

Depends on DS temperature

$$(\sigma v)_{\psi\psi\to\chi\chi} \approx \tilde{N} \frac{g^4}{48\pi} \frac{m_\psi^2}{(m_{Z'}^2 - 4m_\psi^2)^2} v^2$$

$$(\sigma v)_{\psi\psi\to Z'Z'} \approx \frac{g^4}{16\pi\, m_\psi^2} \left(1 - \frac{m_{Z'}^2}{m_\psi^2}\right)^{1/2} \left(1 + \frac{m_\psi^4}{m_{Z'}^4} v^2\right)$$

Extreme limits: $m_{Z'} \gg m_\psi$ or vice versa $\to (\sigma v) \propto v_\phi^{-4}$



$$v_\phi \simeq 10^5\,\mathrm{keV} \left(\frac{m_\psi}{15\,\mathrm{keV}}\right)^{1/2} \left(\frac{3.2}{x_f}\right)^{1/2} \times \begin{cases} 2\tilde{N}^{1/4} & (m_{Z'} \gg m_\psi) \\ 2.4 & (m_{Z'} \ll m_\psi) \end{cases}$$

Constraints

Stability & X-ray bounds

DM decay → ψ lifetime should be larger than the age of the universe

$m_\psi < m_{Z'}$

$\psi \to \nu\chi\chi$

$$\theta^2_{\nu\psi} < 2\times 10^{-16}\left(\frac{15\ \mathrm{keV}}{m_\psi}\right)^5\left(\frac{21}{\tilde{N}}\right)\left(\frac{v_\phi}{2\ \mathrm{GeV}}\right)^4$$

$m_\psi > m_{Z'}$

$\psi \to Z'\nu$

$$\theta^2_{\nu\psi} < 1.2\times 10^{-30}\left(\frac{m_{Z'}}{10\ \mathrm{keV}}\right)^2\left(\frac{10^{-4}}{g}\right)^2\left(\frac{40\ \mathrm{keV}}{m_\psi}\right)^3$$

Sterile νDM mixing with active νs → Observable monochromatic X-ray line

$\psi \to \nu\gamma$

$$\Gamma_{\psi\to\nu\gamma} = \frac{9\,\alpha\,G_F^2}{256\cdot 4\pi^4}\sin^2(2\theta_{\nu\psi})\,m_\psi^5$$

$$\theta^2_{\nu\psi} \lesssim 7.65\times 10^{-13}\left(\frac{15\ \mathrm{keV}}{m_\psi}\right)^5$$

$\theta_{\nu\psi}$ should be really suppressed

Constraints

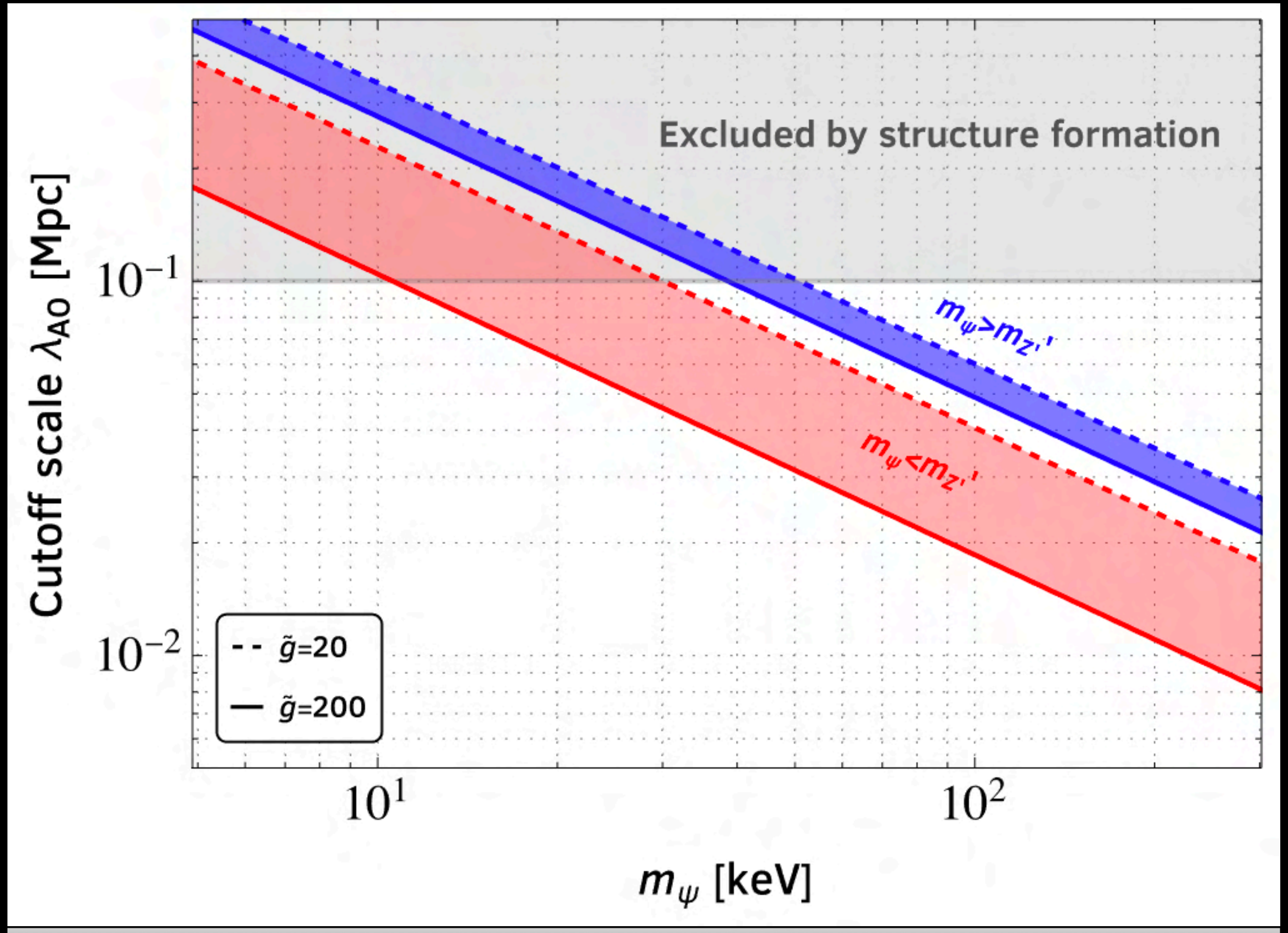
Structure formation

Potentially large free-streaming scale →
Prevent formation of small scale structures

Post freeze-out DM ψ remains in thermal contact with dark radiation χ via elastic processes $\psi\chi \leftrightarrow \psi\chi$

$$M_{\rm hm} = \frac{4\pi}{3}\rho_{\rm DM}\left(\frac{\lambda_{\rm hm}}{2}\right)^3 \approx 1.9\times 10^7\, M_\odot \left(\frac{\lambda_{\rm hm}}{0.1\,{\rm Mpc}}\right)^3$$

Depends on temperature of kinetic decoupling $T_{\rm kd}$



Viable Parameter Space

Putting everything together

Taule, Escudero, Garny: 2207.04062

Thermalisation

νs should thermalise with Z' in $0.7\ \mathrm{MeV} > T_\gamma > 10\ \mathrm{eV}$

BBN constraints

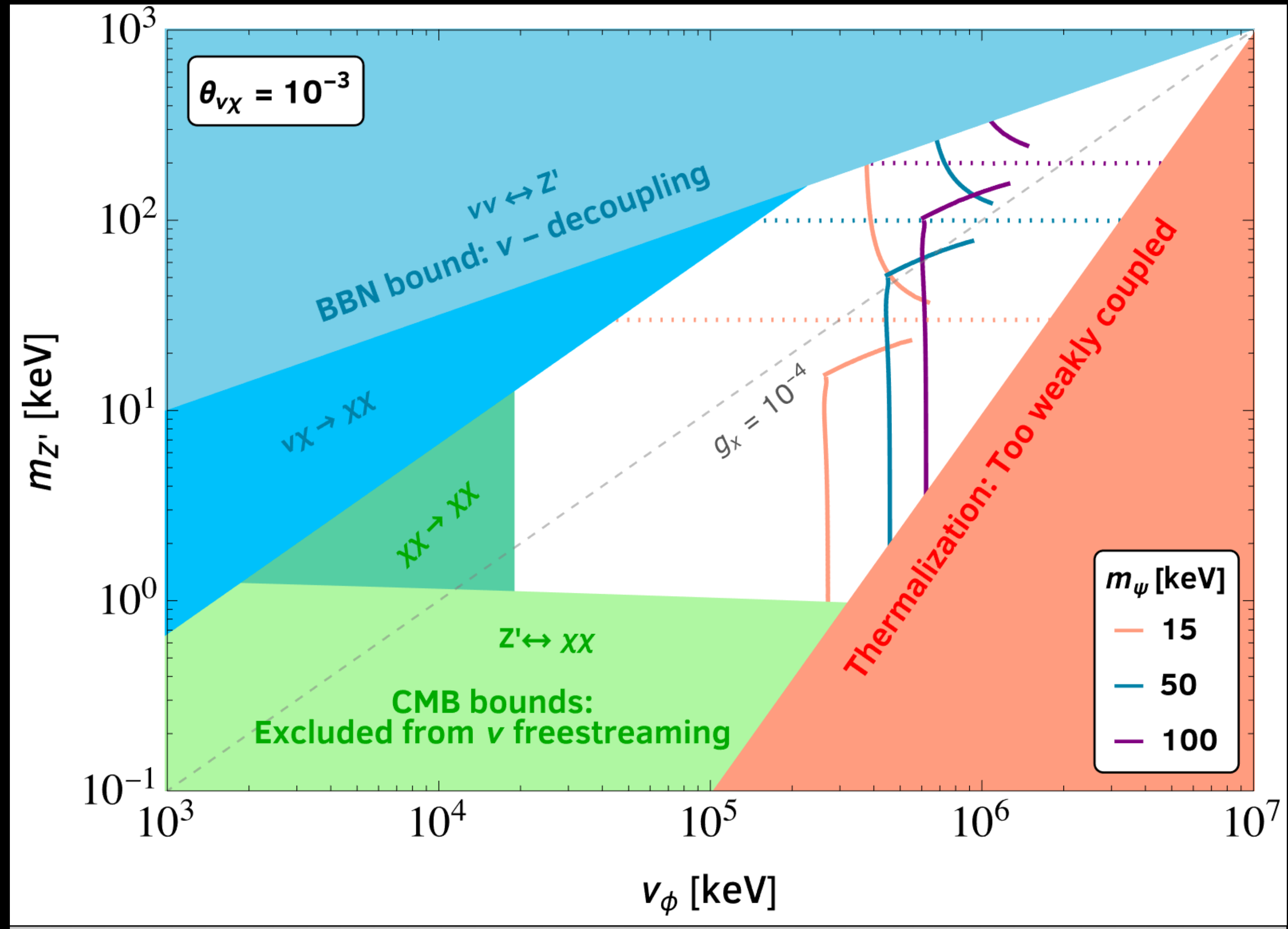
νs should not thermalise with Z'; avoid χ's exponential growth $\nu\chi \leftrightarrow \chi\chi$ at $T_\gamma > 0.7\ \mathrm{MeV}$

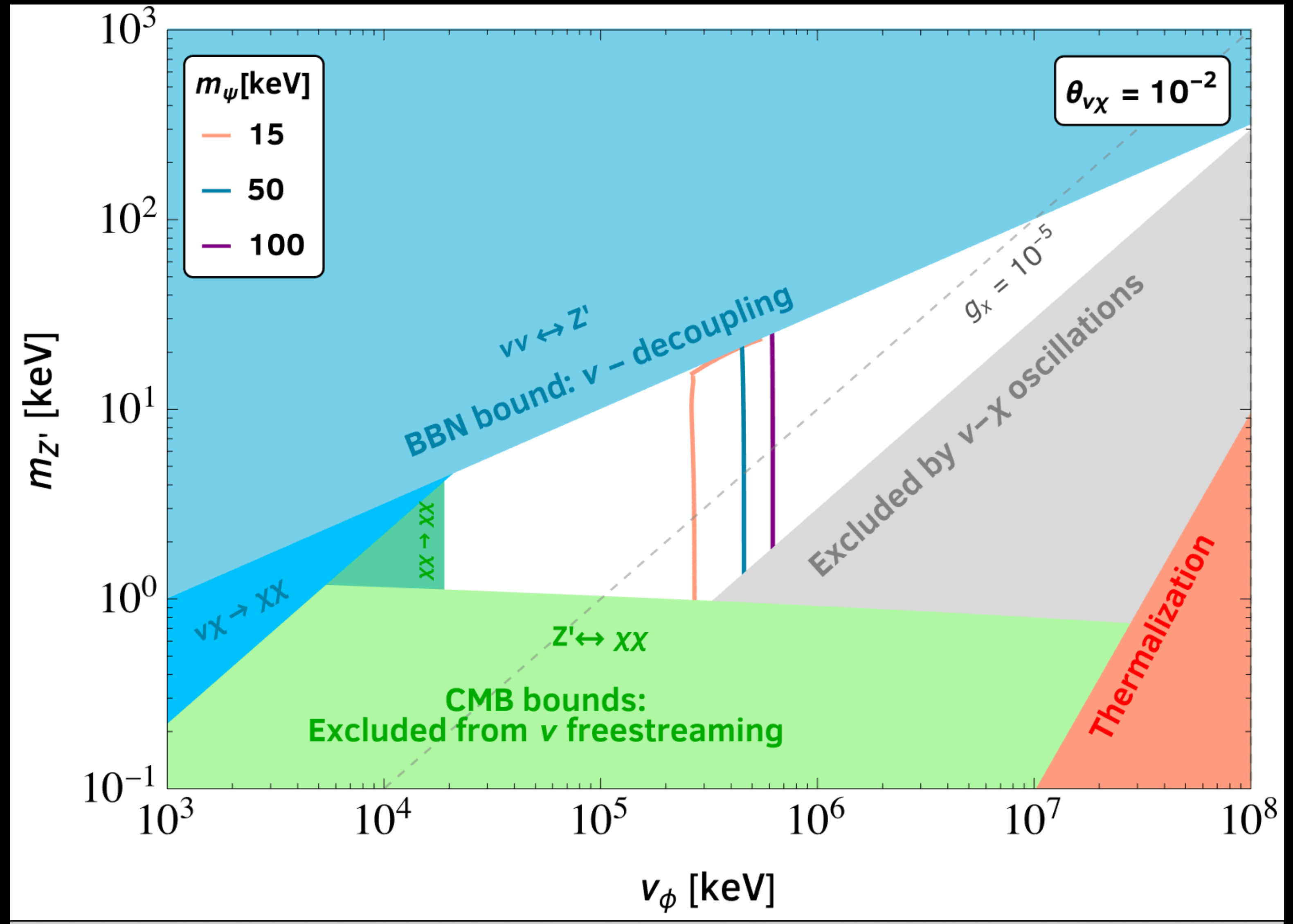
CMB constraints

$\nu\nu \to Z'$ and $Z' \to \chi\chi$ must be inefficient at $z \sim 10^5$; CMB not perturbed by lack of χ free streaming

Active-sterile mixing

Constrain production of χ from oscillations before BBN using $\Delta N_{\mathrm{eff}} < 0.3 \ \to 10^{-4} \leq \theta_{\nu\chi} < 10^{-1}$





Neutrino Mass Suppression

N_{eff} and DS Temperature

New degrees of freedom come into equilibrium with neutrinos at T_ν^{eq} to form a system with T_{eq}

$$\rho_\nu(T_\nu^{\text{eq}}) = \sum_{f=\nu,\chi,\psi} \rho_f(T_{\text{eq}}) + \rho_{Z'}(T_{\text{eq}})$$

System evolves adiabatically from T_{eq} to T_{fin} when ψ, Z' become non-relativistic, use $a_{\text{eq}}^3 s_{\text{eq}}(T_{\text{eq}}) = a_{\text{fin}}^3 s_{\text{fin}}(T_{\text{fin}})$

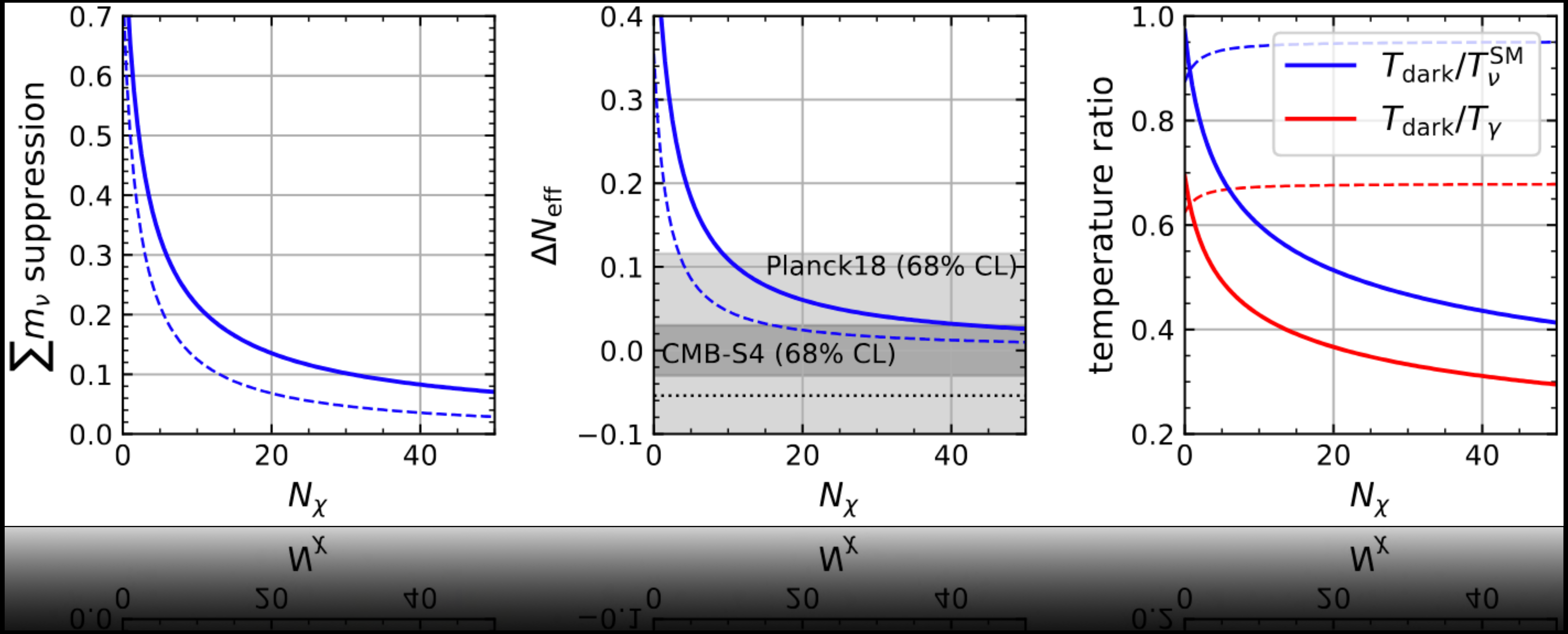
$$\frac{n_\nu}{n_\nu^{\text{SM}}} = \left(\frac{T_{\text{dark}}}{T_\nu^{\text{SM}}}\right)^3 = \frac{g_\nu + \tilde{g} + g_\psi + \frac{8}{7} g_{Z'}}{g_\nu + \tilde{g}} \left(\frac{g_\nu}{g_\nu + \tilde{g} + g_\psi + \frac{8}{7} g_{Z'}}\right)^{3/4}$$

$$N_{\text{eff}} \equiv \frac{8}{7}\left(\frac{11}{4}\right)^{4/3} \frac{\rho_{\text{dark}}}{\rho_\gamma} = \frac{g_\nu + \tilde{g}}{2}\left(\frac{T_{\text{dark}}}{T_\nu^{\text{SM}}}\right)^4$$

$$\left[\sum m_\nu\right]_{\text{eff}} = \sum m_\nu \frac{n_\nu}{n_\nu^{\text{SM}}}$$

Neutrino Mass Suppression

N_{eff} and DS Temperature



M [GeV]	M' [GeV]	m_D [GeV]	κ' [GeV]	Λ [GeV]	v_ϕ [GeV]	m_ψ [keV]	$m_{Z'}$ [keV]	$g = m_{Z'}/v_\phi$	$\theta_{\nu\chi}$	N_χ	n_ν/n_ν^{SM}	ΔN_{eff}
10^{11}	10^2	4.47	0.043	0.004	0.5	18.5	100	2×10^{-4}	10^{-3}	10	0.216	0.109
10^{12}	10^3	14.14	0.23	0.141	0.8	53	77	9.6×10^{-5}	10^{-2}	10	0.216	0.109
10^{13}	10^2	44.7	0.1	0.044	0.6	100	32	5.3×10^{-5}	10^{-3}	20	0.135	0.060

Summary

It's a wrap

Comparing cosmology and laboratory bounds on $\sum m_\nu$ → Hints of new physics

The cosmological neutrino mass bound can be relaxed with a light dark sector

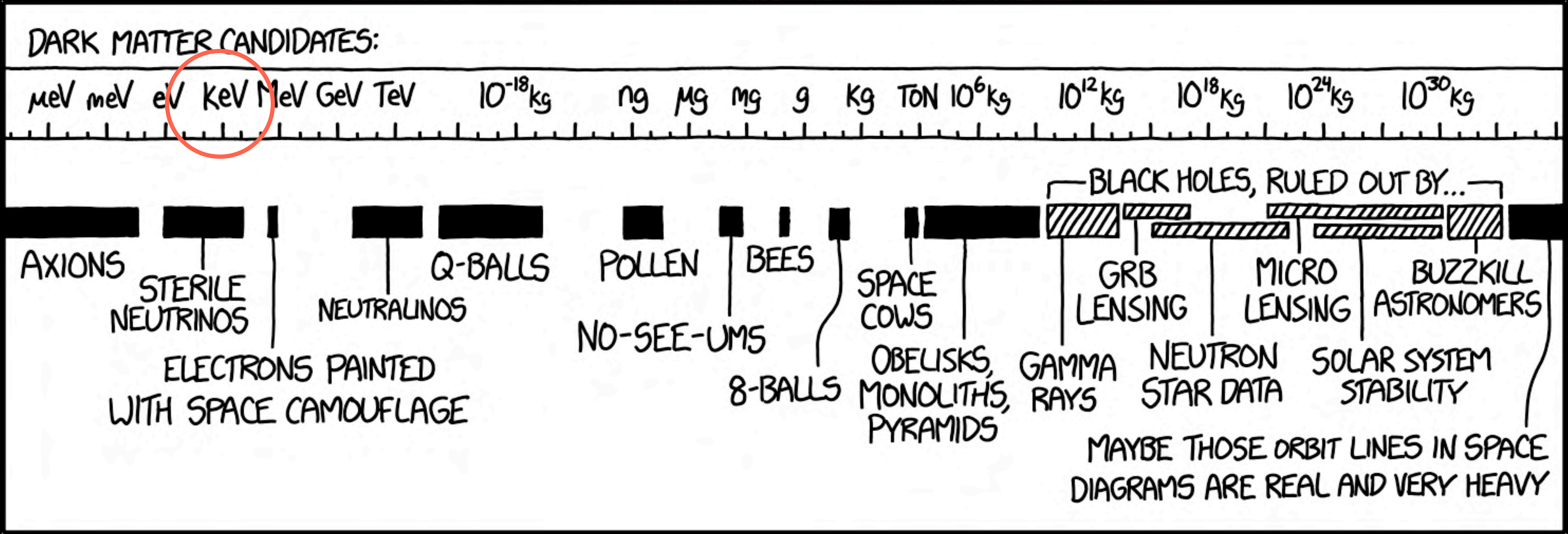
We embed the mechanism in an extended seesaw model → ν masses, DM and leptogenesis

Model has a complex particle content though not so numerous free parameters → $m_\psi, m_{Z'}, v_\phi, \theta_{\nu\chi}, N_\chi$

Dark sector particles play no role above $T_\gamma \sim 1$ MeV and come into equilibrium after ν-decoupling

DM thermalises with the DS and then freezes out → Abundance set by DS gauge interactions, not by mixing with SM neutrinos

Signatures of the model → Slightly increased $N_{\rm eff}$ at late times, Suppressed matter power spectrum at small scales (warm DM) xkcd



Thank You

Backup

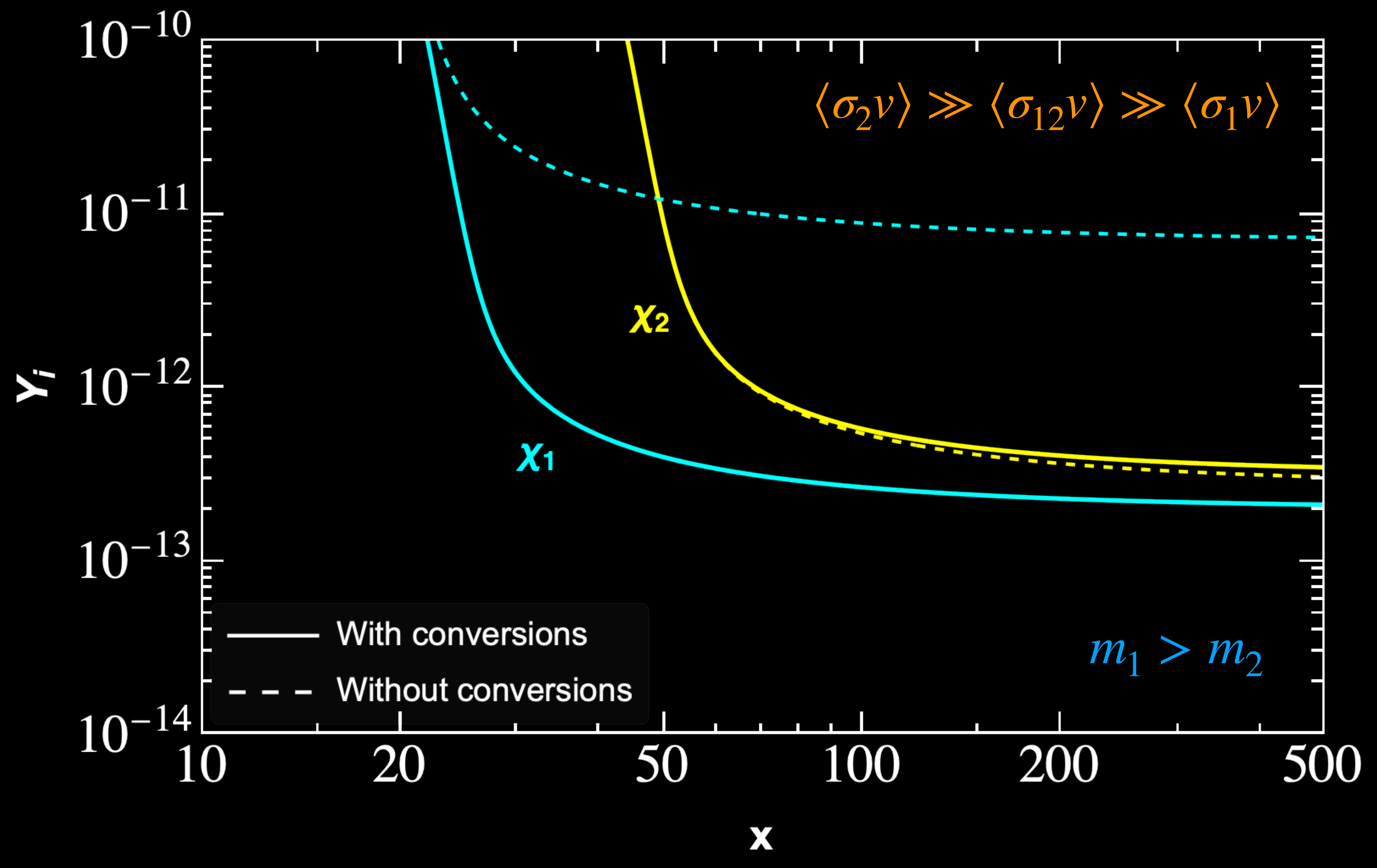
Symmetric Components

2DM

$$\frac{dY_1}{dx} = -\frac{s}{Hx}\left[\langle\sigma_1 v\rangle\,(Y_1^2 - \bar{Y}_1^2) + \langle\sigma_{12} v\rangle(Y_1^2 - \frac{Y_2^2}{\bar{Y}_2^2}\bar{Y}_1^2)\right] \qquad \frac{dY_2}{dx} = -\frac{s}{Hx}\left[\langle\sigma_2 v\rangle\,(Y_2^2 - \bar{Y}_2^2) - \langle\sigma_{12} v\rangle(Y_1^2 - \frac{Y_2^2}{\bar{Y}_2^2}\bar{Y}_1^2)\right]$$

No conversions: DM abundance dominated by the particle with *smallest* $\langle\sigma v\rangle$

$$\Omega h^2 \propto \frac{1}{\langle\sigma_1 v\rangle} + \frac{1}{\langle\sigma_2 v\rangle} \equiv \frac{1}{\langle\sigma v\rangle_{\rm eff}} \simeq \frac{1}{2.2\cdot 10^{-26}\,{\rm cm^3/s}}$$



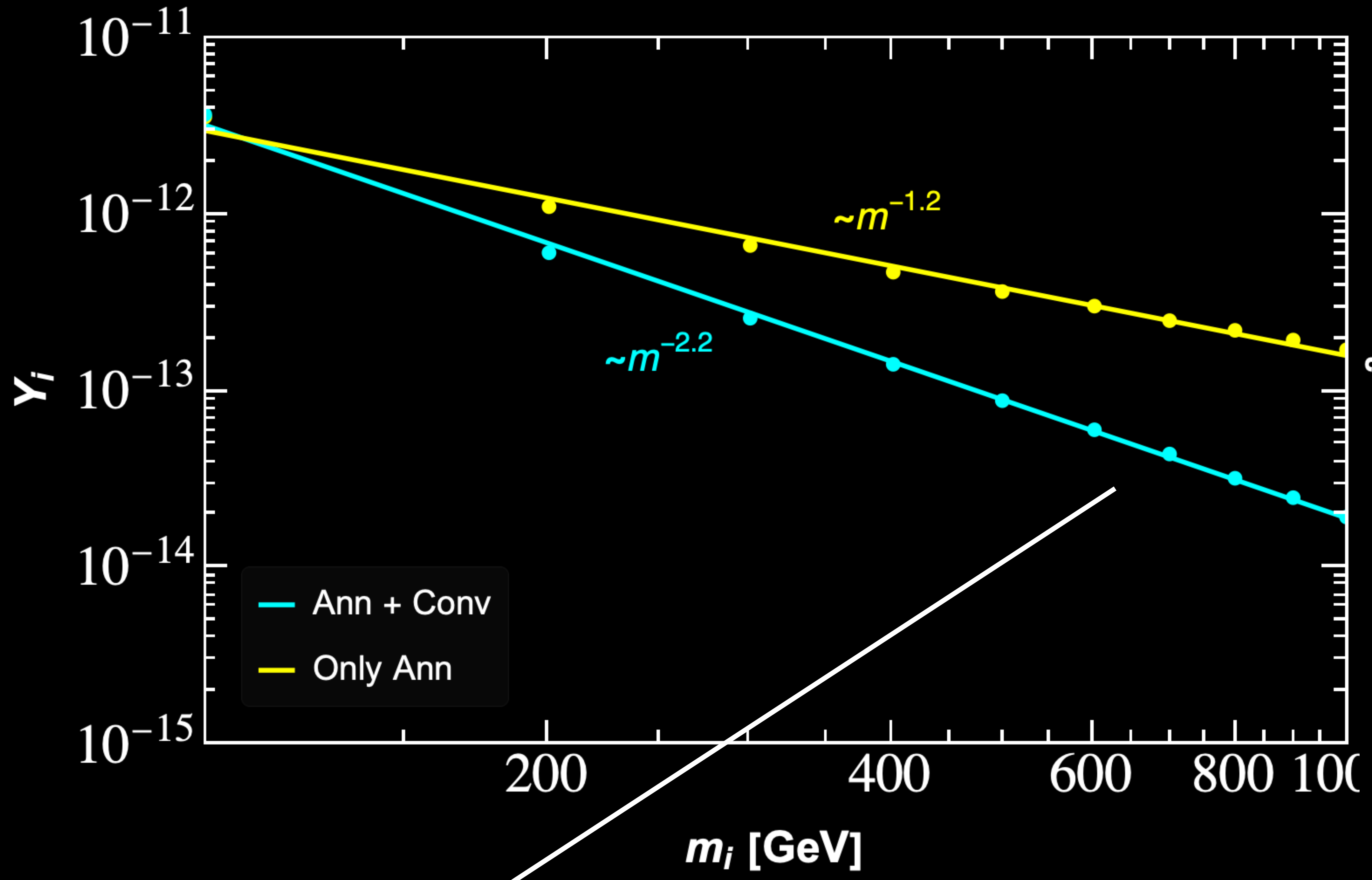
With conversions: heavier components have *reduced* abundance

$$\Omega h^2 \propto \frac{1}{\langle\sigma_1 v\rangle + \langle\sigma_{12} v\rangle} + \frac{1}{\langle\sigma_2 v\rangle} \equiv \frac{1}{\langle\sigma v\rangle_{\rm eff}}$$

Aoki et al: 1207.3318
Bhattacharya et al: 1607.08461

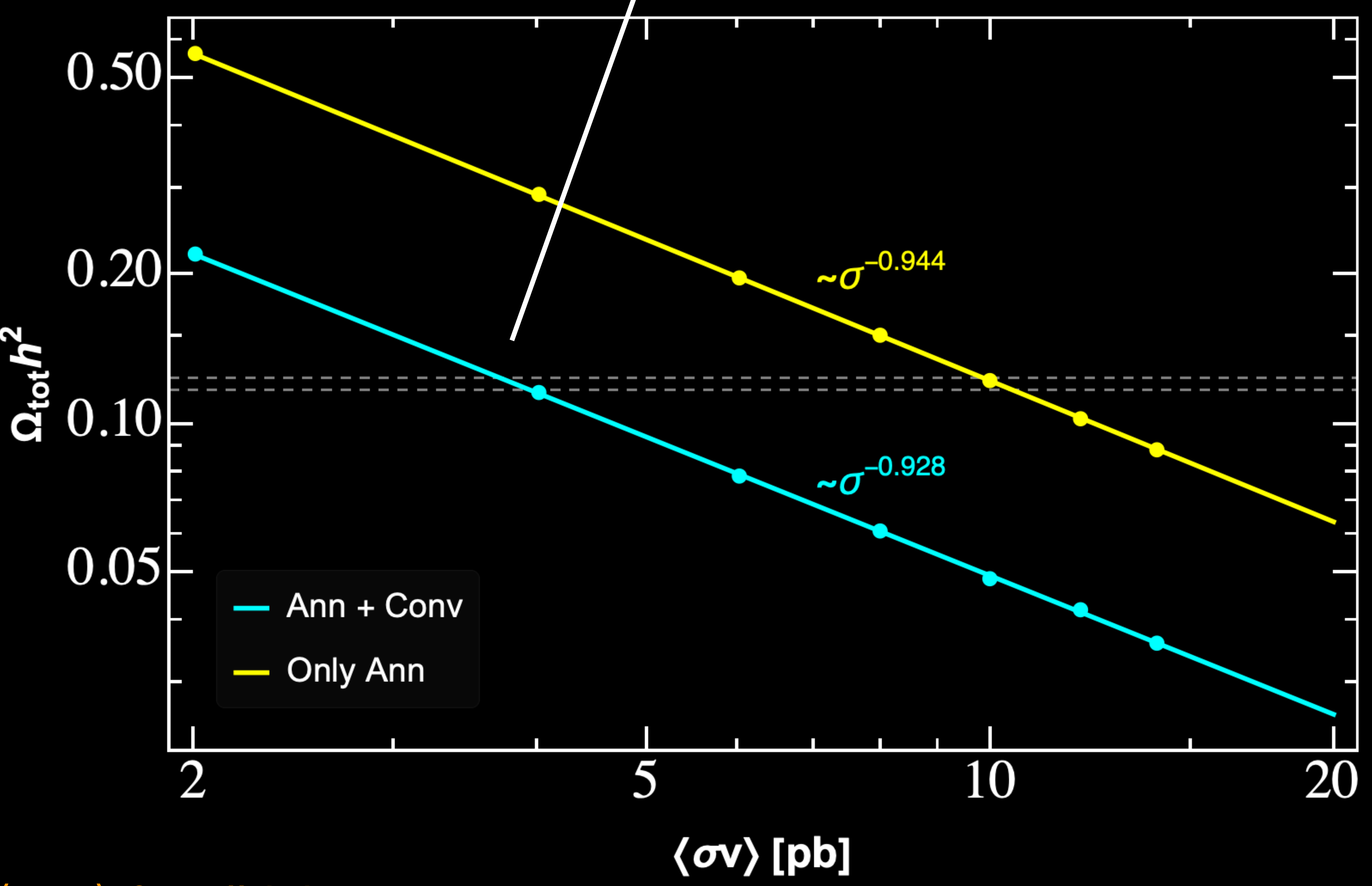
Symmetric Components

10DM



Heavier component yields suppressed due to conversions

Smaller annihilation cross-sections with conversions: relaxes limits!



$\langle\sigma_i v\rangle = \langle\sigma_{ij} v\rangle$ for all i, j

Model

Masses, Mixing and Parameters

$$m_\chi = 0\,,$$

$$m_\nu = \frac{(m_D\kappa' - m_D{}'\kappa)^2 + (m_D{}'\Lambda - m_D\Lambda')^2 + (\kappa'\Lambda - \kappa\Lambda')^2}{M'(m_D^2 + \kappa^2 + \Lambda^2) + M(m_D{}'^2 + \kappa'^2 + \Lambda'^2)}\,,$$

$$m_\psi \approx \frac{m_D^2 + \kappa^2 + \Lambda^2}{M} + \frac{m_D{}'^2 + \kappa'^2 + \Lambda'^2}{M'}\,,$$

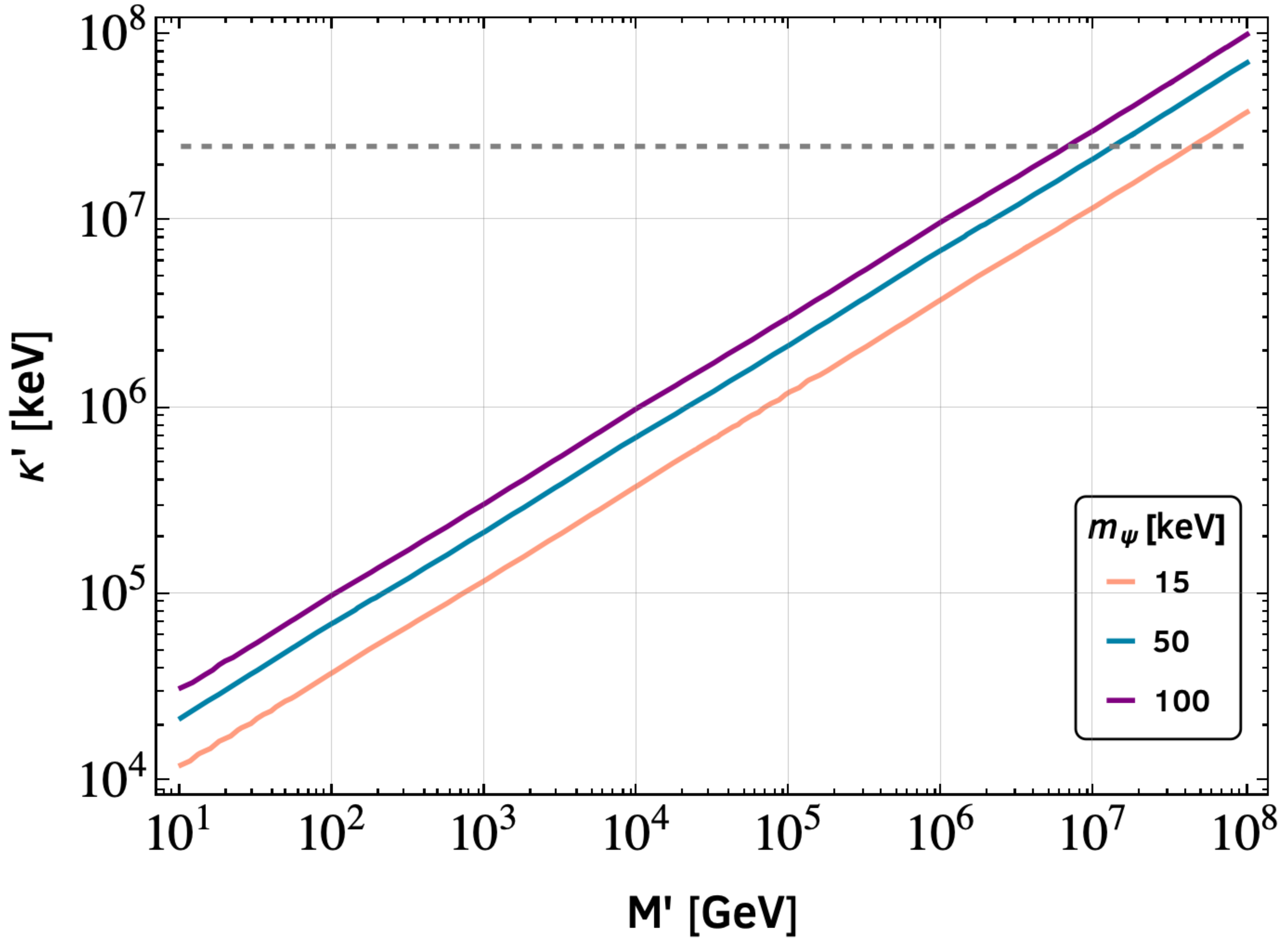
$$m_{N'} \approx M'\,,$$

$$m_N \approx M\,.$$

$$\theta_{\nu N} = \frac{m_D}{M}\,, \quad \theta_{\nu\chi} = \frac{\Lambda}{m_D}\,, \quad \theta_{N\chi} = \frac{\Lambda}{M} \sim 0\,,$$

$$\theta_{\nu N'} = \frac{m_D{}'}{M'}\,, \quad \theta_{\nu\psi} = \frac{m_D{}'}{\kappa'}\,, \quad \theta_{N'\psi} = \frac{\kappa'}{M'}\,.$$

$$m_\psi \sim 10\ \text{keV} \left(\frac{\kappa'}{10^4\ \text{keV}}\right)^2 \left(\frac{10\ \text{GeV}}{M'}\right)$$



DS Freeze-out

Analytical Solution

$$\Omega_\psi h^2 \simeq x_f \frac{10^{-10}\,\mathrm{GeV}^{-2}}{\langle \sigma v \rangle_{\rm tot}}$$

$x_f = m_\psi / T_{\gamma,f}$

DS Freeze-out temperature $\quad x_f \approx \xi \ln[0.04\delta(\delta+2)(g_\psi/g_{\rm eff}^{1/2})\beta\xi^{3/2}] - \frac{1}{2}\xi \ln\{\xi \ln[0.04\delta(\delta+2)(g_\psi/g_{\rm eff}^{1/2})\beta\xi^{3/2}]\}$

$\xi \equiv T_{\rm dark}/T_\gamma$

$g_{\rm eff} = g_\gamma + g_{\rm dark}\xi^4$

$\beta = M_{\rm pl}\, m_\psi \langle \sigma v \rangle_{\rm tot}^{x=1}$

DS Freeze-out

Numerical Solution

$$\frac{dY_\nu}{dx} = \frac{\langle \Gamma_\nu \rangle}{Hx} \left(Y_{Z'} - Y_{Z'}^{\rm eq} \frac{Y_\nu^2}{Y_\nu^{\rm eq\,2}} \right) ,$$

$$\frac{dY_{Z'}}{dx} = \sum_{i=\nu,\chi,\psi} -\frac{\langle \Gamma_i \rangle}{Hx} \left(Y_{Z'} - Y_{Z'}^{\rm eq} \frac{Y_i^2}{Y_i^{\rm eq\,2}} \right) + \frac{s\langle \sigma v \rangle_{\psi\psi\to Z'Z'}}{Hx} \left(Y_\psi^2 - \frac{Y_\psi^{\rm eq\,2}}{Y_{Z'}^{\rm eq\,2}} Y_{Z'}^2 \right) ,$$

$$\frac{dY_\chi}{dx} = \frac{\langle \Gamma_\chi \rangle}{Hx} \left(Y_{Z'} - Y_{Z'}^{\rm eq} \frac{Y_\chi^2}{Y_\chi^{\rm eq\,2}} \right) + \frac{s\langle \sigma v \rangle_{\psi\psi\to \chi\chi}}{Hx} \left(Y_\psi^2 - \frac{Y_\psi^{\rm eq\,2}}{Y_\chi^{\rm eq\,2}} Y_\chi^2 \right) ,$$

$$\frac{dY_\psi}{dx} = \frac{\langle \Gamma_\psi \rangle}{Hx} \left(Y_{Z'} - Y_{Z'}^{\rm eq} \frac{Y_\psi^2}{Y_\psi^{\rm eq\,2}} \right) - \frac{s\langle \sigma v \rangle_{\psi\psi\to \chi\chi}}{Hx} \left(Y_\psi^2 - \frac{Y_\psi^{\rm eq\,2}}{Y_\chi^{\rm eq\,2}} Y_\chi^2 \right) - \frac{s\langle \sigma v \rangle_{\psi\psi\to Z'Z'}}{Hx} \left(Y_\psi^2 - \frac{Y_\psi^{\rm eq\,2}}{Y_{Z'}^{\rm eq\,2}} Y_{Z'}^2 \right)$$

Constraints

Structure formation

Free-streaming $$\lambda_{\rm FS} \approx \frac{1}{2}\int_{t_{\rm kd}}^{t_{\rm MRE}} dt \frac{v_\psi}{a(t)} \approx \frac{1}{2}\left(\frac{4\pi^3 g_{\rm eff}}{135}\right)^{-1/2} \sqrt{\frac{\xi}{T_{\rm kd} m_\psi}} \frac{M_{\rm pl}}{T_0} \log \frac{T_{\rm kd}}{T_{\rm MRE}}$$

Acoustic oscillations $$\lambda_{\rm AO} = \int_0^{t_{\rm kd}} \frac{dt}{a(t)} = \frac{1}{aH}\bigg|_{\rm kd} \approx \left(\frac{4\pi^3 g_{\rm eff}}{45}\right)^{-1/2} \frac{M_{\rm pl}}{T_{\rm kd} T_0}$$

$$\lambda_{\rm cutoff} = \max(\lambda_{\rm FS}, \lambda_{\rm AO}) < 0.1\,\mathrm{Mpc}$$